

The study of the longitude development of muons in the air shower

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Outline

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02

Study of muon production depth by simulation

- Introduction of data simulation
- Reconstruction method of Muon Production Depth(MPD)

03

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Relevant factors of X_{\max}^{μ} reconstruction

- Time resolution of muon detector
- Core and angular reconstruction resolution

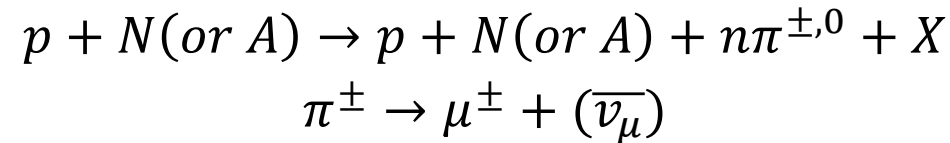
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Summary and Next plan

Physical Motivation

□ Muon production in air shower

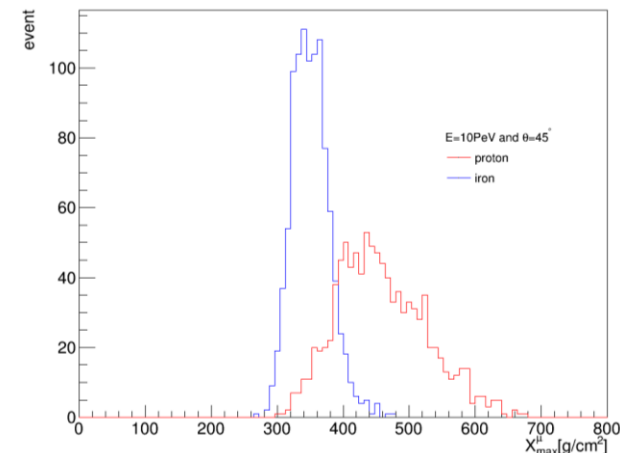
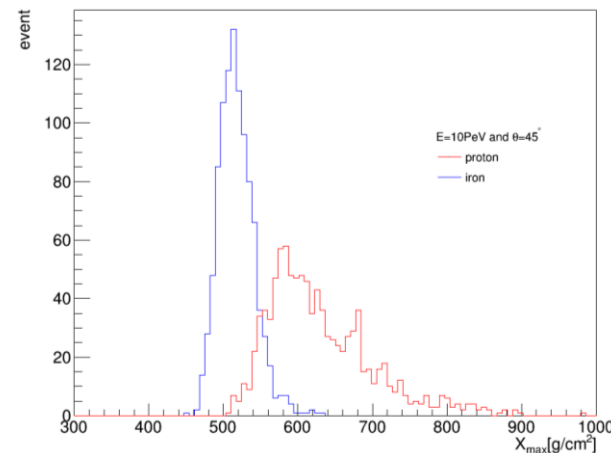
- ✓ The main contributions to the muon component in air showers are the decay of charged pions (π^+ , π^-) and kaons (K^+ , K^- , K_L^0), and other charmed particles, such as D^\pm , D^0 , J/ψ



- ✓ The production depth of muon can trace the positions of charged pions, so the longitude development of muons preserves the information of primary particles and plays an important role in the study of composition identification.

□ Purpose to study muon

- ✓ Measure the cosmic ray energy spectrum and mass composition
- ✓ Help to study the hadronic interaction models
- ✓ Provide insight on whether new physics phenomena take place

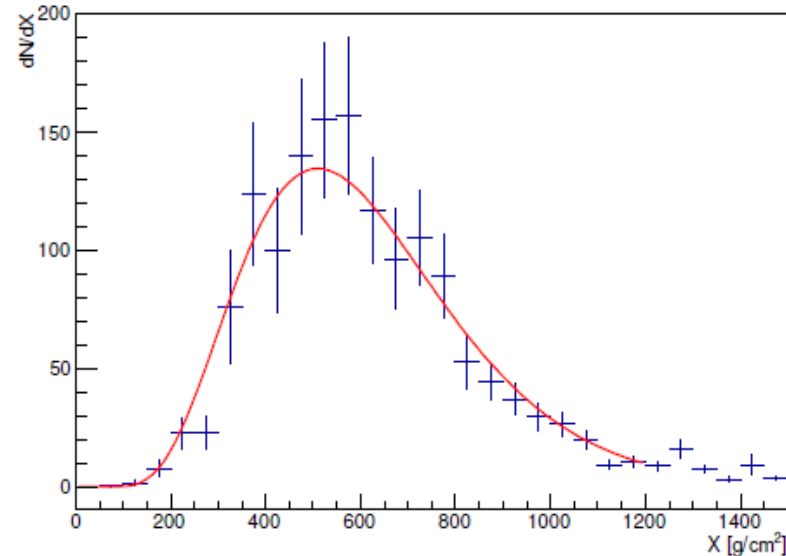
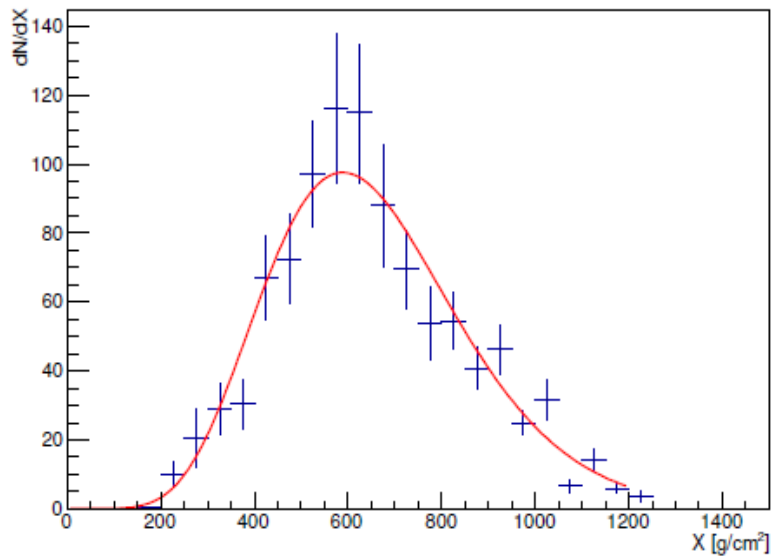


- X_{\max} : the air depth of the shower maximum development
- X_{\max}^μ : the depth along the shower axis where the production of muons reaches maximum

X_{\max}^μ can be used to identify mass compositions of cosmic rays.

Physical Motivation

- Method 1: By tracking the trajectories of the muon
 - KASCADE-Grande (reaches up to 10^{18} eV)
- Method 2: By the arrival time of muons
 - Surface detector of Pierre Auger Observatory (reaches up to 10^{20} eV)



We try to reconstruct the muon production depth by the arrival time of muon in 10^{16} eV

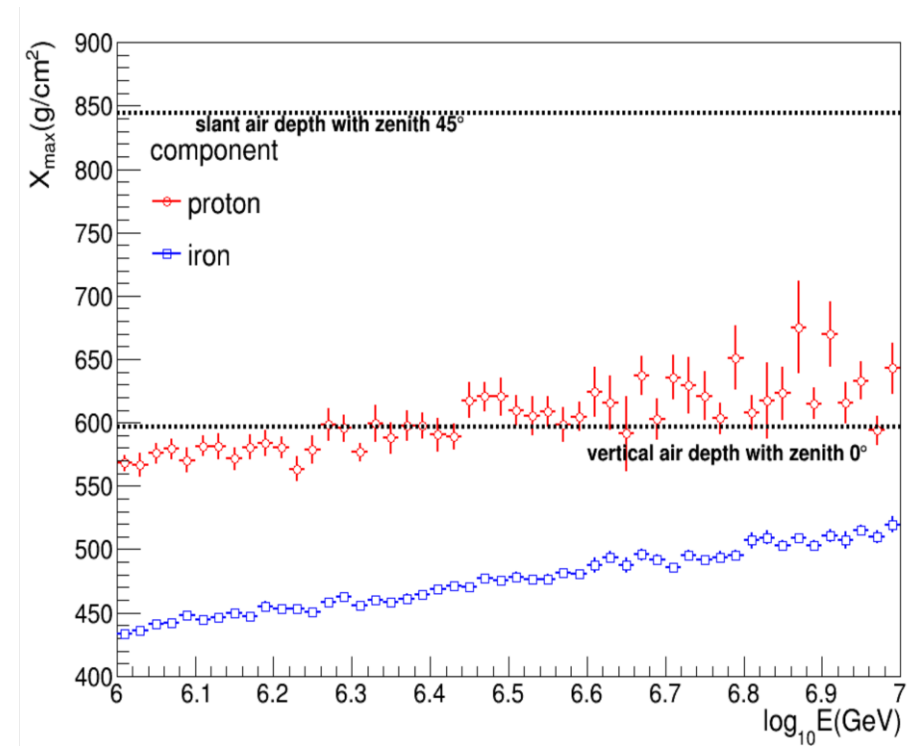
Fig. 8: Left) The reconstructed MPD distribution for a proton shower with $\theta = 48^\circ$ and $E = 10^{19.8}$ eV simulated using QGSJetII-04 model. Right) The reconstructed MPD distribution for a proton shower with $\theta = 58^\circ$ and $E = 10^{19.6}$ eV simulated using QGSJetII-04 model. The USP function fits are also shown.

Introduction of data simulation

➤ Simulation with CORSIKA (v76400)

- EPOS-LHC+GHEISHA
- Altitude: 4410m (vertical air depth= $600\text{g}/\text{cm}^2$)
- Energy: 10PeV; 5PeV; 1PeV
- Zenith angle(θ): fixed 45°
- MUPROD: additional information of muon

Number of shower	10PeV	5PeV	1PeV
Proton	1000	2000	2000
Iron	1000	2000	2000

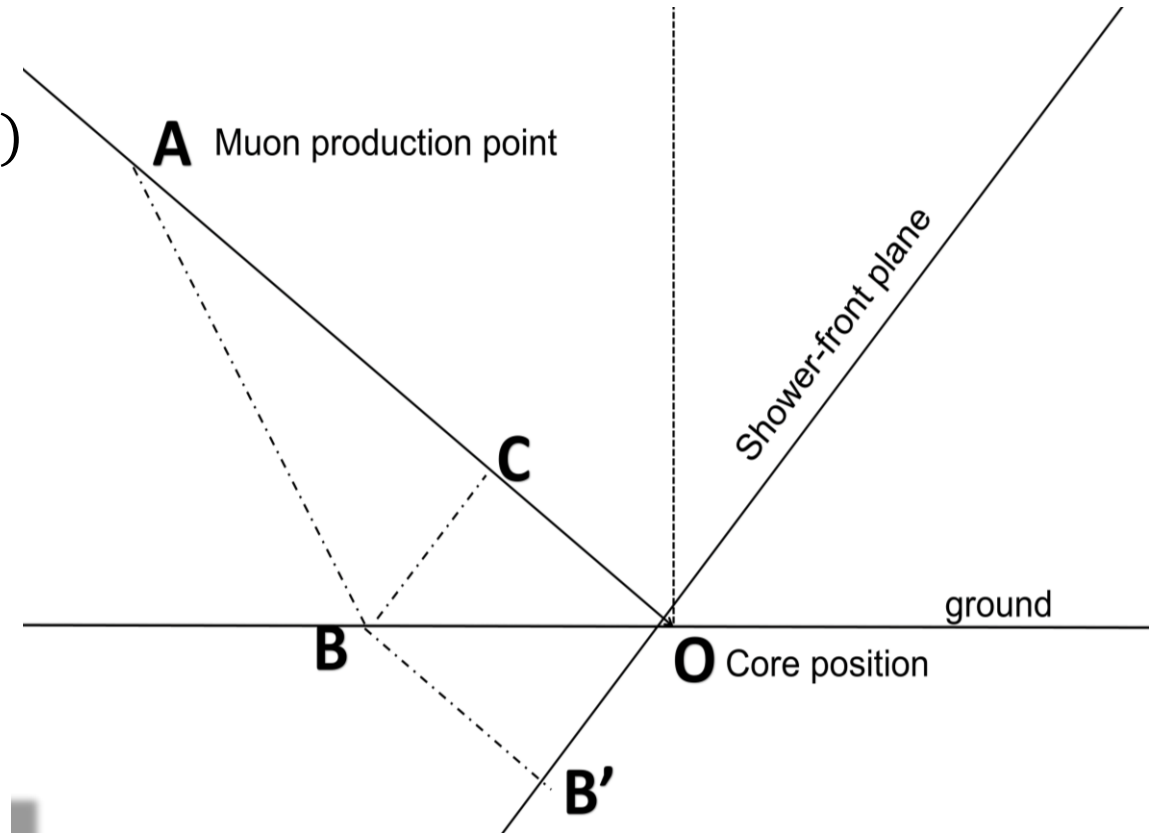


*** In order to ensure the shower full development, the zenith angle is selected as 45°**

Reconstruction method of Muon Production Depth(MPD)

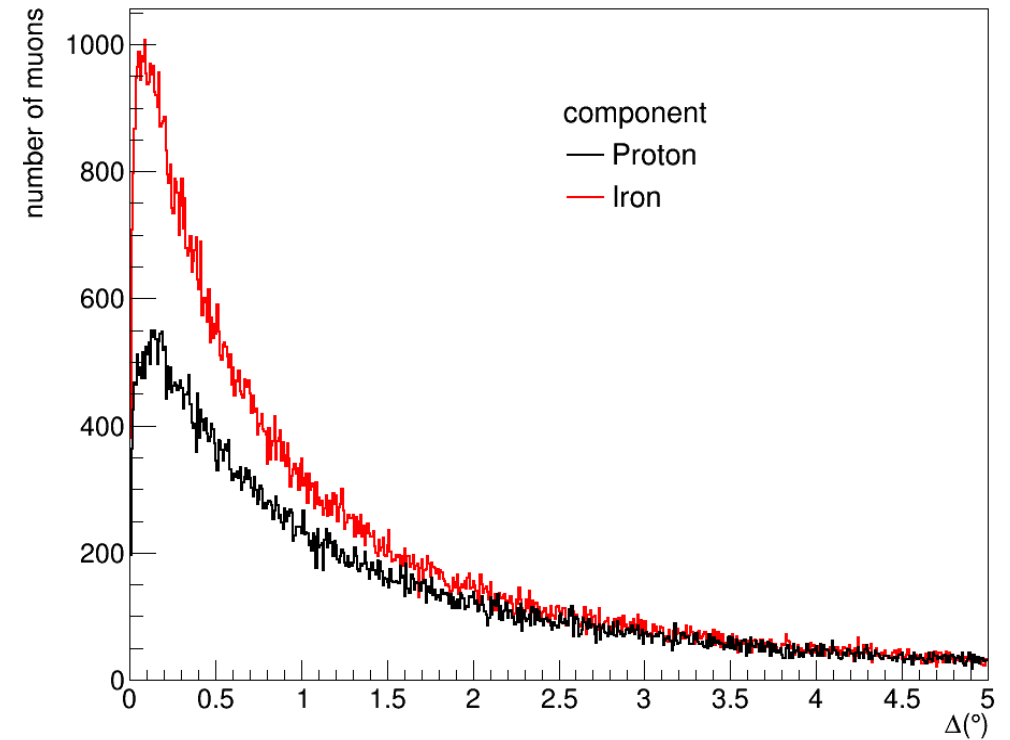
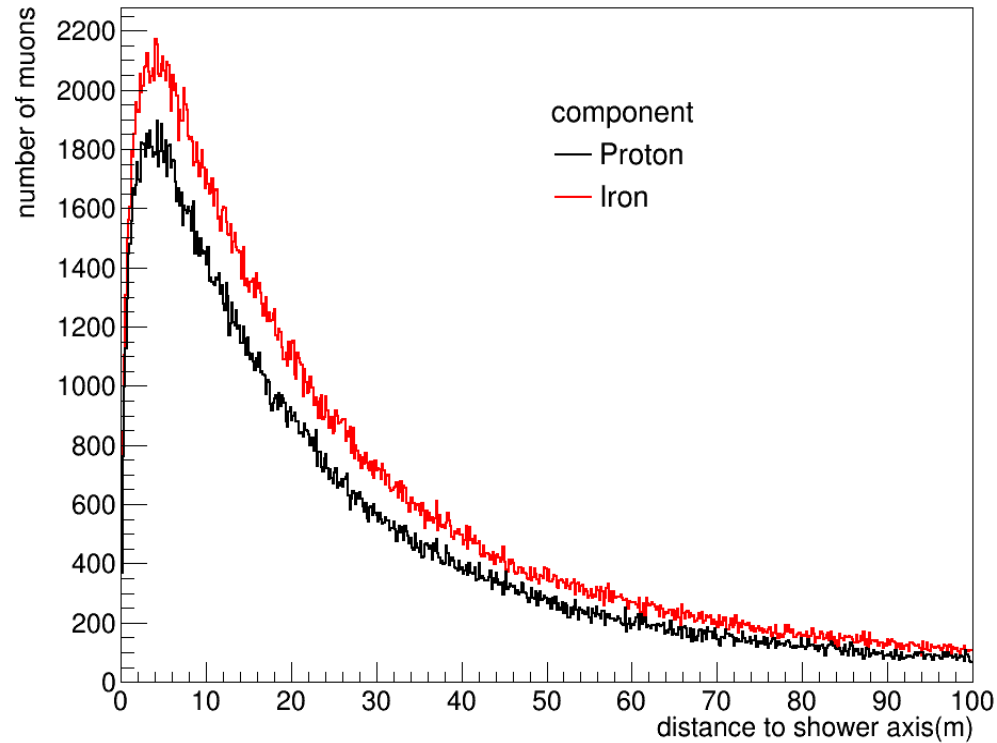
- Considering **geometric delay** as : $t_g = (t_{AB} - t_{AC})$
- $\overline{AB} = \sqrt{\overline{BC}^2 + (\overline{AO} - \overline{BB'})^2}$
- $\overline{AB} = \overline{AC} + c \times t_g$
- Correction by the path traveled by the parent mesons : $t_\pi \sim 3ns$

$$\overline{AO} = \frac{1}{2} \left(\frac{\overline{BC}^2}{c \times (t_g)} - c \times (t_g) \right) - c \times t_\pi$$



- The muons are produced on the shower axis
- Muons travel in straight lines at the speed of light c
- Geometric delay is dominant

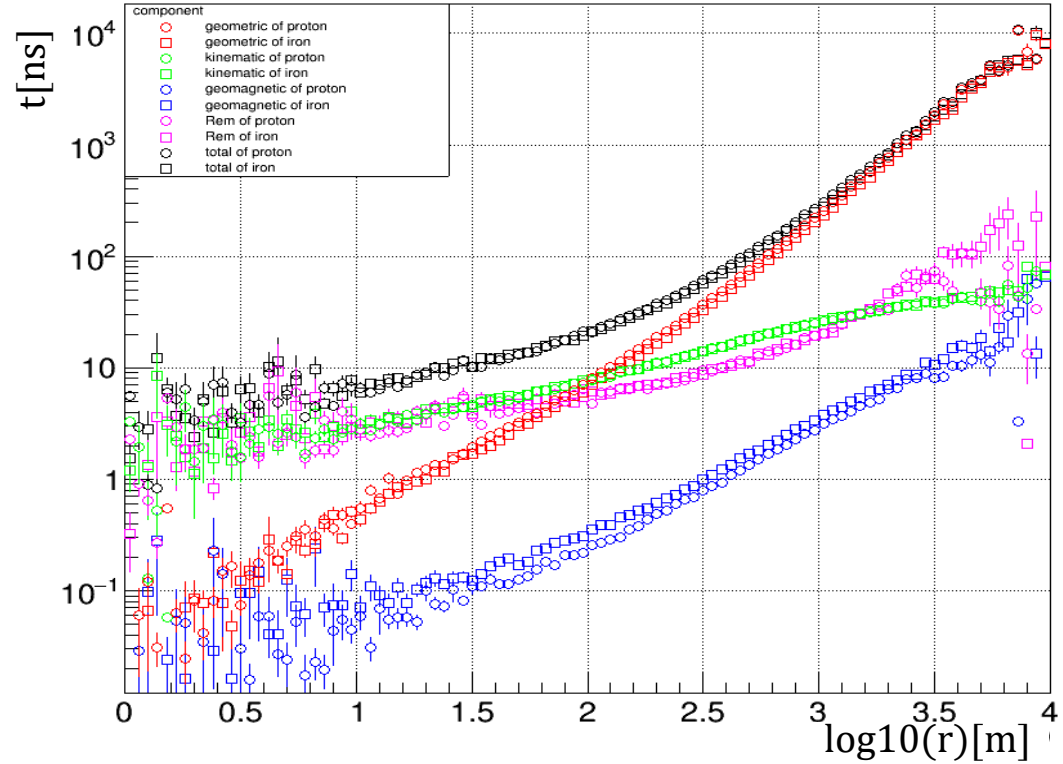
Reconstruction method of Muon Production Depth(MPD)



* most muons have small distance to shower axis and also lie within small Δ° (The angular difference between the production and incidence at the ground)

✓ **Approximation: The muons are produced on the shower axis and travel in straight lines**

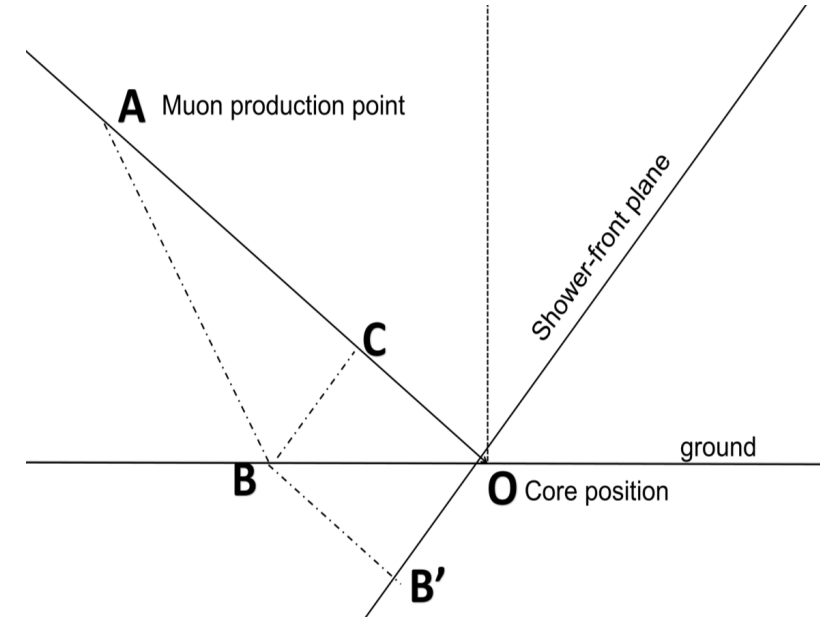
Reconstruction method of Muon Production Depth(MPD)



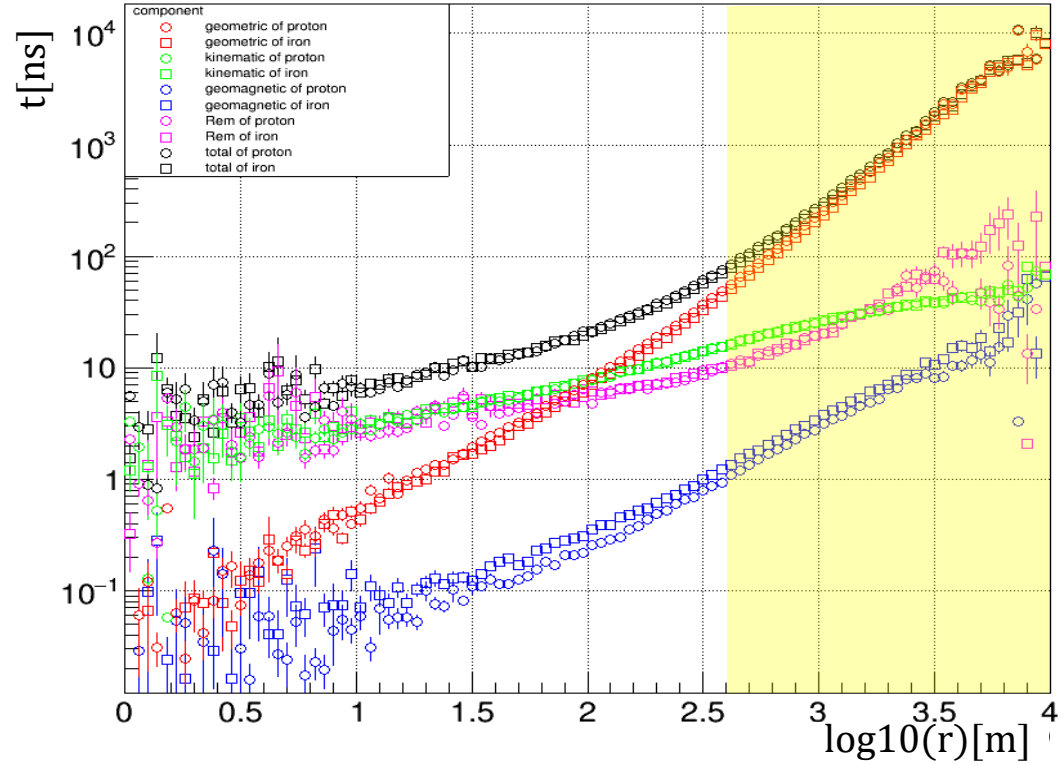
- t_{total} : total(measurement) time delay
- t_g : **geometric delay**
- t_ϵ : **kinematic delay**
- t_B : **geomagnetic delay**
- t_{Rem} : **multiple scattering**

r : distance from muon position to core on SFP

- ✓ **The geometric delay is dominant ($v_\mu \approx c$)**
- * $t_B + t_{Rem} < t_{total} \times 10\%$ when $r > 400m (10^{2.6})$
- * t_g is dominant when $r > 400m$



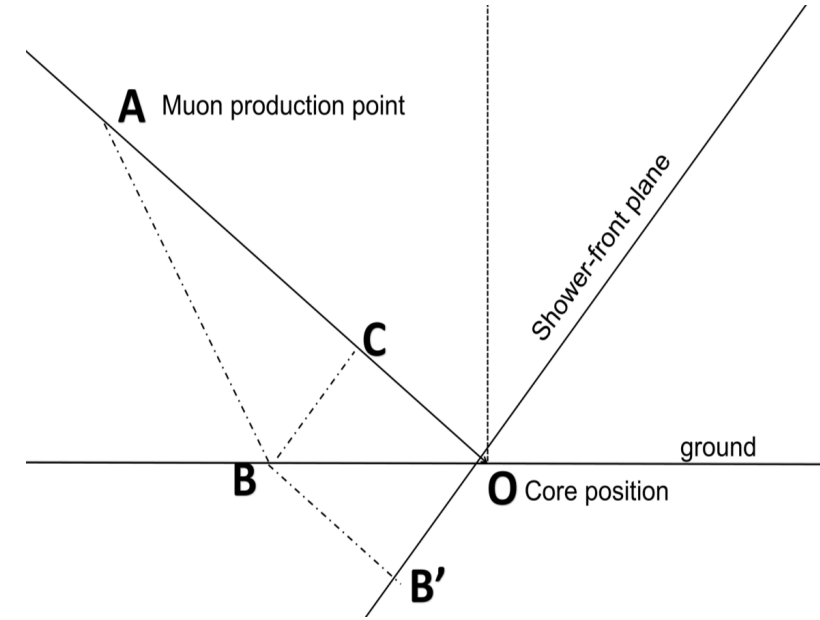
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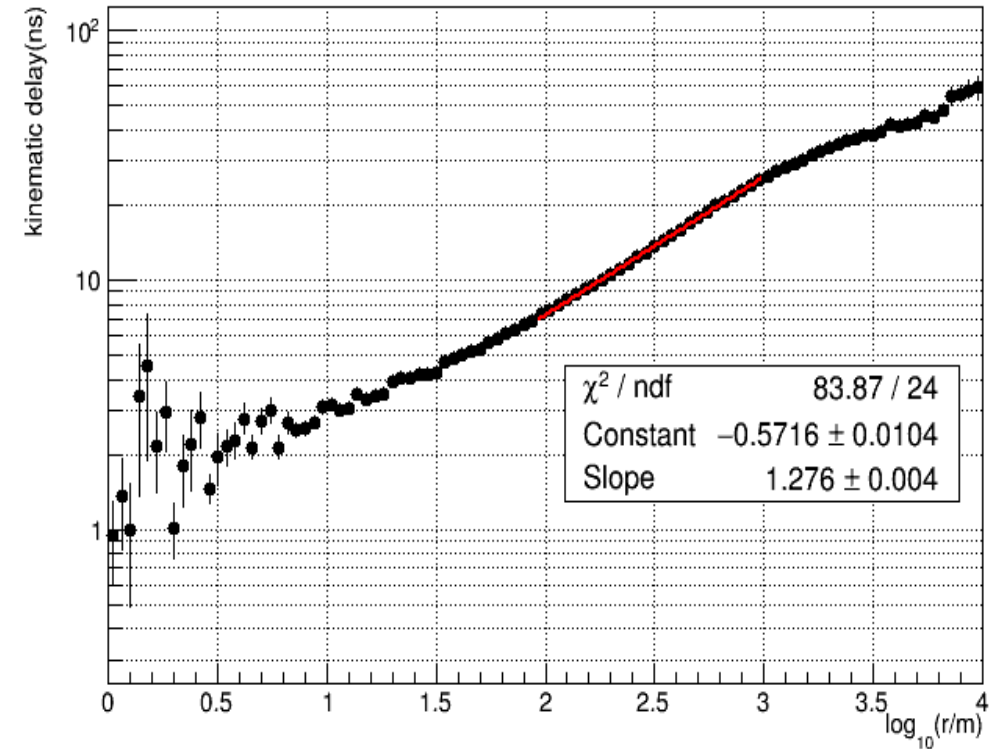
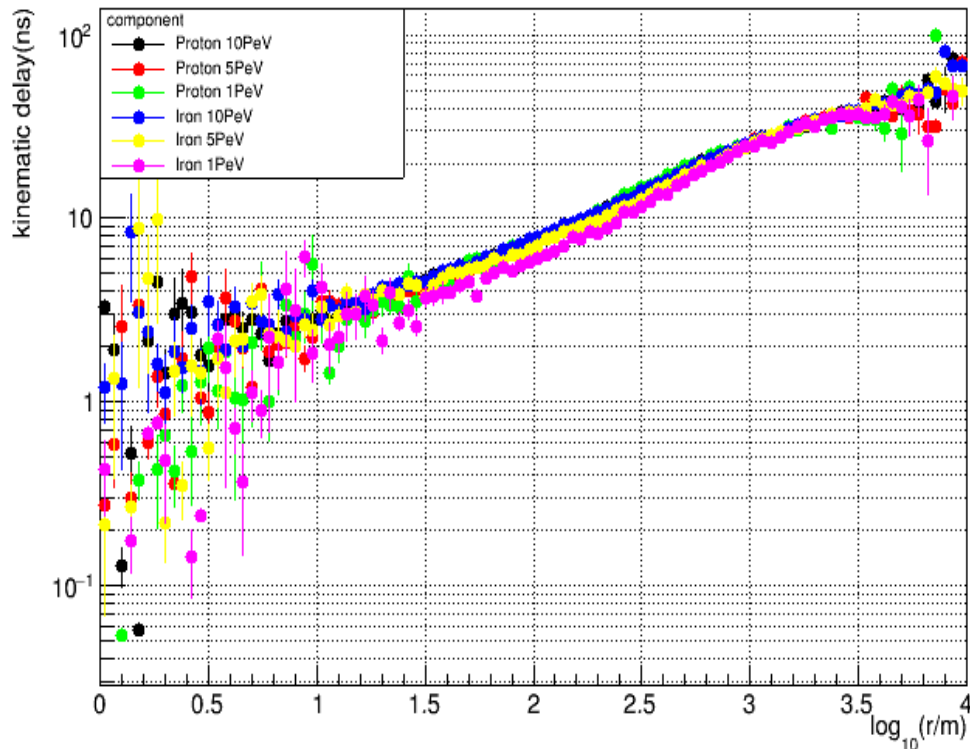
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Reconstruction method of Muon Production Depth(MPD)

◆ Correction for t_g



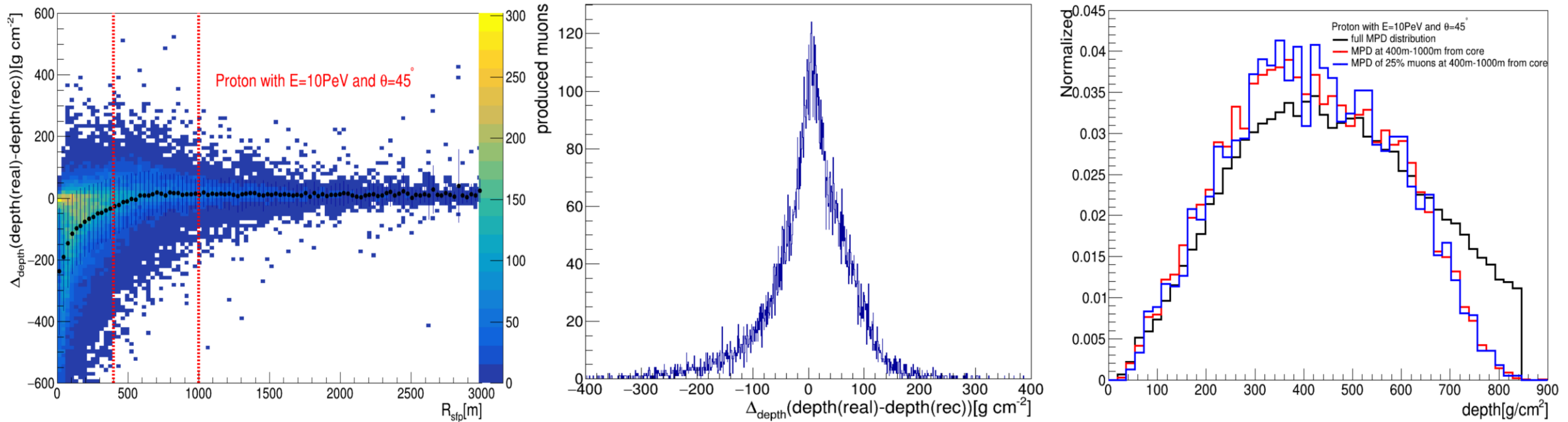
- The relationship between the t_ϵ and r of different primaries with different energy is similar
- The relationship between t_ϵ and r of different compositions and energies can be fitted with an exponential function:

$$t_\epsilon = \exp(-0.57 + 1.276 \times \log_{10}(r))$$

$$\checkmark t_g = t_{\text{total}} - t_\epsilon$$

Features of the muon profiles

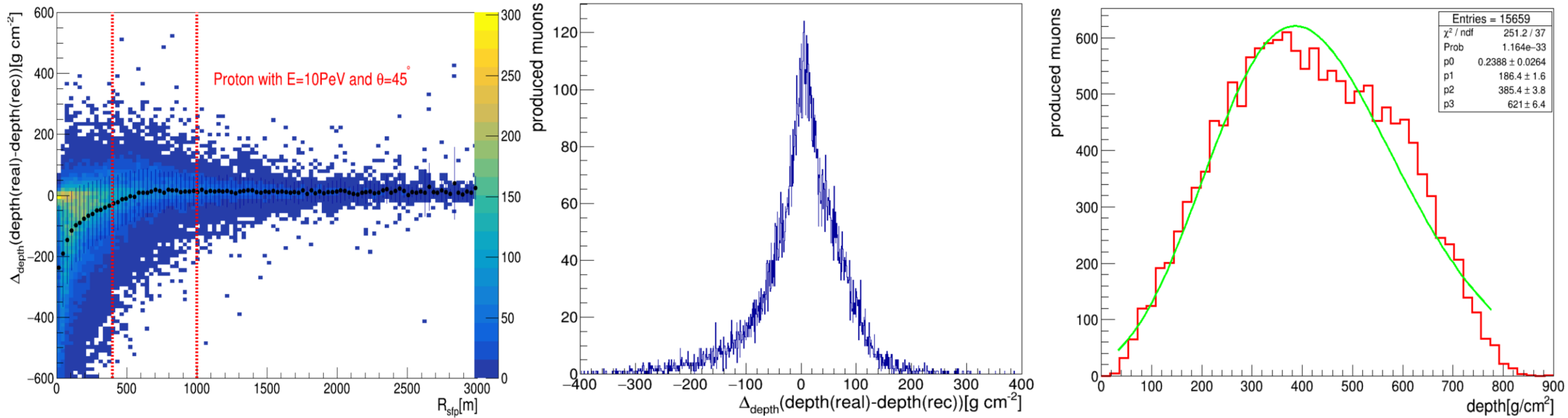
◆ Fit the longitudinal profile of the muon



- The differences between the MPD distributions at different distance to the core and different coverage are small

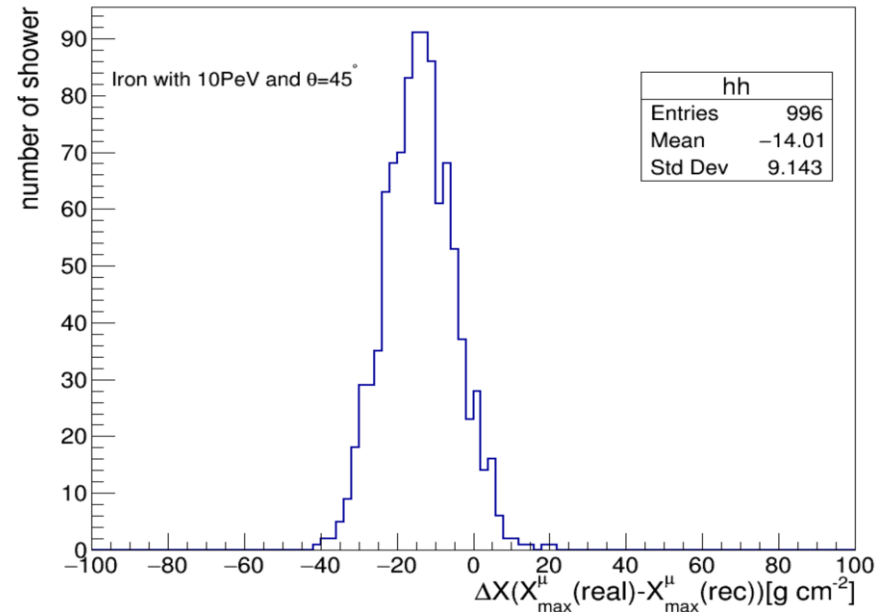
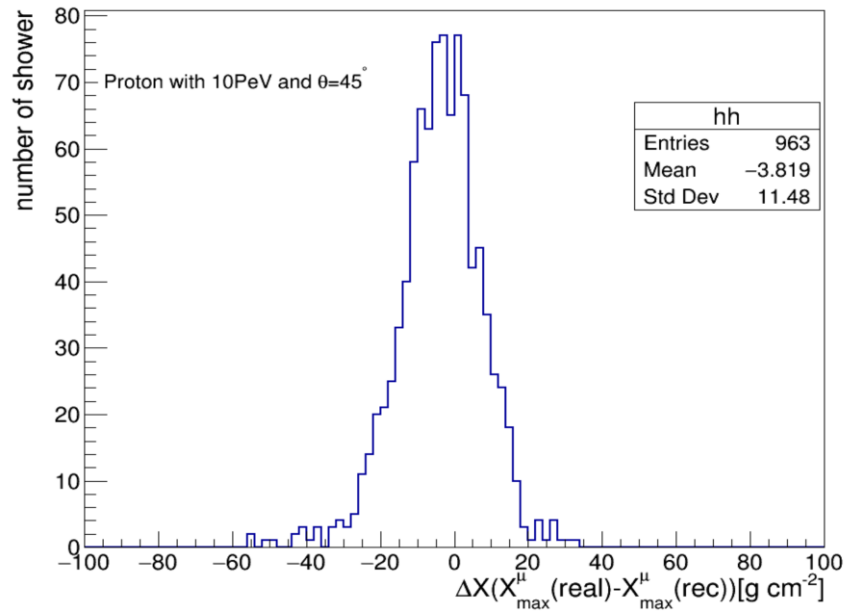
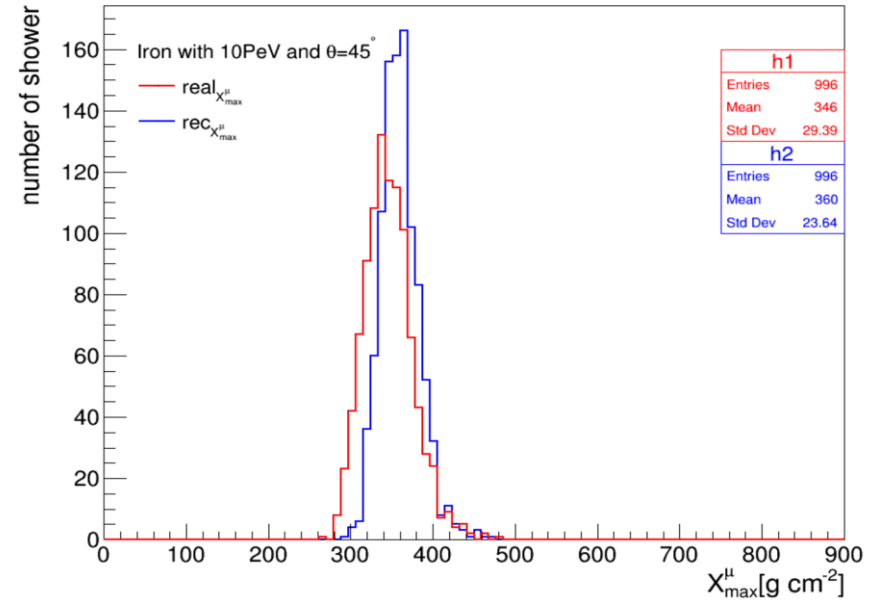
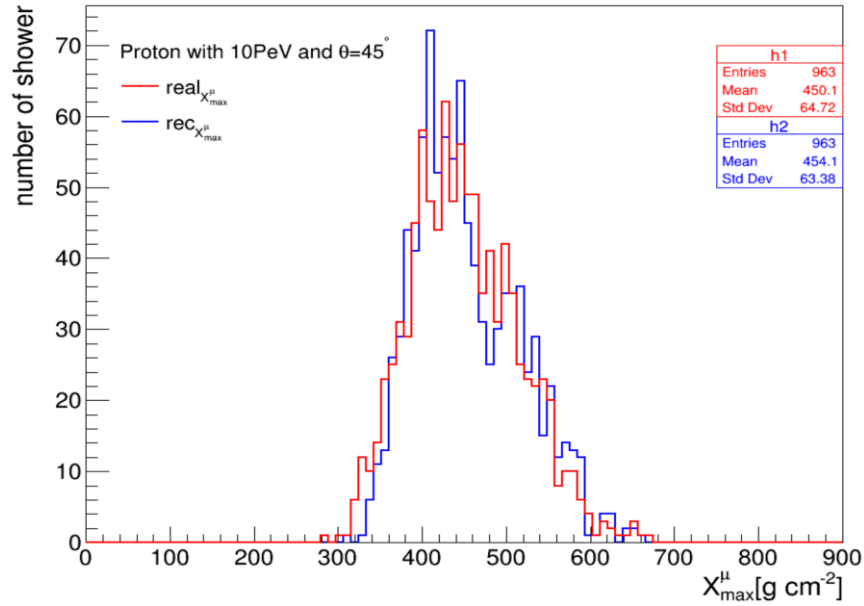
Features of the muon profiles

◆ Fit the longitudinal profile of the muon

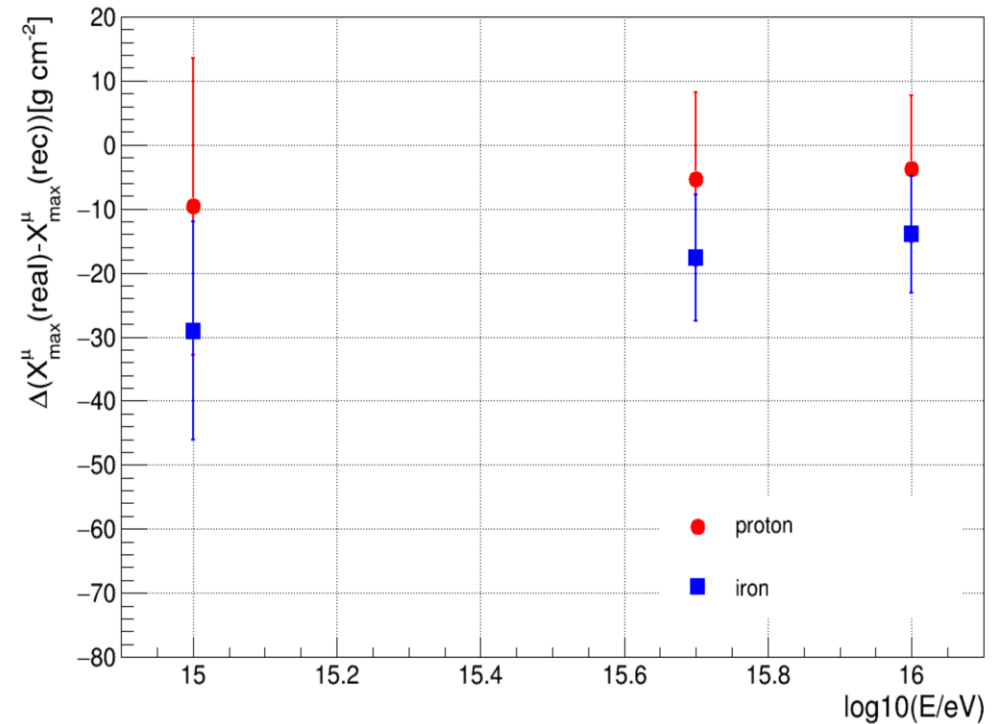
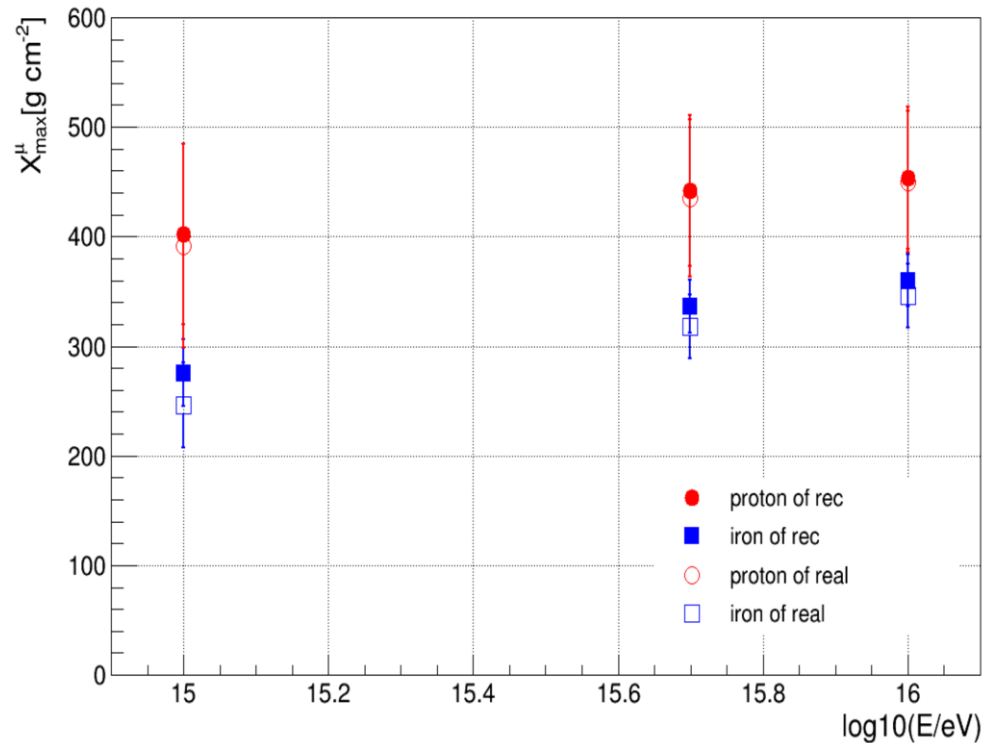


$$\text{Fit formula: } \frac{dN}{dX} = N_{max} \left(1 + \frac{R}{L} (X - X_{max}^\mu)\right)^{R-2} \exp\left(-\frac{X - X_{max}^\mu}{LR}\right)$$

Features of the muon profiles

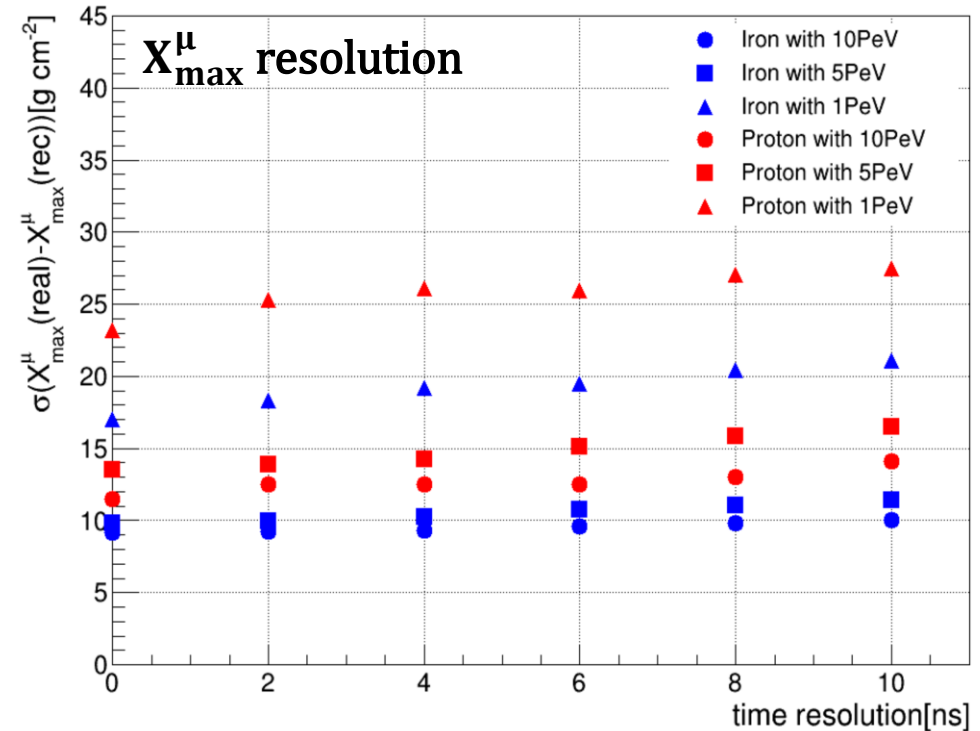
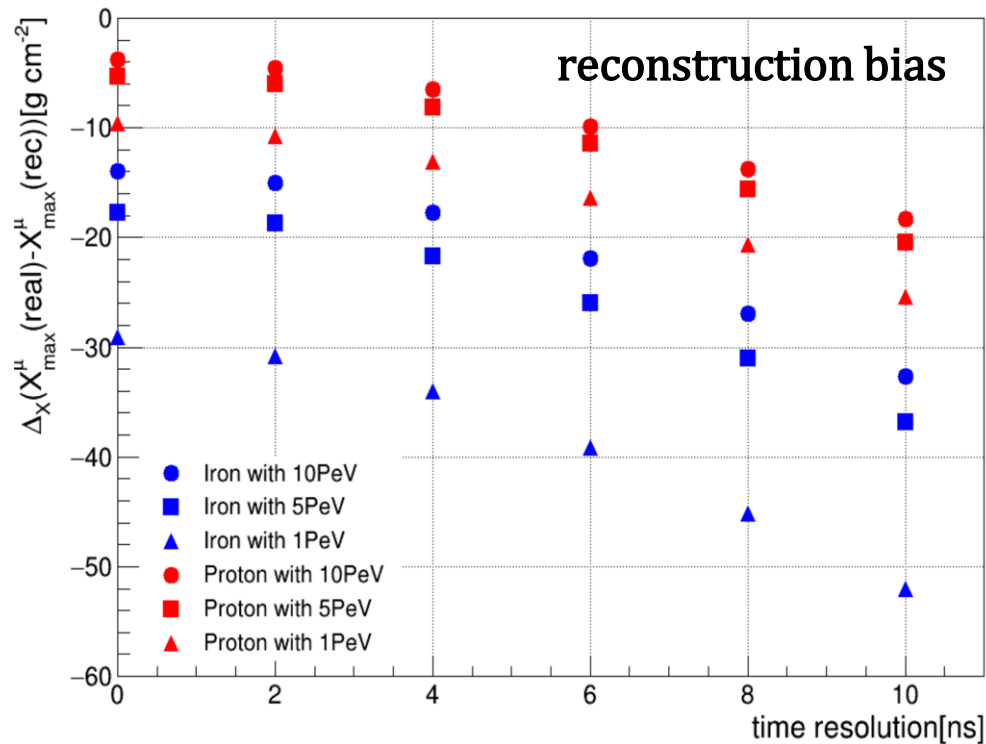


The X_{\max}^{μ} study with different energies



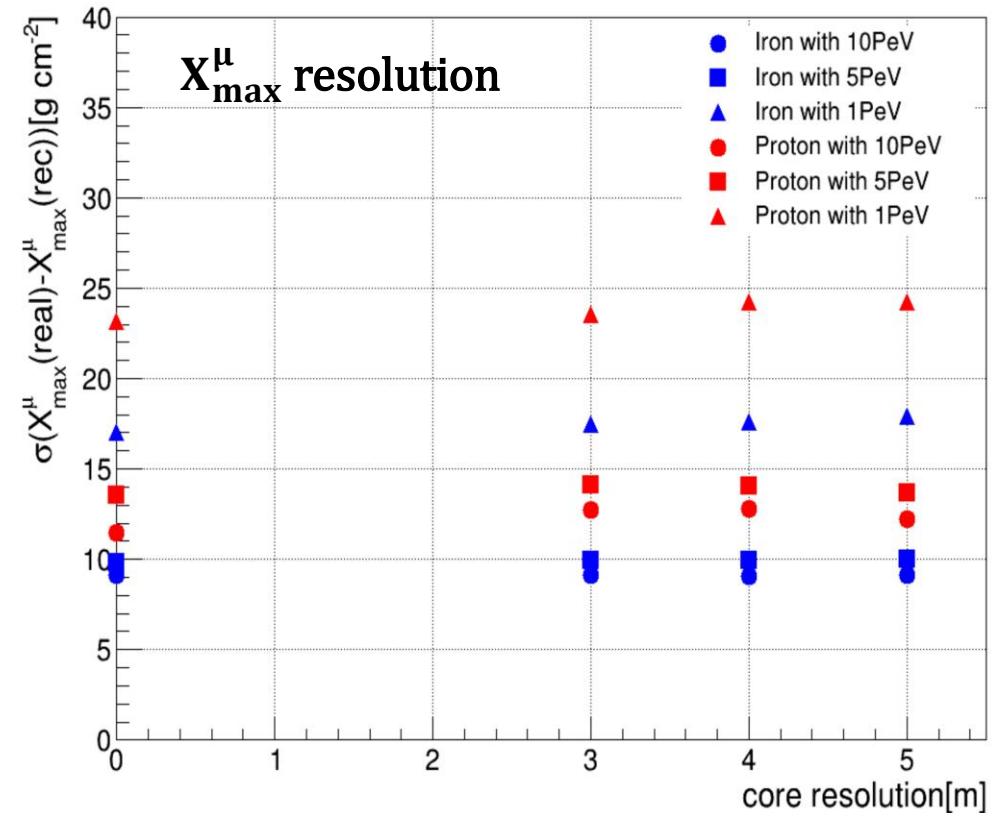
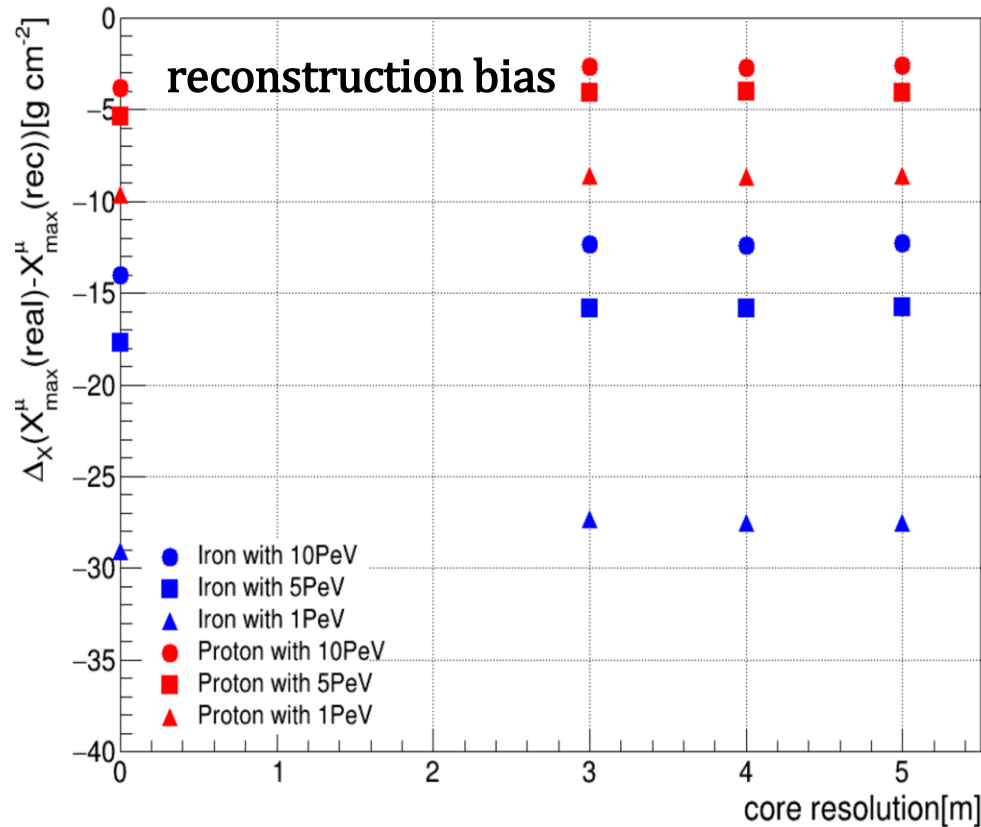
- The reconstruction bias of X_{\max}^{μ} for Proton is smaller than that for Iron
- The X_{\max}^{μ} resolution improves with energy increasing
- X_{\max}^{μ} can be used to identify mass compositions of cosmic rays.

The X_{\max}^{μ} study with different time resolution



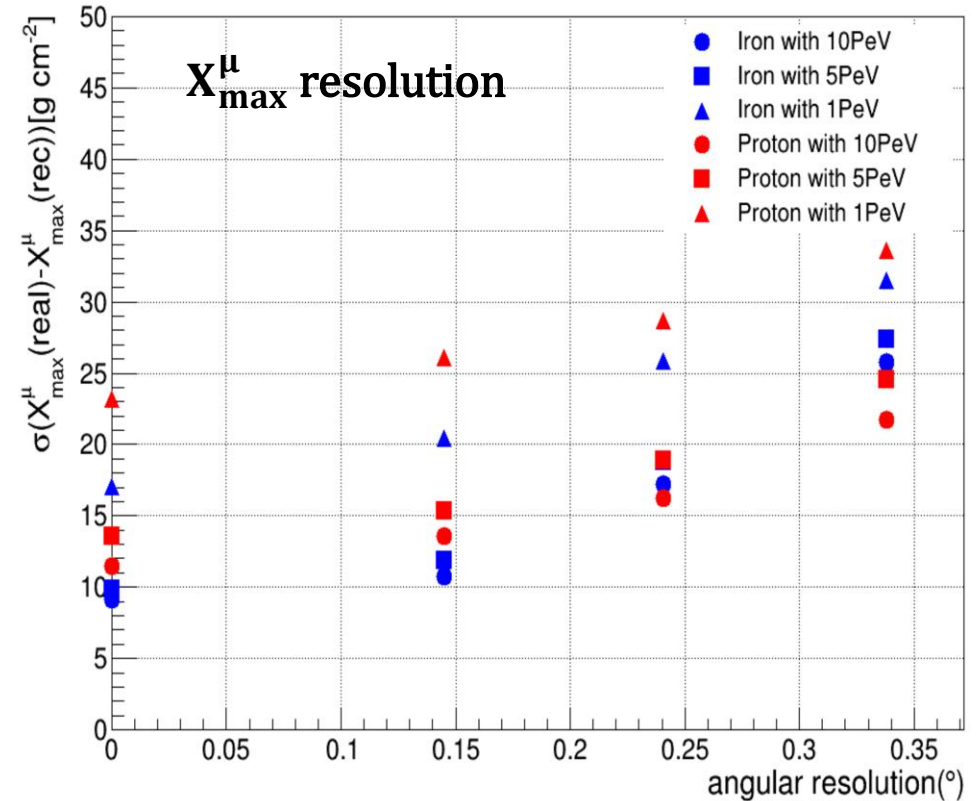
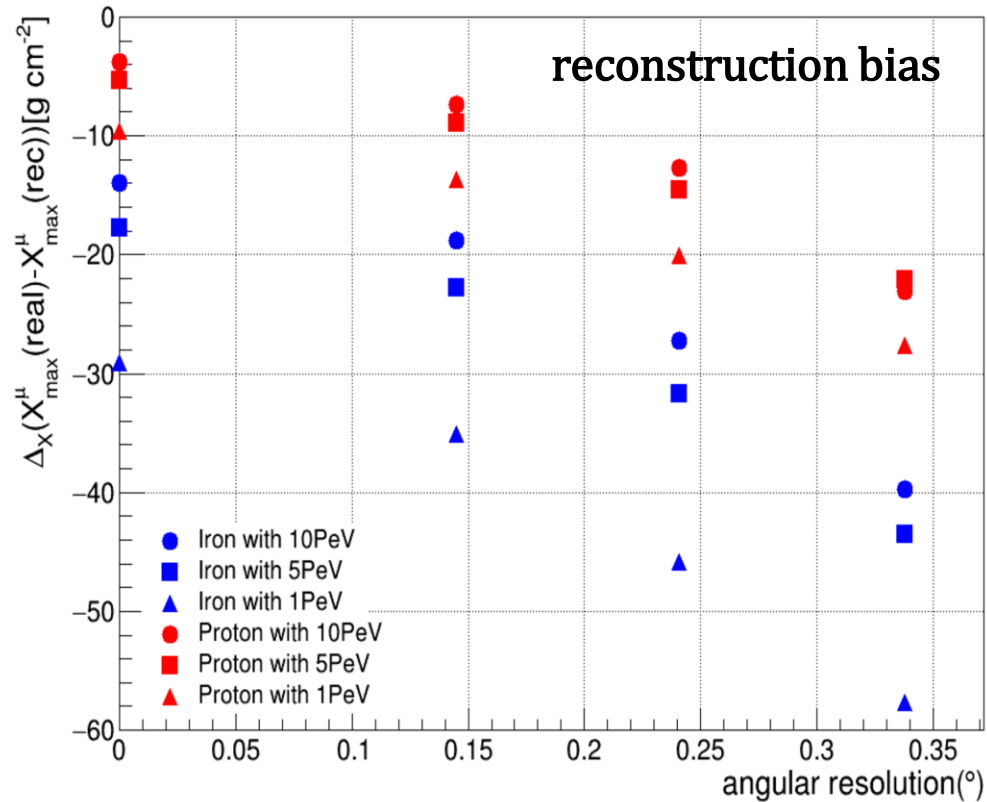
- The time resolution can affect reconstruction bias of X_{\max}^{μ}
- The time resolution has little influence on X_{\max}^{μ} resolution

The X_{\max}^{μ} study with different core resolution



- The core resolution has little influence on the reconstruction of MPD

The X_{\max}^{μ} study with different angular resolution



- The angular resolution has a great impact on the reconstruction of MPD

Summary and Next plan

Summary

- The longitude development in the air shower and X_{\max}^{μ} can be reconstructed according to the geometry effect
- The time resolution and core resolution have little influence on the X_{\max}^{μ} resolution
- The time resolution and angular resolution also affect on the reconstruction bias of X_{\max}^{μ}

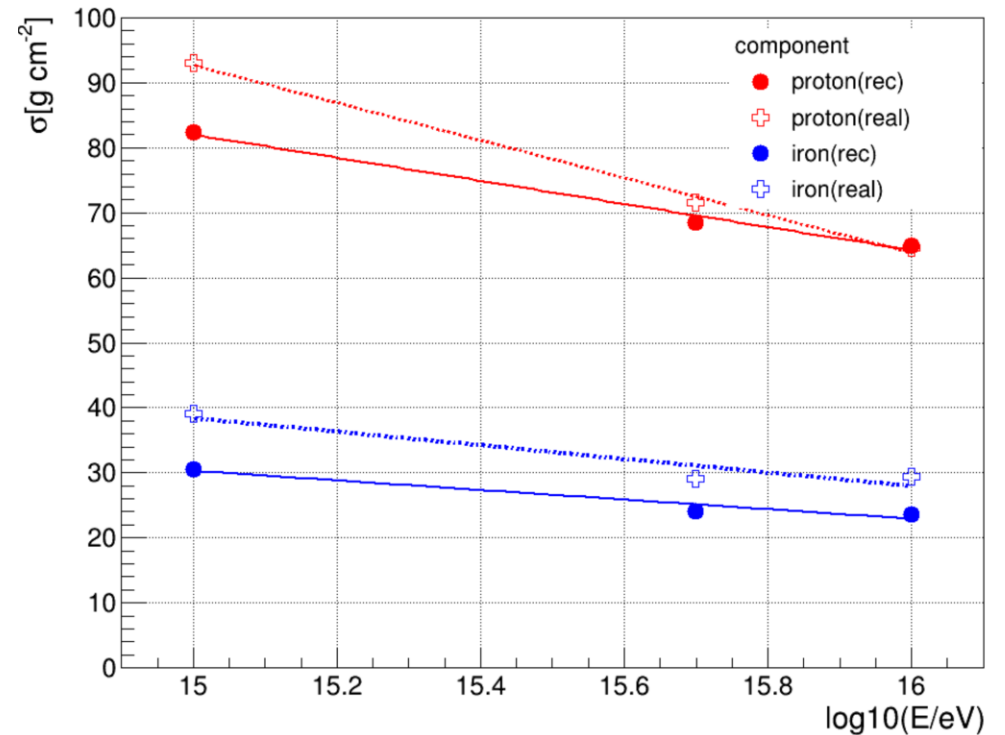
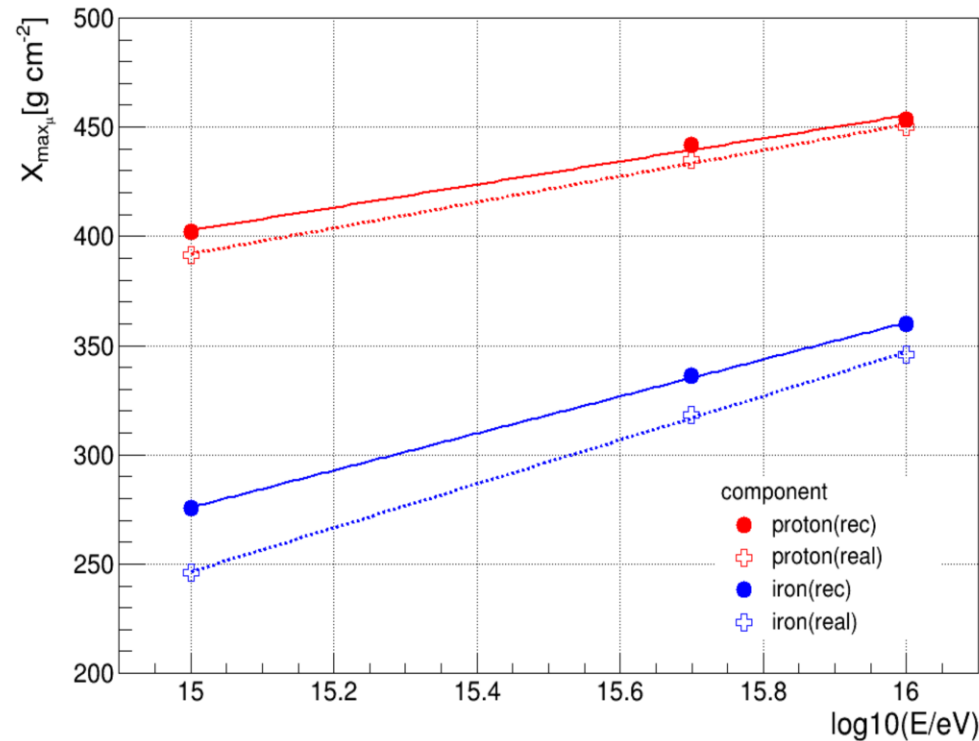
Next plan

- The well performance muon detector in LHAASO will be used to measure the MPD
- Identify mass compositions of cosmic rays by X_{\max}^{μ}
- Measure $\langle \ln A \rangle$ by the X_{\max}^{μ} and it's fluctuations

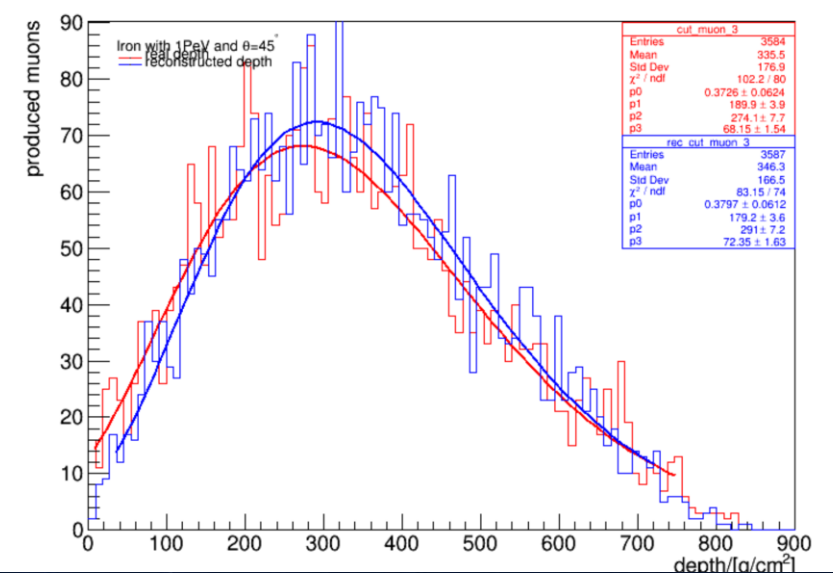
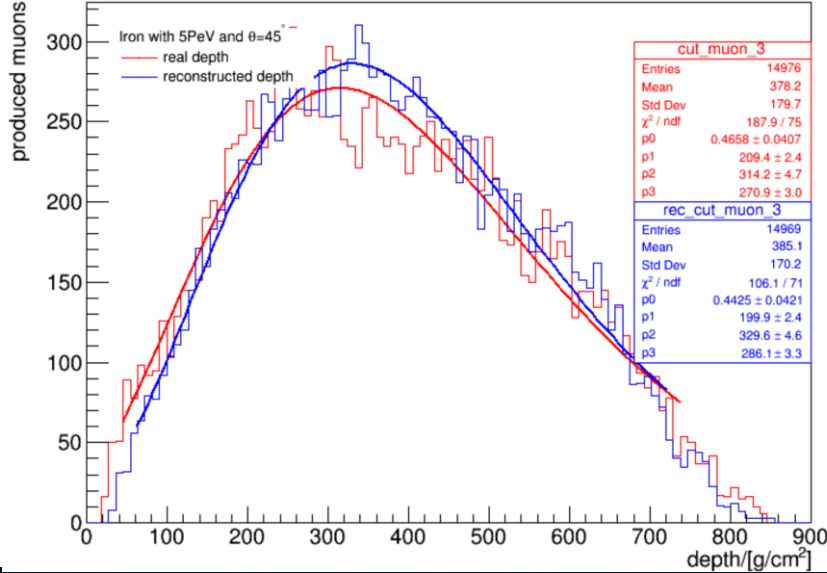
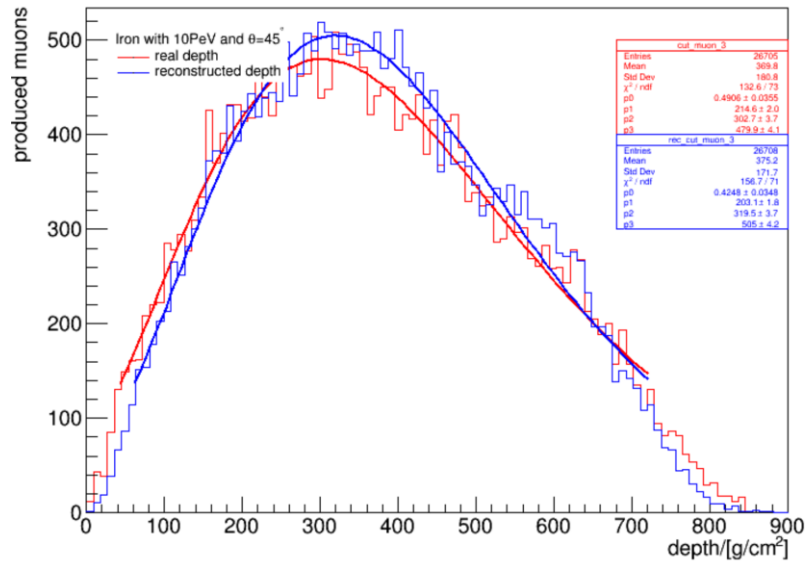
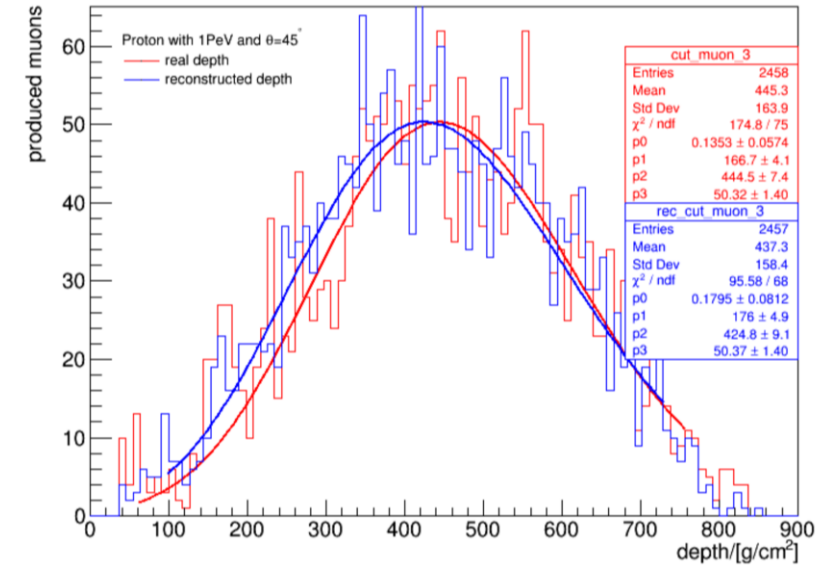
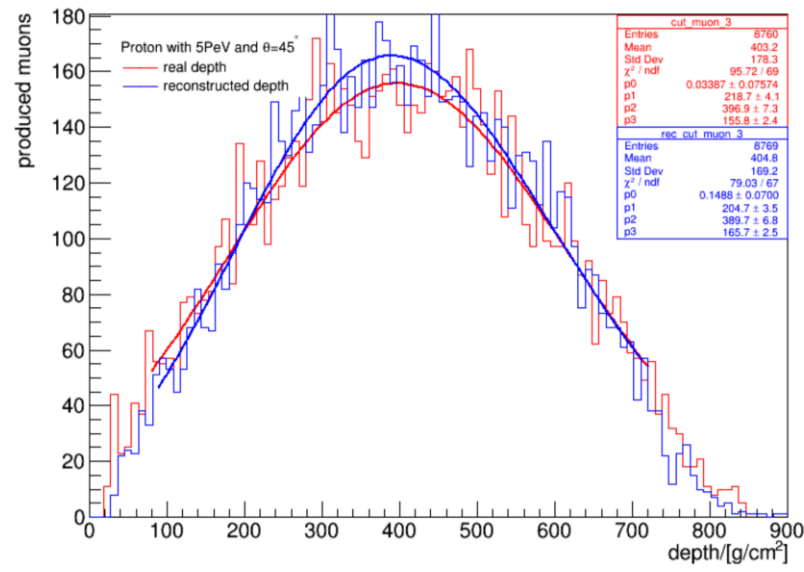
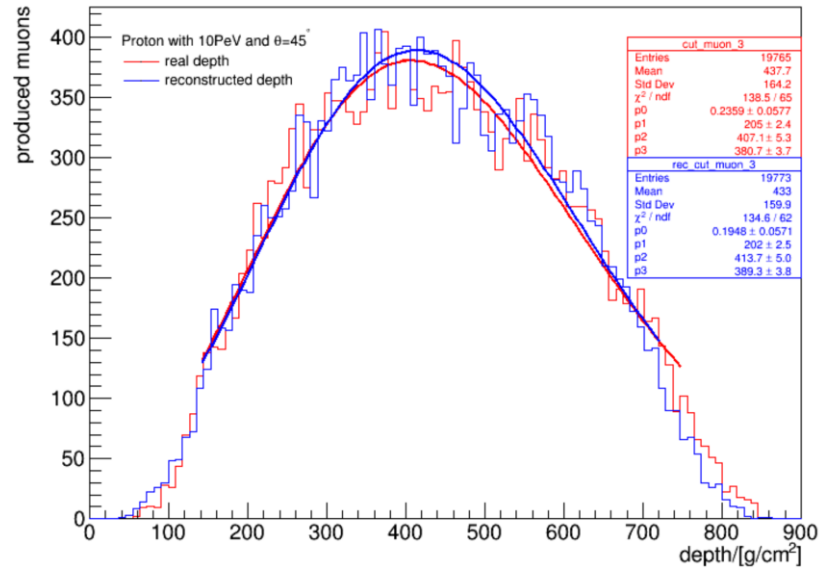
THANK YOU

BACK UP

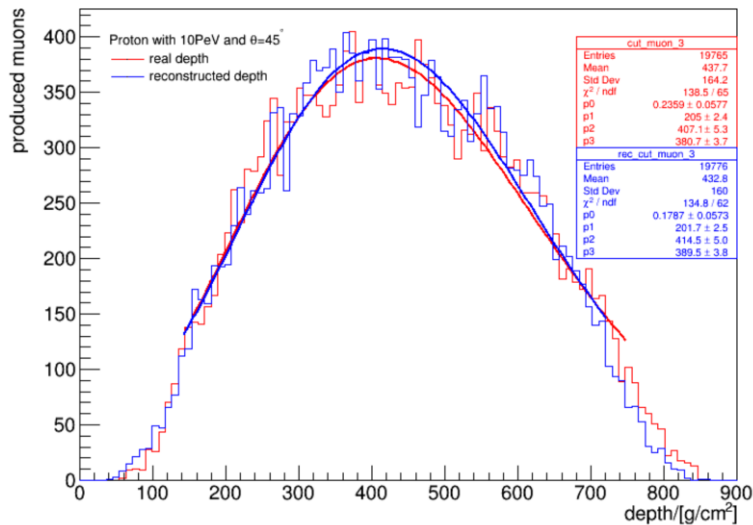
◆ Measure the mass of cosmic rays by the reconstructed muon profile



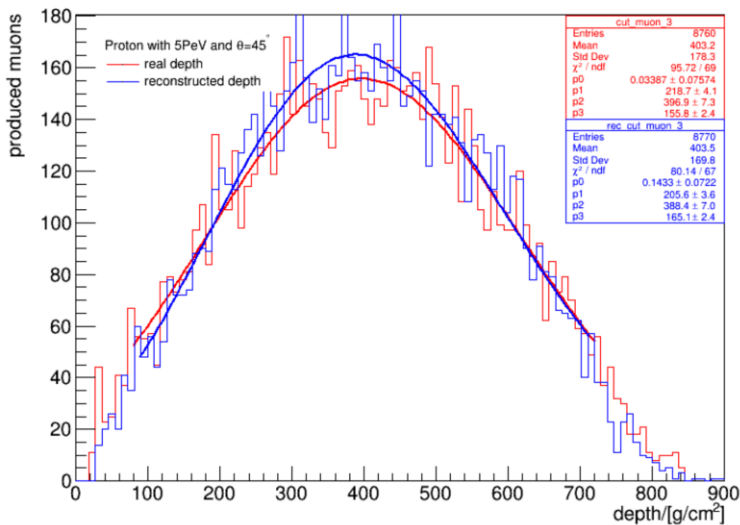
- * Measure $\langle \ln A \rangle$ by learning the longitudinal development of muons
- * Identify mass compositions of cosmic rays.



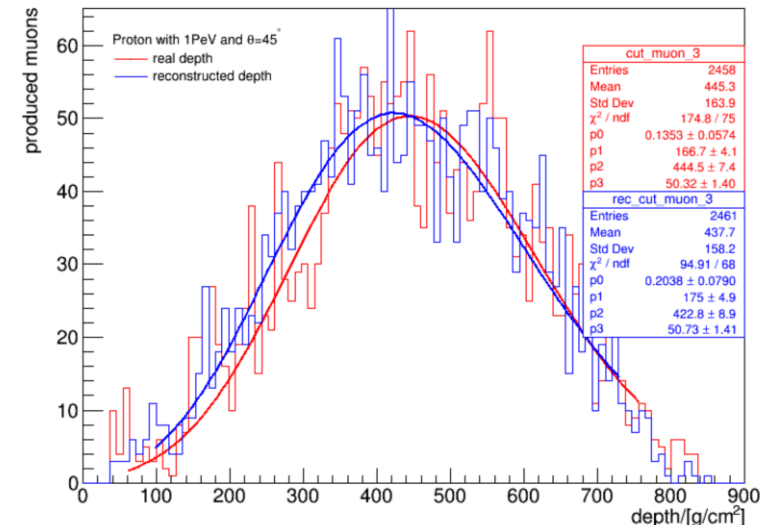
survial muon(>1GeV && 400m<rsfp<1km)



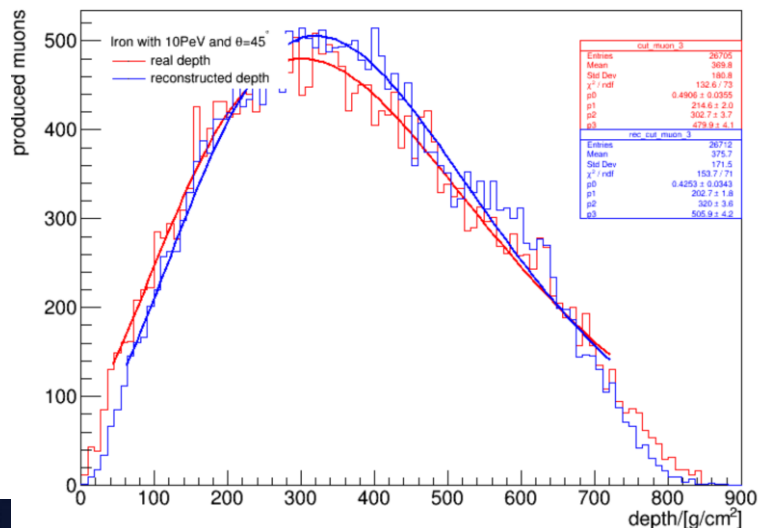
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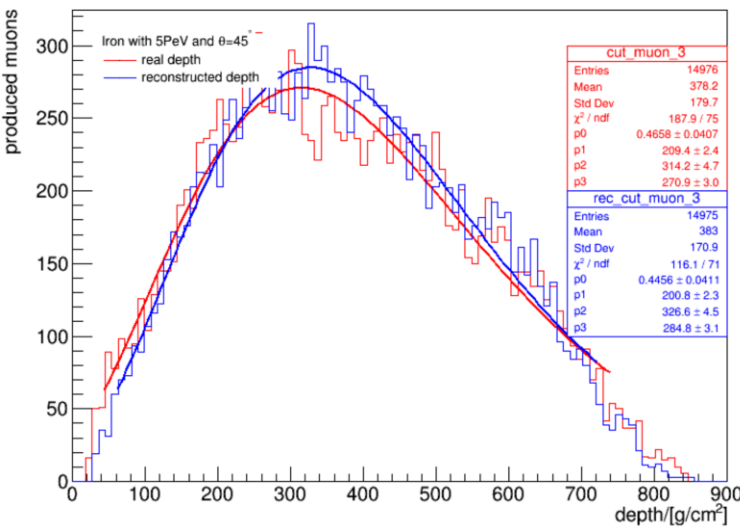
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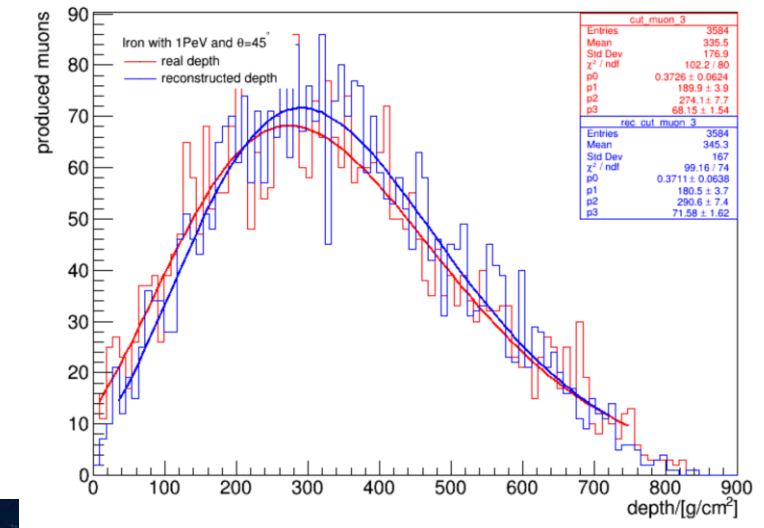
survial muon(>1GeV && 400m<rsfp<1km)



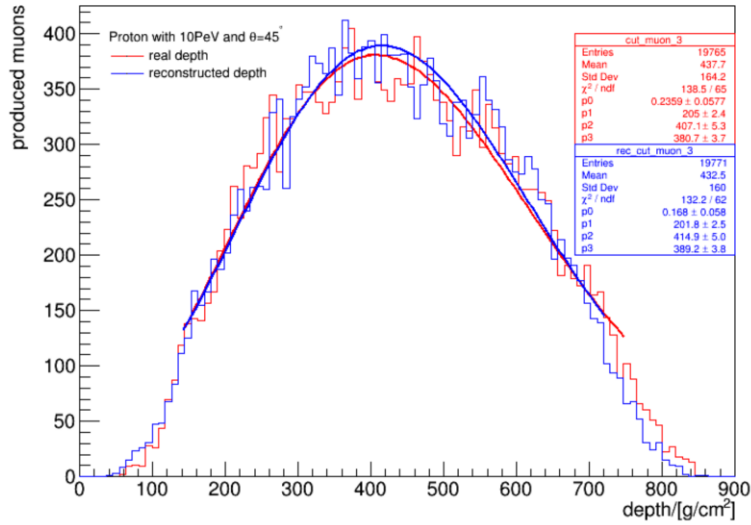
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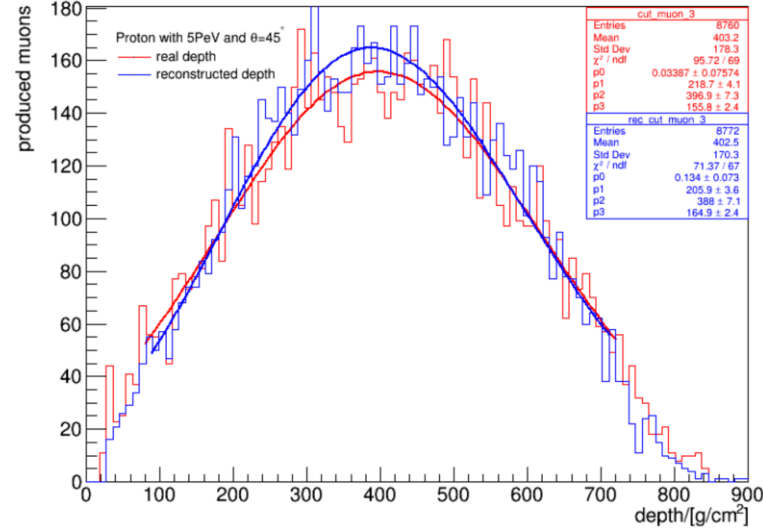
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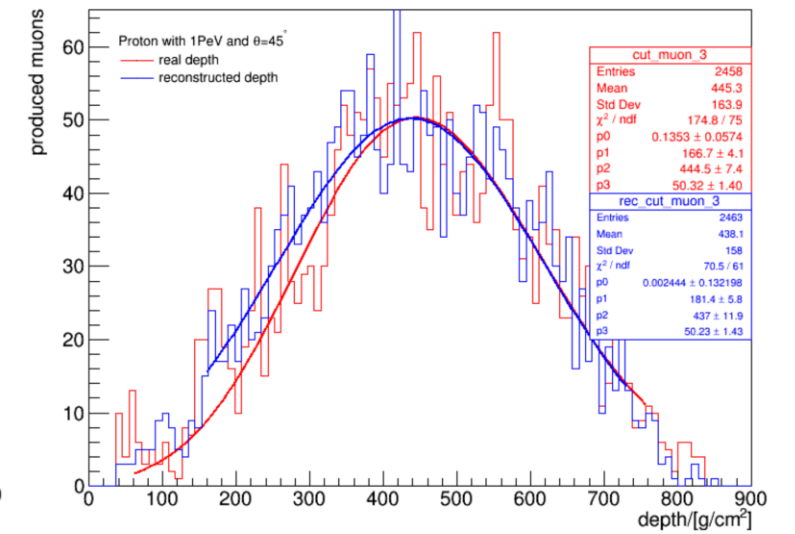
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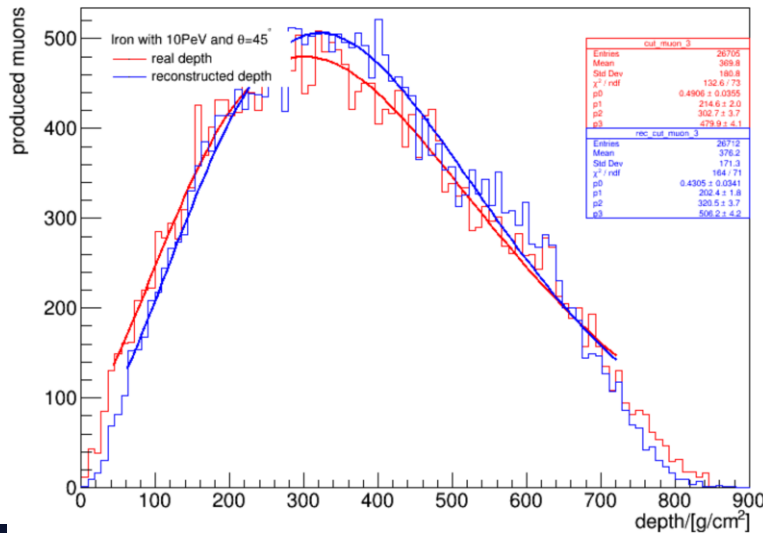
survial muon(>1GeV && 400m<rsfp<1km)



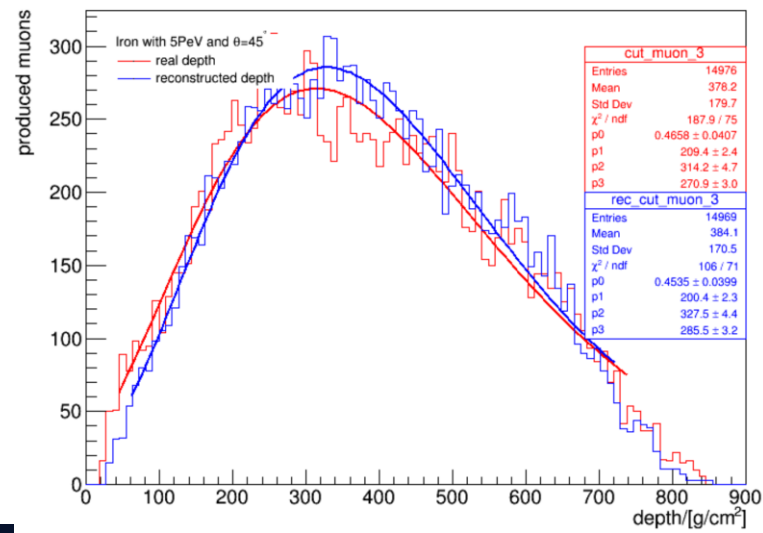
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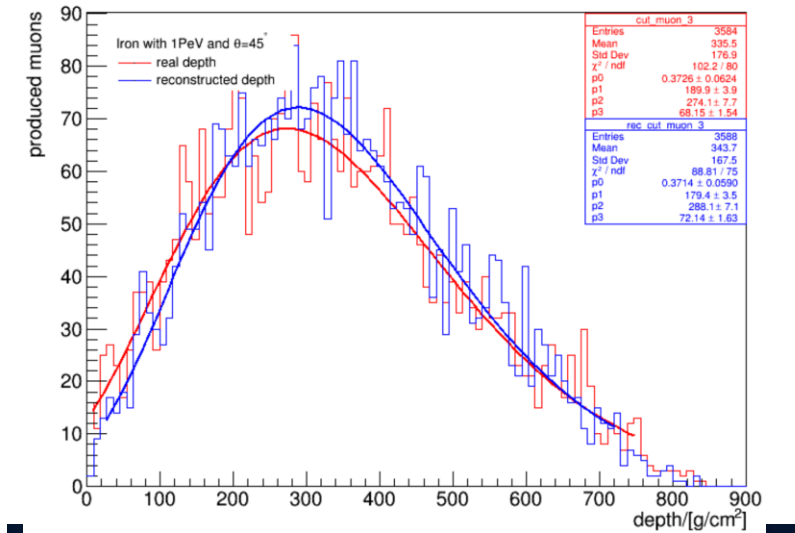
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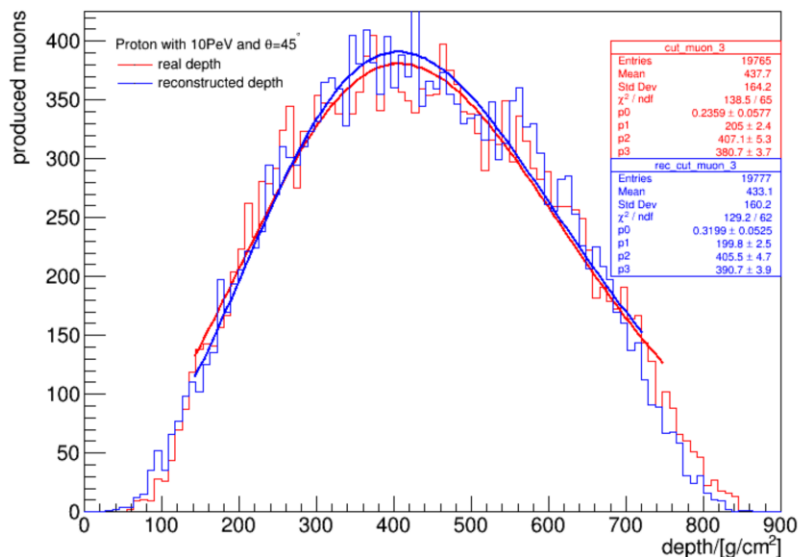
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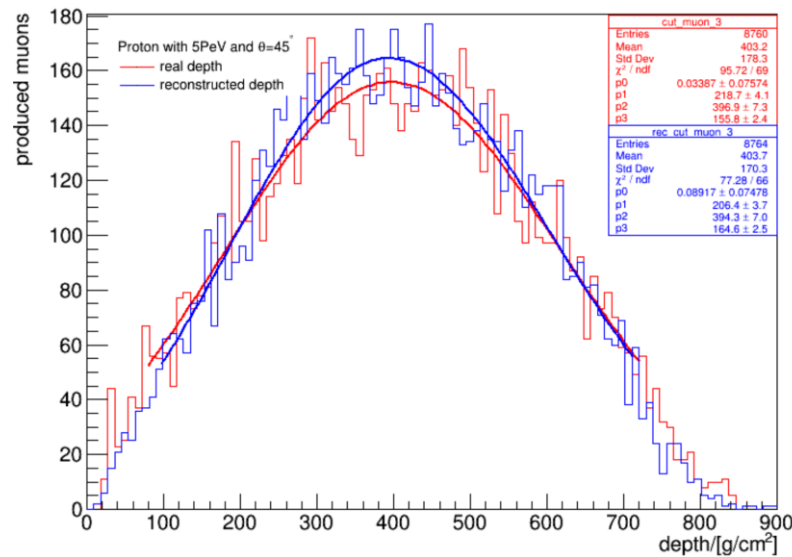
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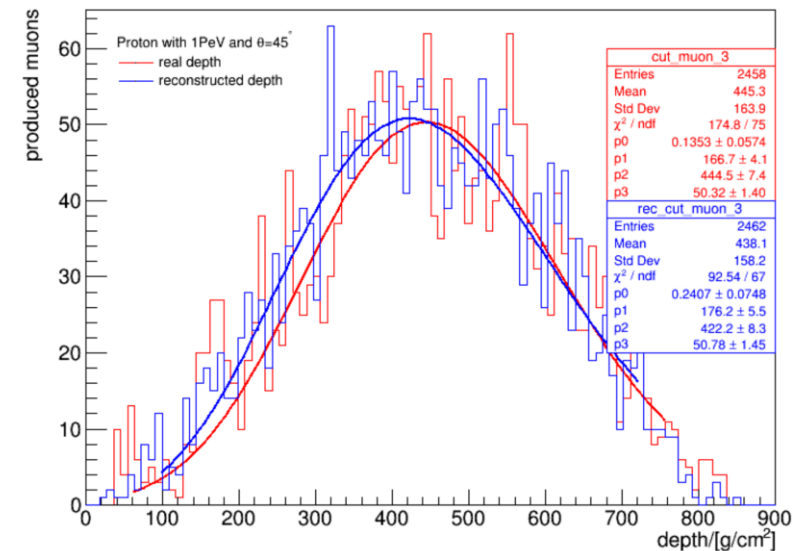
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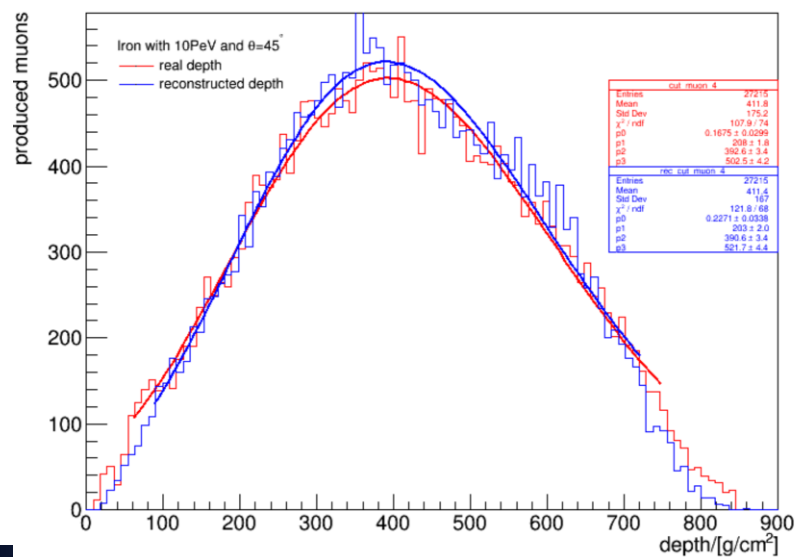
survial muon(>1GeV && 400m<rsfp<1km)



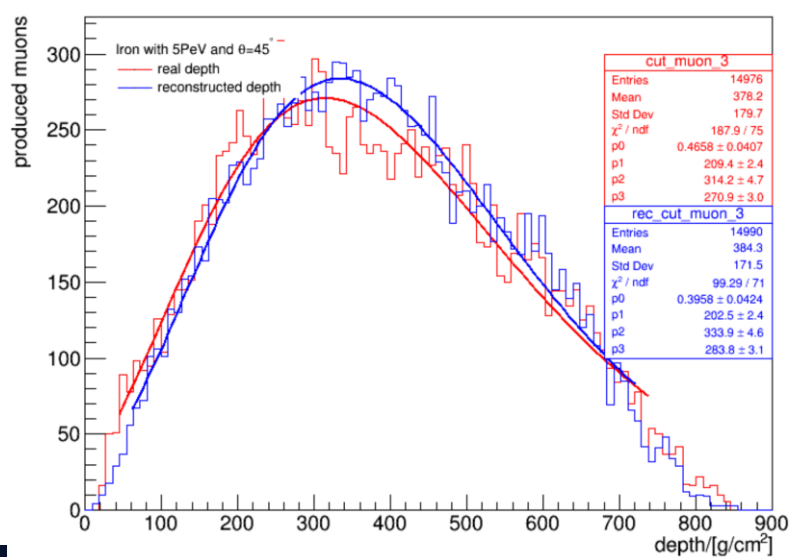
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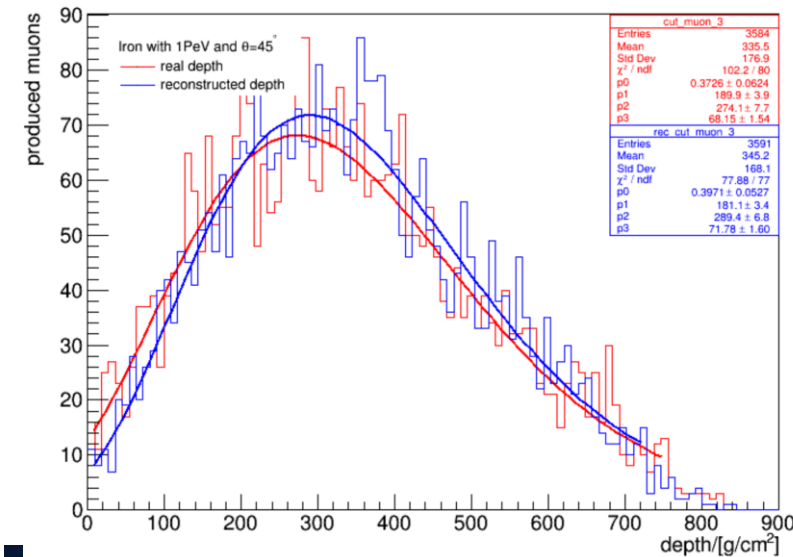
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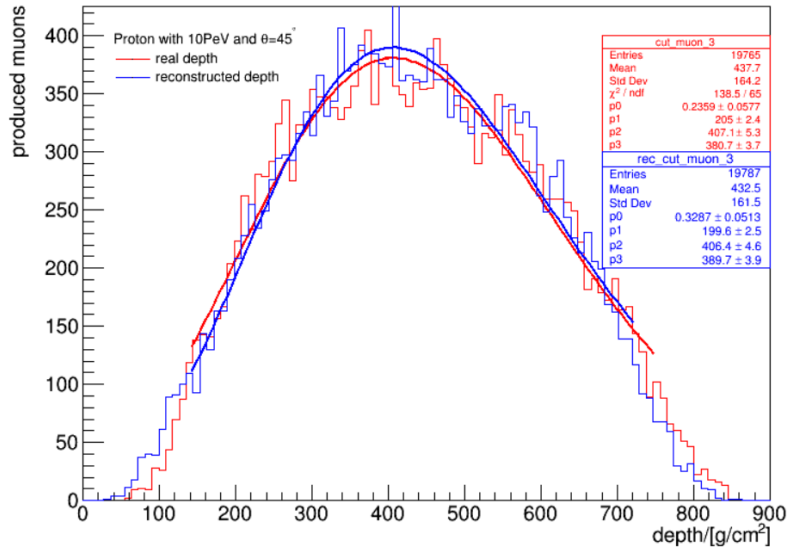
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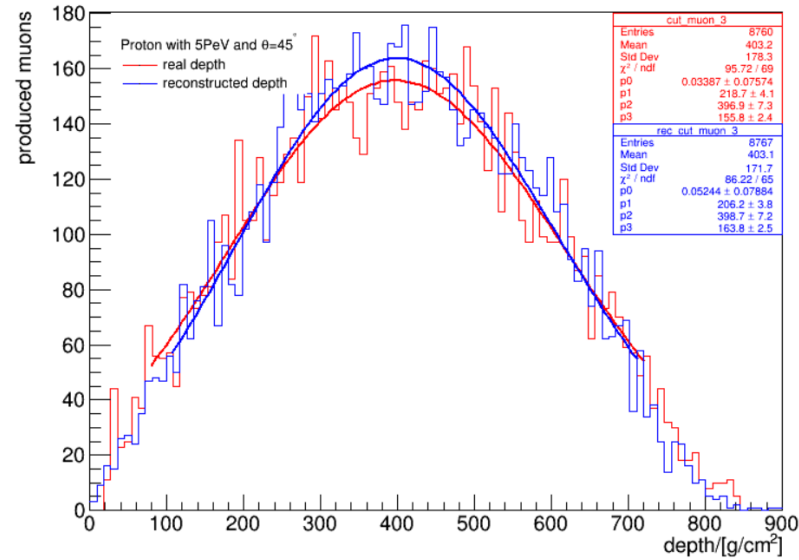
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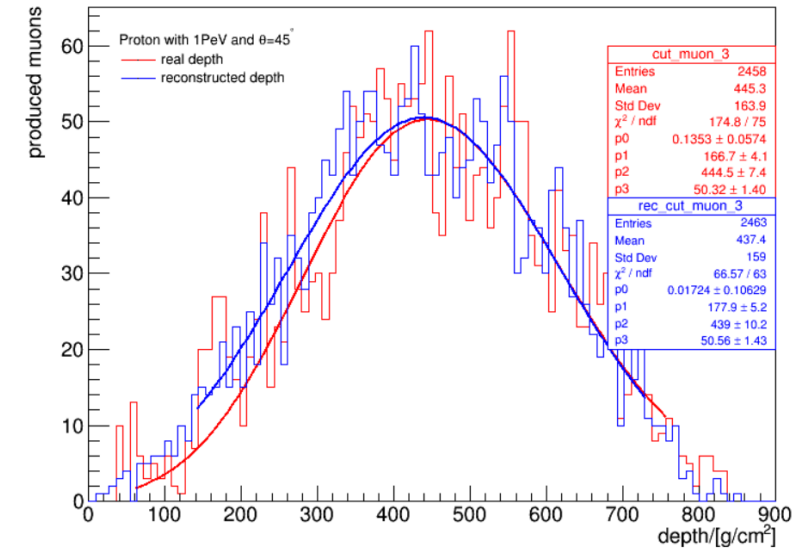
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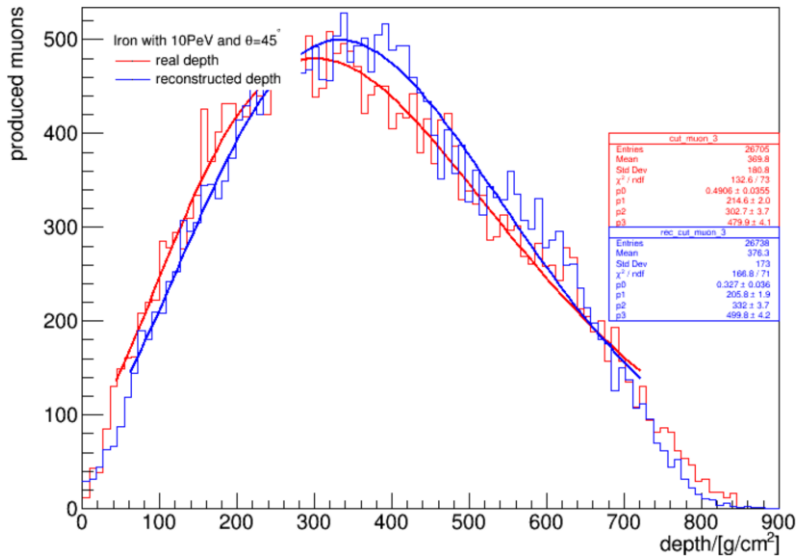
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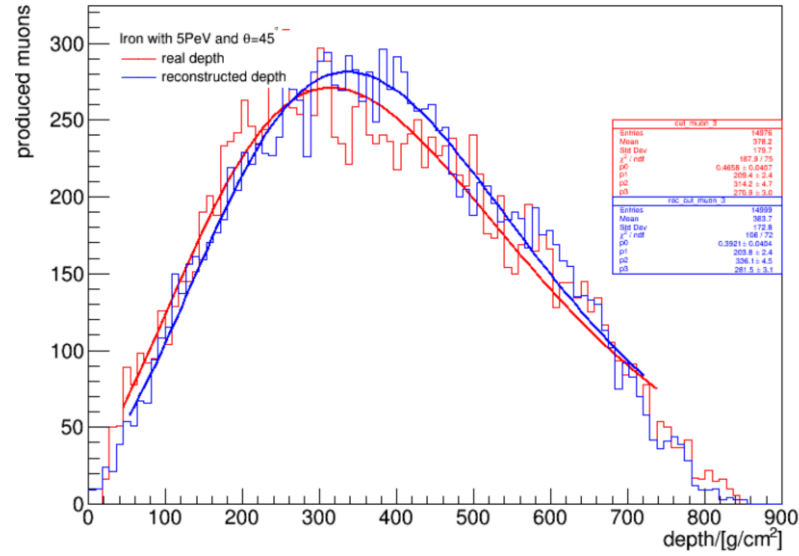
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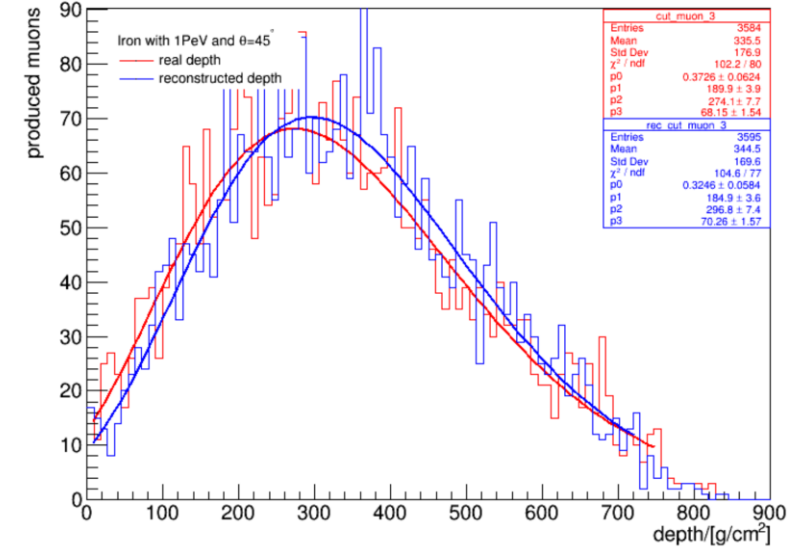
survial muon(>1GeV && 400m<rsfp<1km)



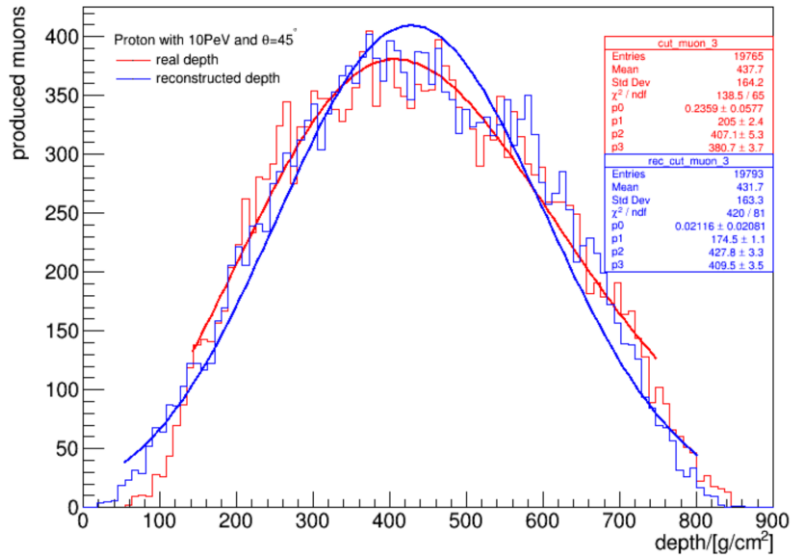
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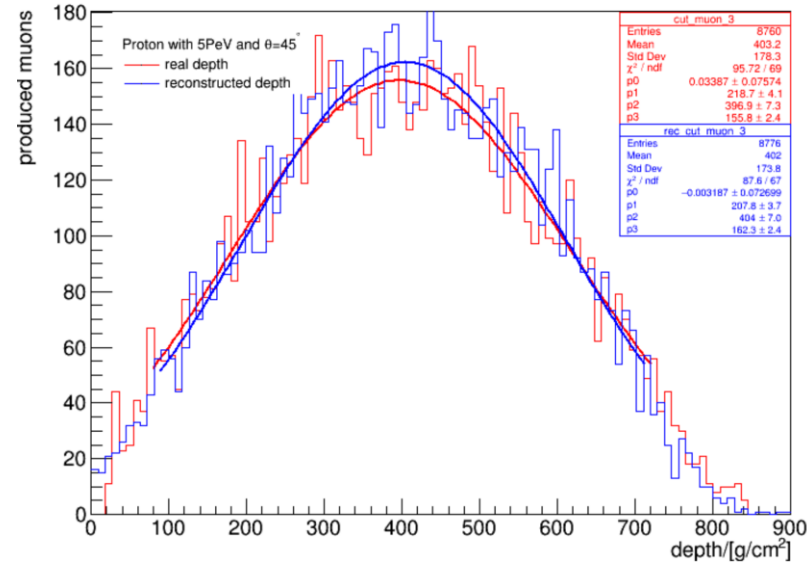
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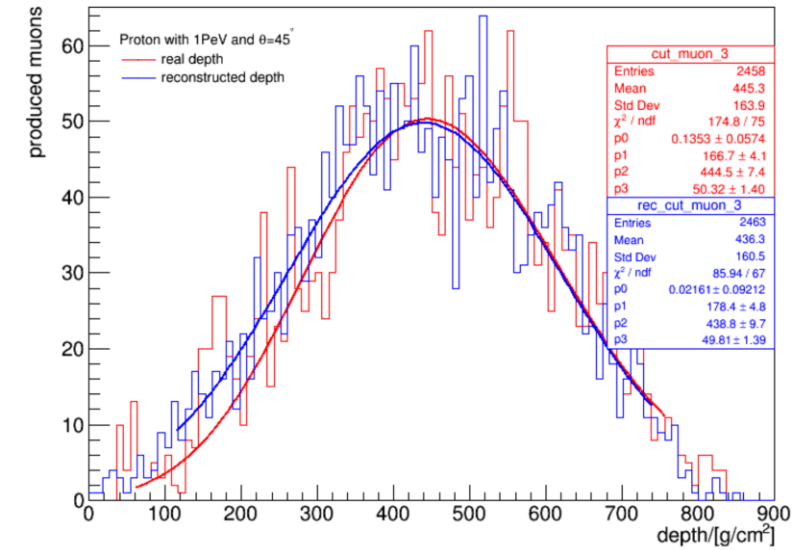
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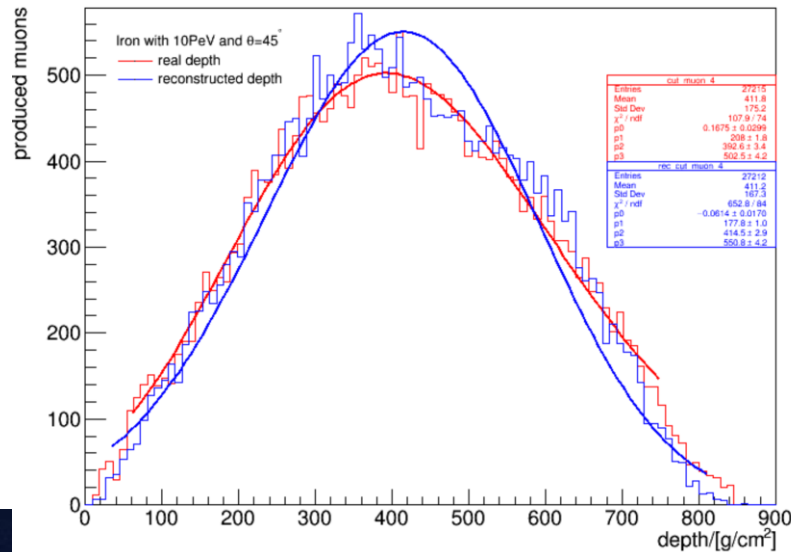
survial muon(>1GeV && 400m<rsfp<1km)



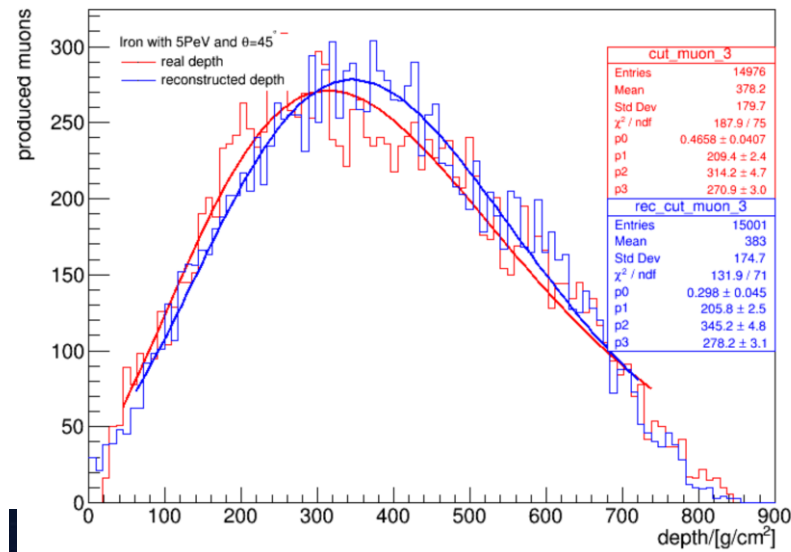
survial muon(>1GeV && 400m<rsfp<1km)



survial muon(>1GeV && 400m<rsfp<1km)



survial muon(>1GeV && 400m<rsfp<1km)



survial muon(>1GeV && 400m<rsfp<1km)

