21cm limits on (*the energy injection of*) dark matter and primordial black holes

& highlights

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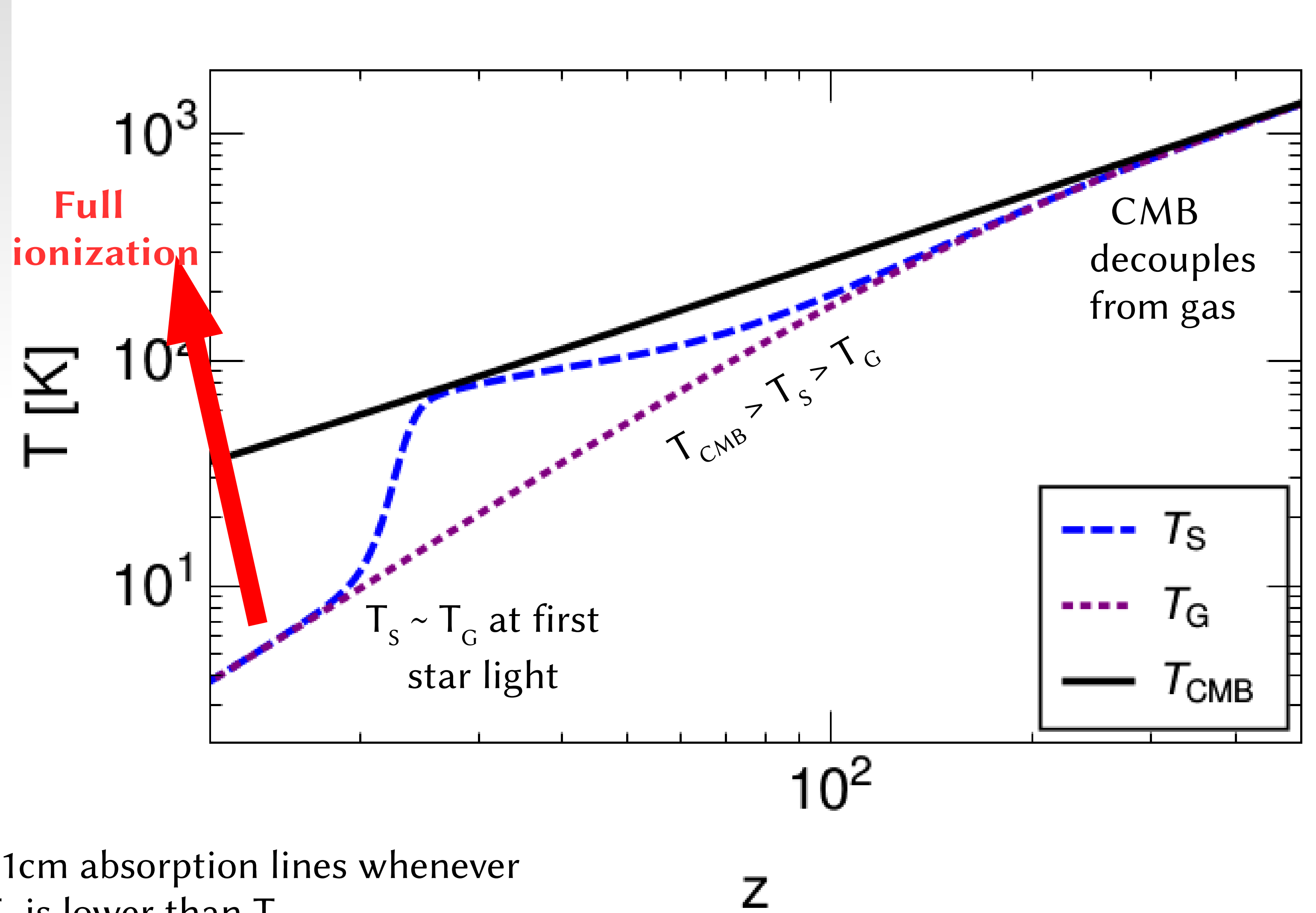
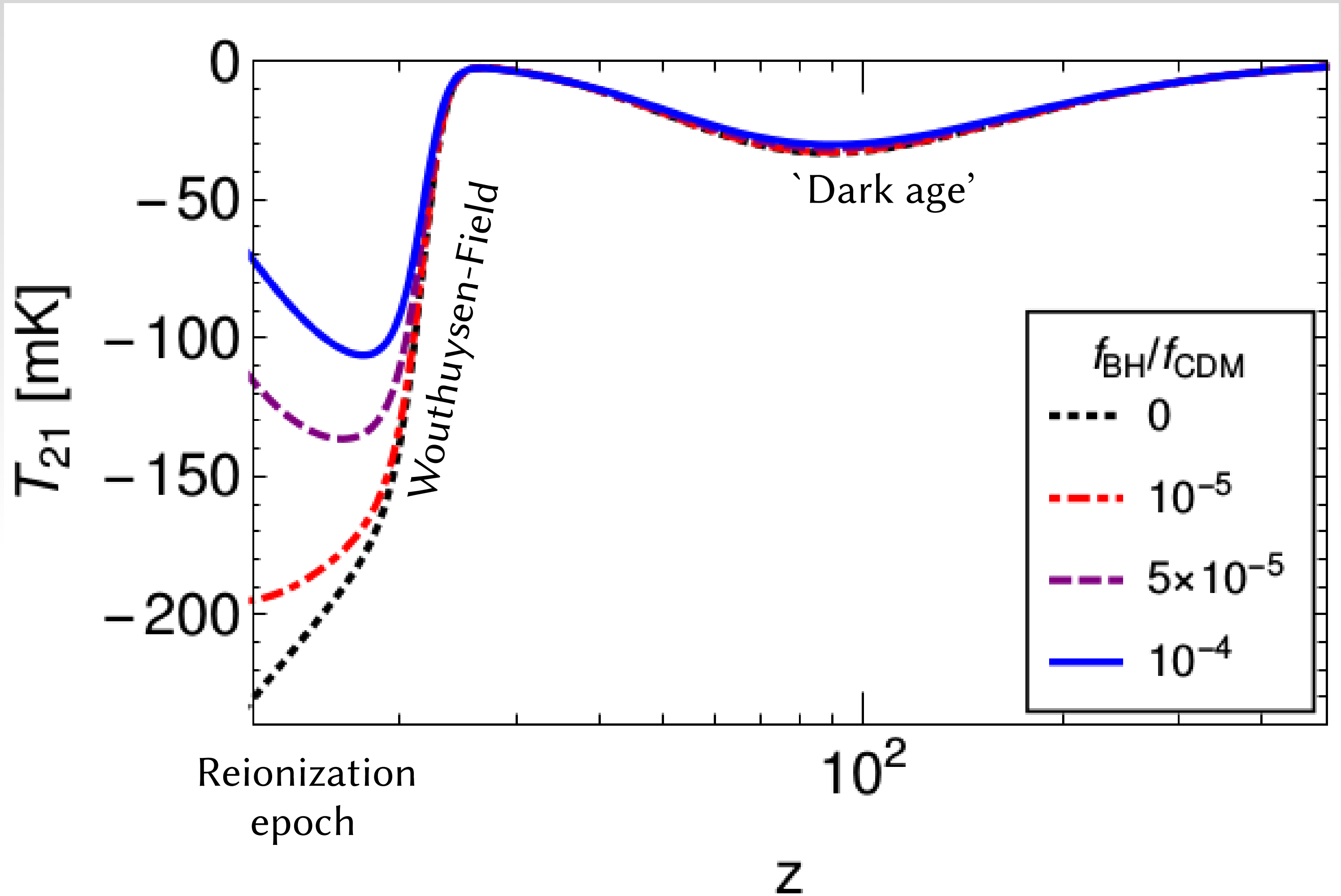
What does 21cm tell us…

- The cosmological reionization history (mean temperature)
- Distribution of neutral Hydrogen gas (temperature map & power spectrum)
  at high reshift (~ reionization epoch, $10<z<30$)
  & late time (post-reionization, $z<1$ galactic gas)

Reionization in the `standard' astrophysics



21cm absorption lines whenever
T_S is lower than T_{CMB}.



The average `brightness temperature',
ignoring over-density and comoving velocity gradients

$$T_{21} \approx 0.023\mathrm{K} \cdot x_{\mathrm{H_I}}(z) \left(\frac{0.15}{\Omega_\mathrm{m}} \cdot \frac{1+z}{10} \right)^{\frac{1}{2}} \frac{\Omega_\mathrm{b} h}{0.02} \left(1 - \frac{T_\mathrm{CMB}}{T_\mathrm{S}} \right)$$

First claim of observation: The EDGES 21cm result

J. D. Bowman, A. E. E. Rogers, R. A. Monsalve,
T. J. Mozdzen, and N. Mahesh, Nature 555, 67 (2018).

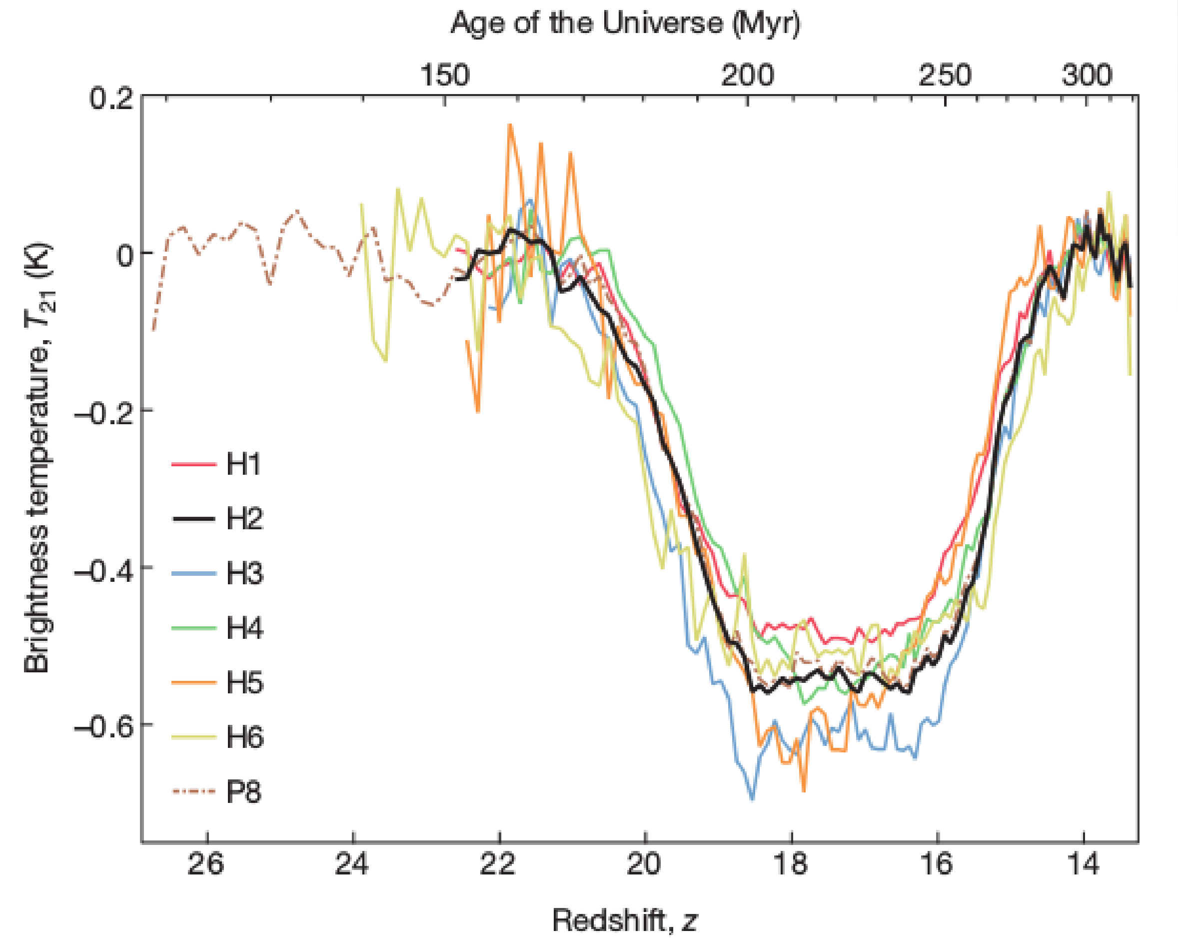


Figure 2 | Best-fitting 21-cm absorption profiles for each hardware case.

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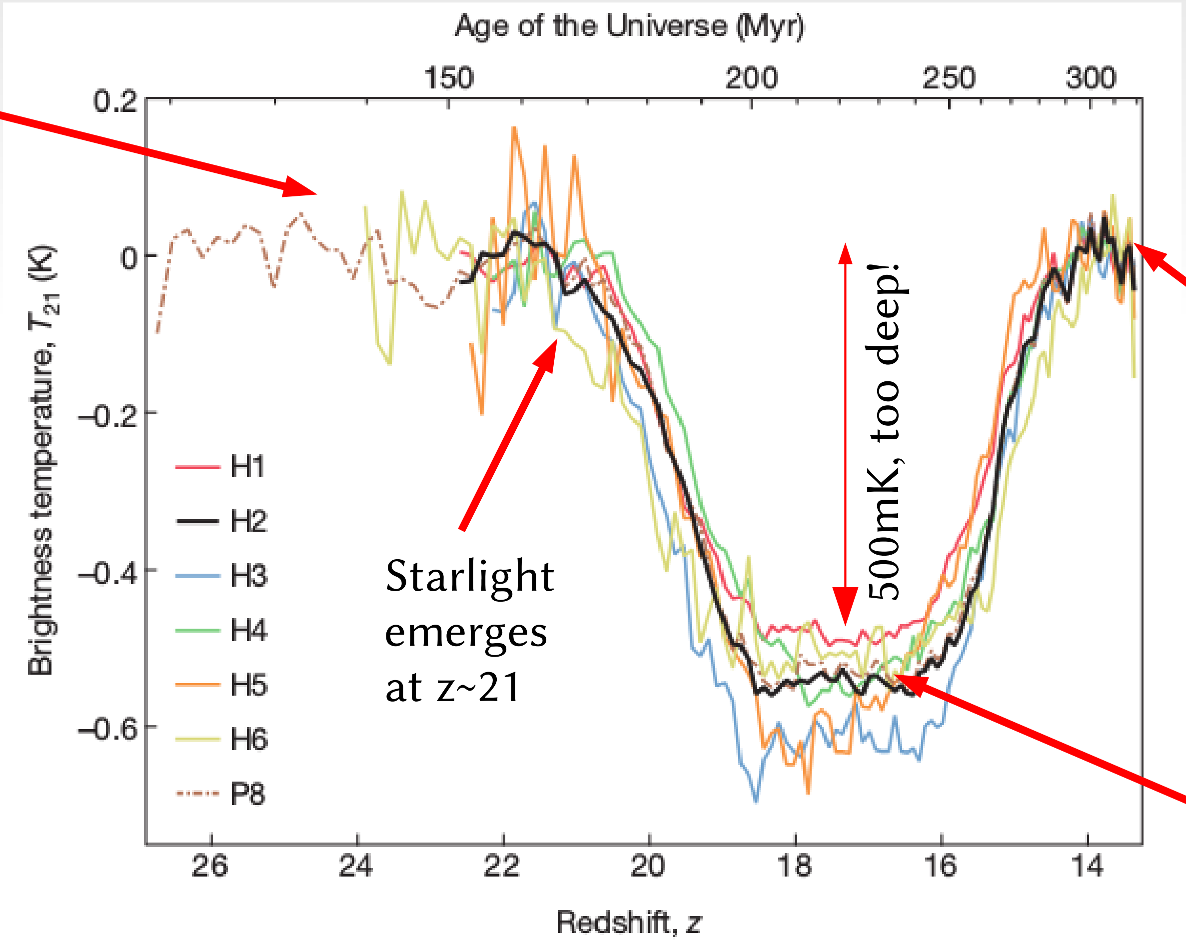
$T_S \sim T_{CMB}$
no signal
for z>21



Starlight
emerges
at z~21

Full ionization
at z~14?

How can this be a
flat valley?

Figure 2 | Best-fitting 21-cm absorption profiles for each hardware case.

The dark matter connection

- Dark matter can cool down T_G via collision

Barkana, et.al. :
A millicharged DM
with $\sigma \sim v^{-4}$

- Dark matter can affect T_S via SD scattering
- Non-CDM affects the 21cm temperature & power spectrum

M. Sitwell, et.al, MNRAS, 438, 2664

- Dark matter, or black holes, can indirectly ruin the 21cm signal if they have emissions, and …

G. D'Amico, et.al.,1803.03629 (DM annilihation)
S. Clark, et.al. 1803.09390 (DM decay & PBH rad.)
K. Cheung, et.al, 1803.09398 (DM annihilation)
A. Hektor, et.al,1803.09697 (PBH accretion)
T. Slatyer, C.-L. Wu, 1803.09734 (DM ann. & decay)

- enhance ionized fraction of gas, affect CMB propagation

T. Slatyer, 1506.03811
T. Slatyer and C.-L. Wu, 1610.06933.
S. Clark, et.al. 1612.07738

The scattering (H, He cooling, etc) mediator must be screened by SM particles

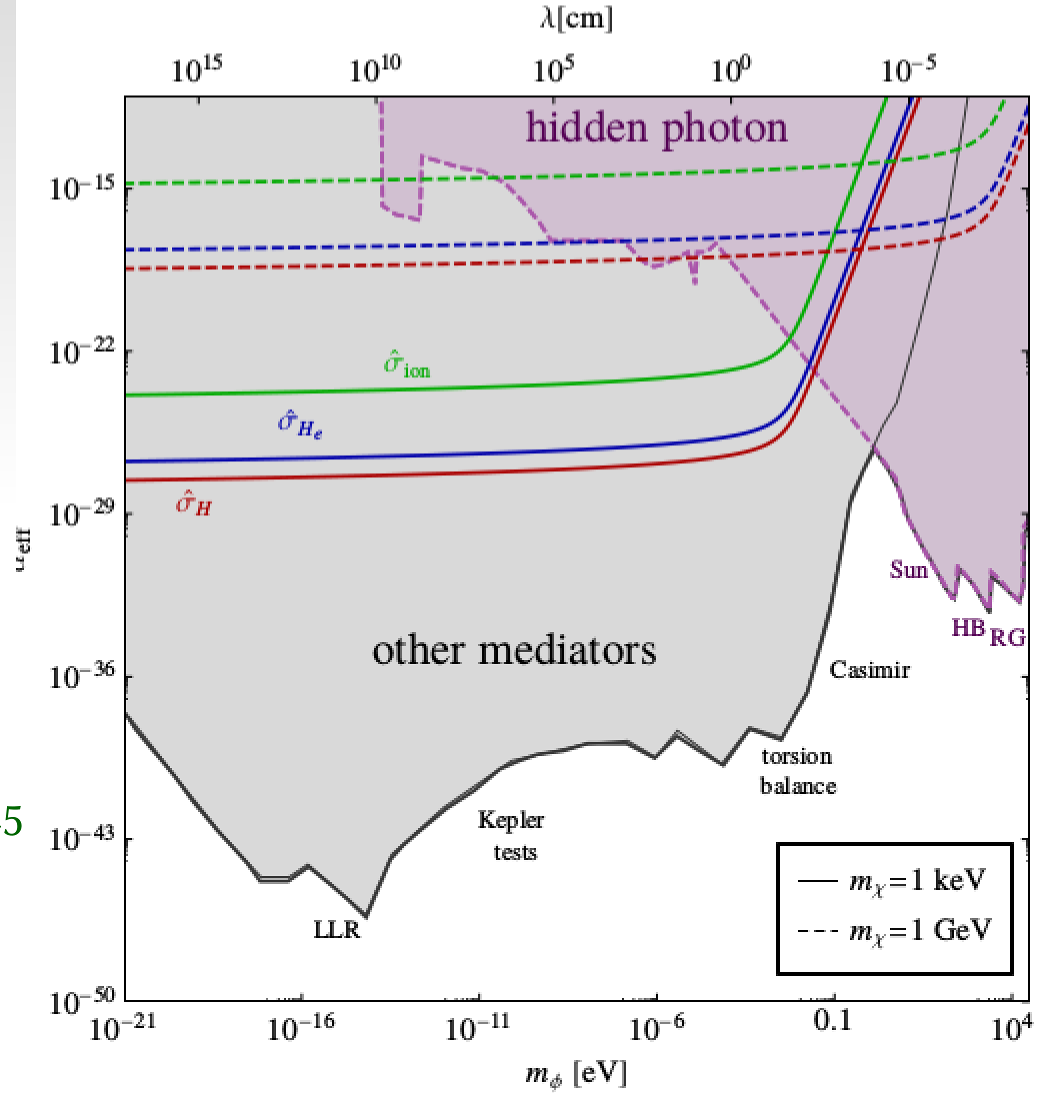
viable candidates: photon & friends

R. Barkana, et.al. 1803.03091

A few more DM cases,

S. Fraser et.al 1803.03245

Z. Kang, 1803.04928

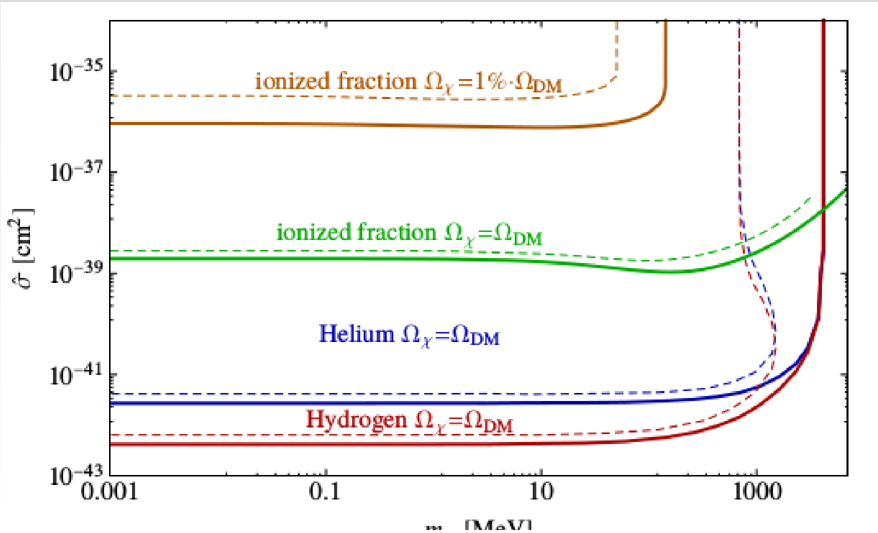


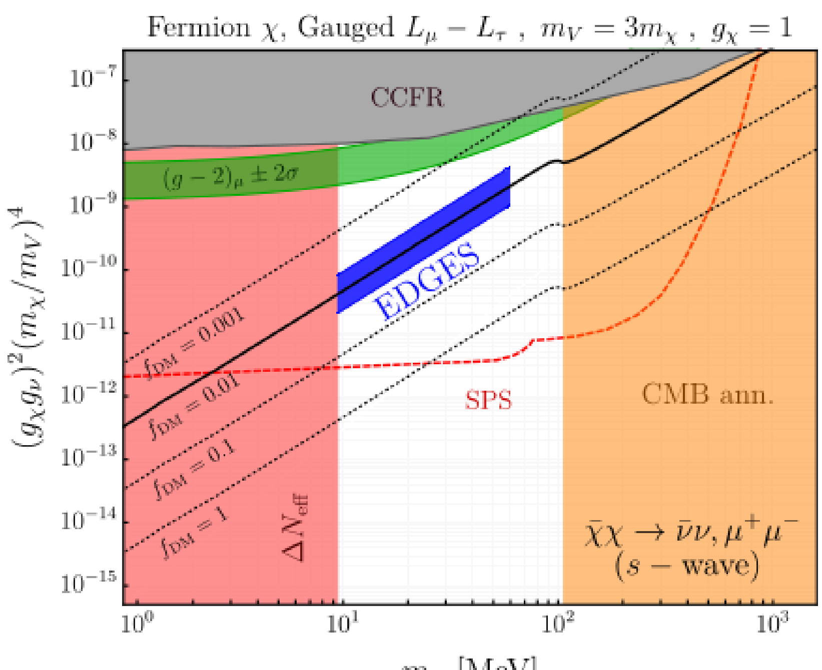
Dark matter as an explanation to the EDGES data

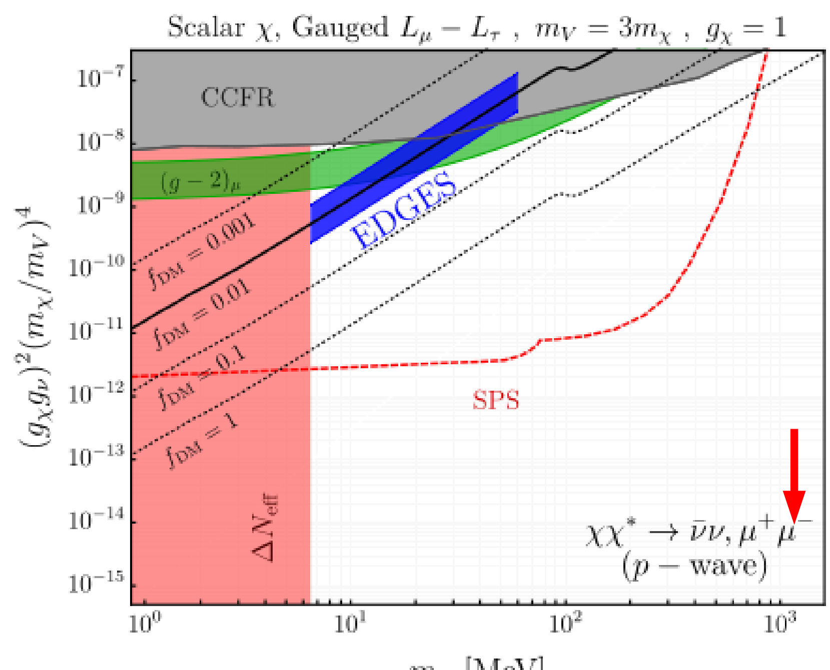
Cooling H_I gas needs a very large scattering cross-section need to avoid constraint from direct detection

– velocity dependence?

R. Barkana, et.al. 1803.03091







You don't want this channel

Millicharged DM need to be at percent level abundance, and relic density requires alternative annihiliation / depletion procecess A.Berlin, et.al. 1803.02804

A few non-DM scenarios:
Earlier T_G-T_{CMB} decoupling

Charge sequestration:

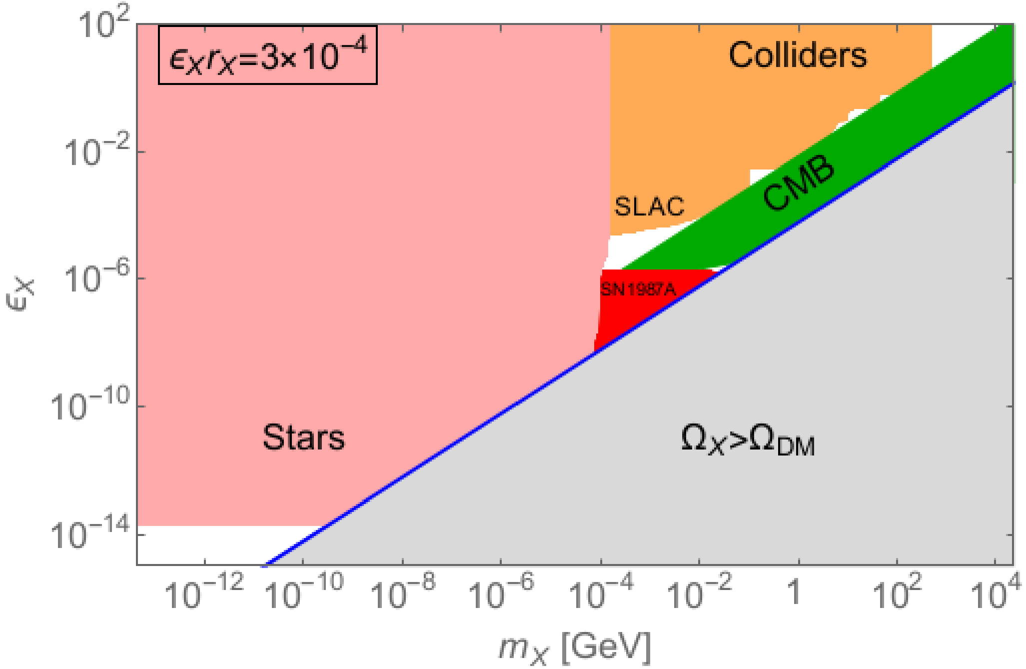
A. Falkowski, K. Petraki, 1803.10096

A negatively charged (-ϵ), stable particle to replace some electrons in the Universe

* reduces x_e during recombination
* fasten CMB – HI decoupling

$$x_p = x_e + \epsilon_X r_X$$

$$x_p \equiv n_p/n_B,\ x_e \equiv n_e/n_B,\ r_X \equiv n_X/n_B$$

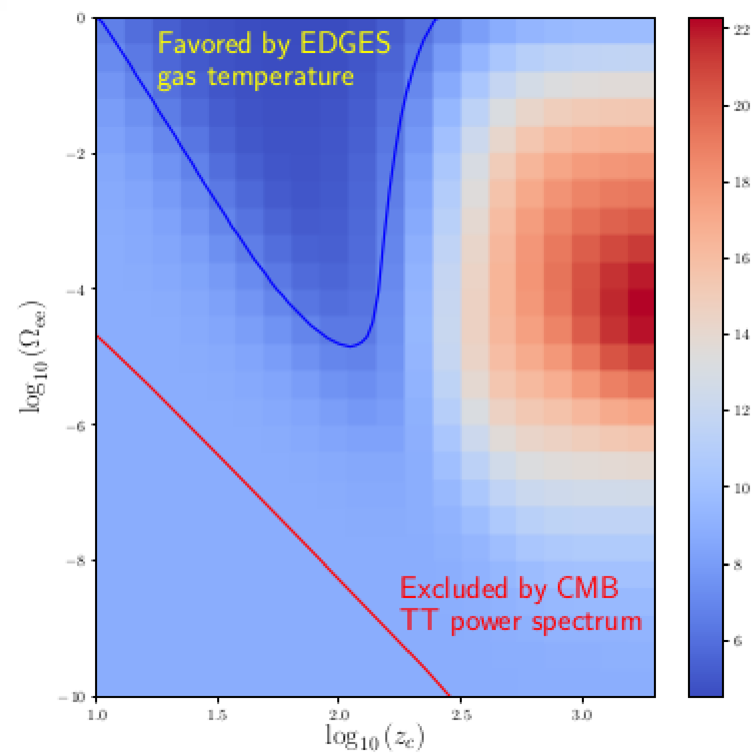


\`Early' dark energy?

J. Colin Hill, E. Baxter. 1803.07555

Addition dark energy component with *w*=-1 and it decays away by z_c ~ 20-1000.

* faster expansion rate \`early on'
* Earlier CMB – HI decoupling



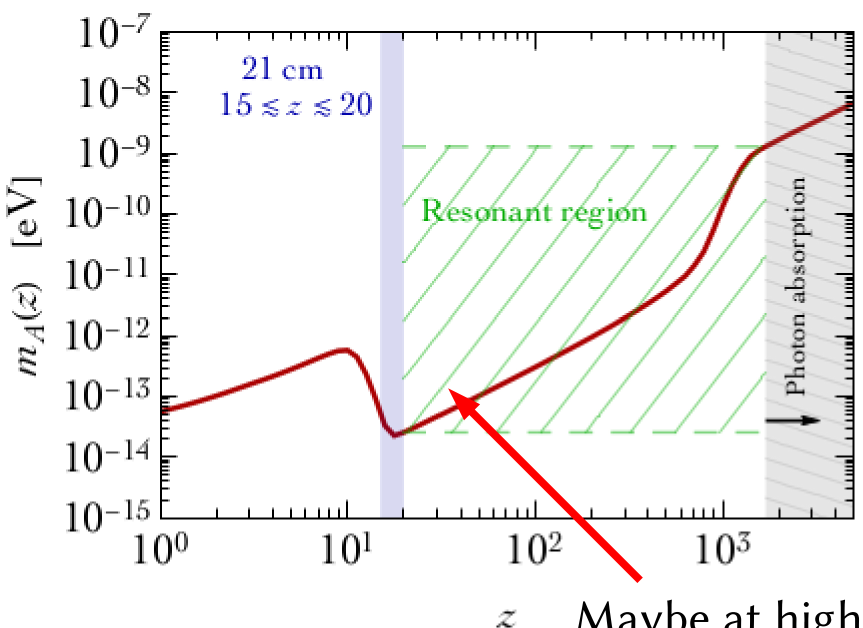
A few non-DM scenarios:

Modification in CMB:

M. Pospelov, et.al. 1803.07048

Order~1 increase in the Raileigh-Jeans tail of CMB: e.g. osc. with a very light dark photon A' via mixing $\epsilon F'_{\mu\nu} F_{\mu\nu}$

Can be of interest in precision 21cm tests, when A' - A oscillate resonantly at the effective A mass (in plasma) $m_A \sim m_{A'}$



Maybe at higher z, too

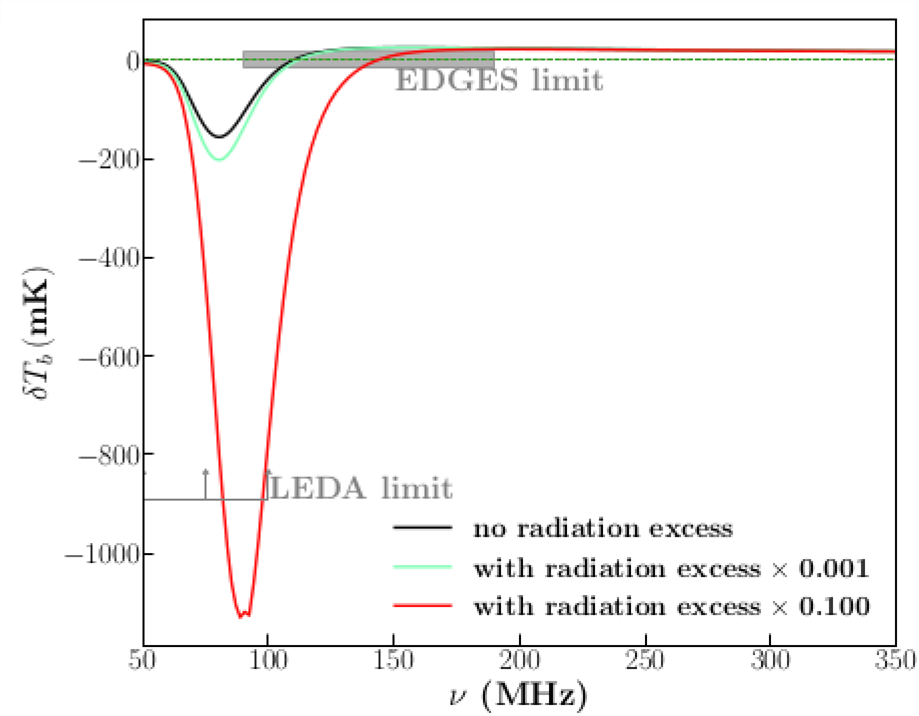
Radio-wavelength backgrounds

C. Feng, G. Holder. 1801.05396

other than CMB is detected by ARCADE

– non black body CMB, enhancing the signal

– EDGES can place a bound on early rad. fields

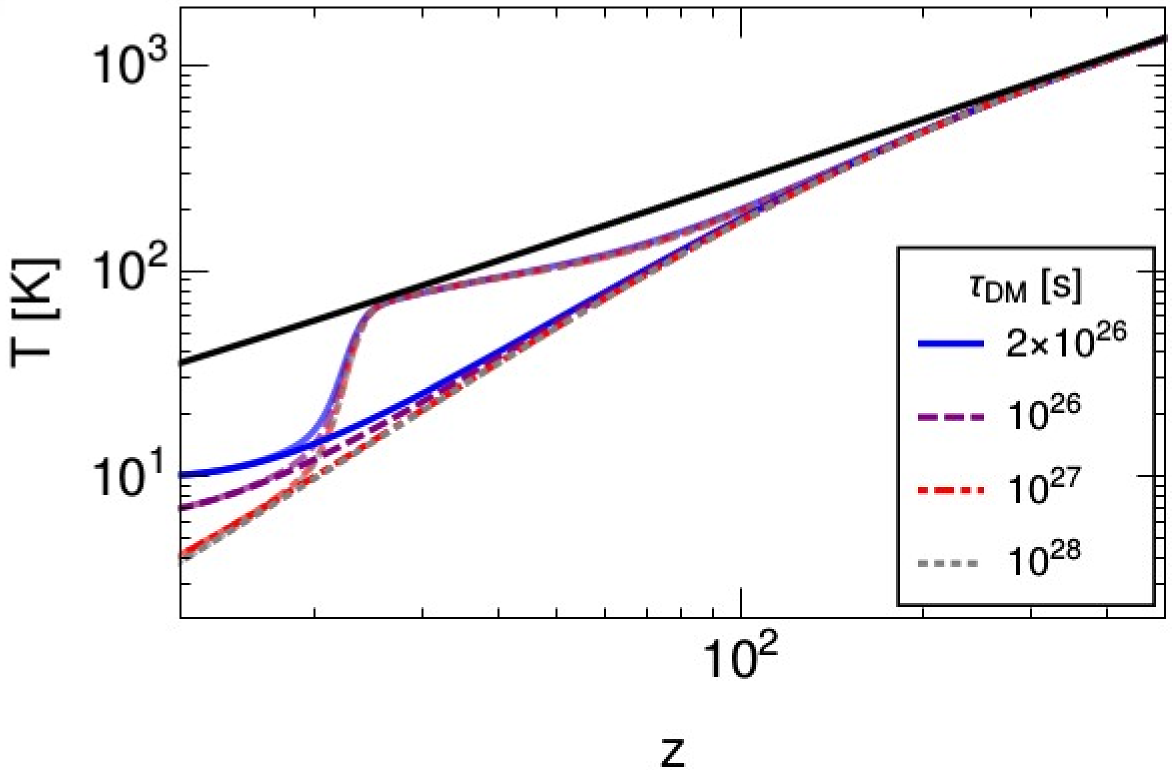


What 21cm can be sensitive to...

- Injection of energy into intergalactic gas suppresses 21cm during reionization epoch
- The observation of (any) 21cm signal means an upper limit
- EDGES may have a lot systematics, future 21cm experiments awaits.

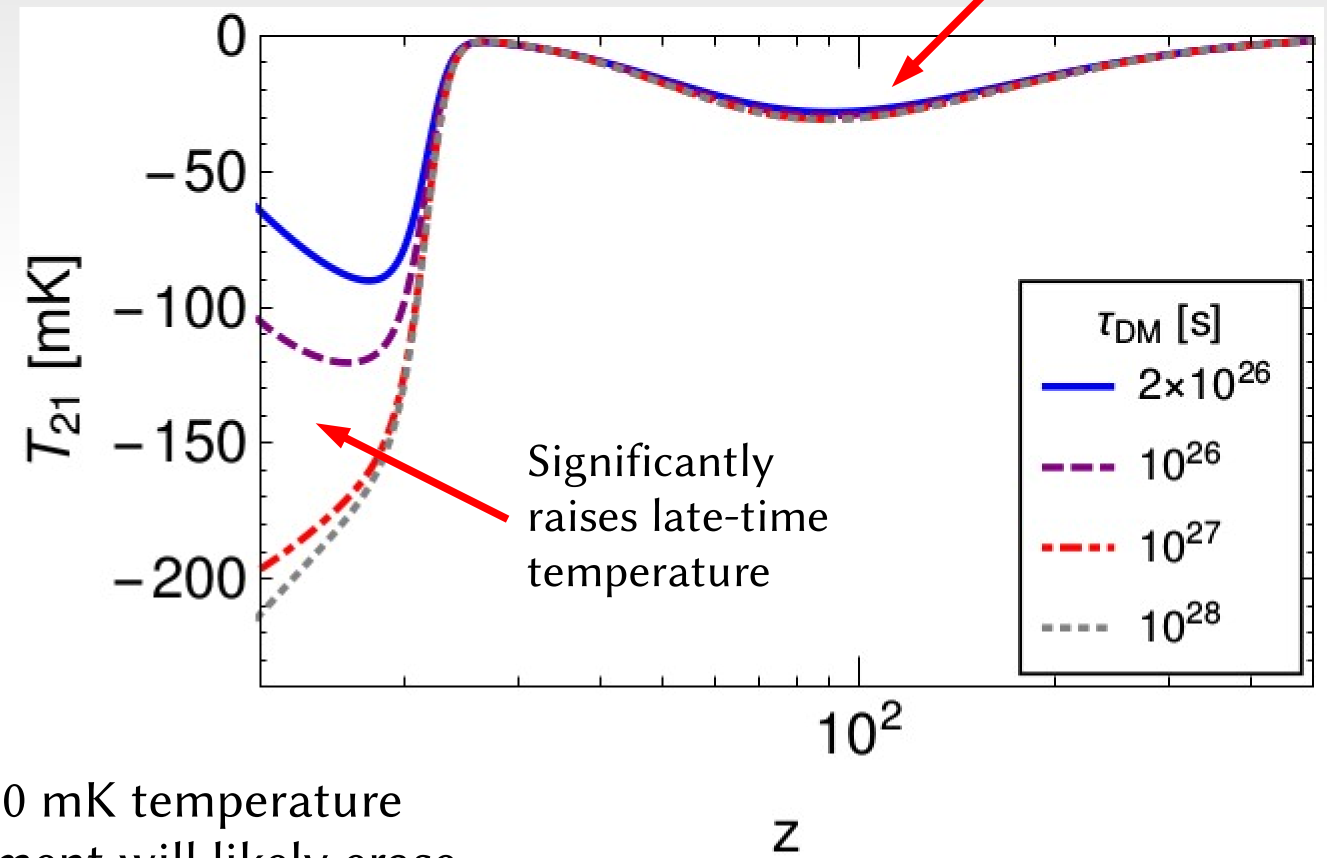
Energy injection effects

- (Historic, cumulative) high-energy injection of electrons/photons
- Increased x_e
- Later CMB – gas decoupling
- Higher T_G
- May reduce or wipe out the 21cm signal



21cm signal suppression by injection

Limited early time effects leave `dark age' alone



Assuming 100% DM decays as DM⟶ e+e-

A 100~200 mK temperature enhancement will likely erase the expected 21cm signal in standard astrophysics

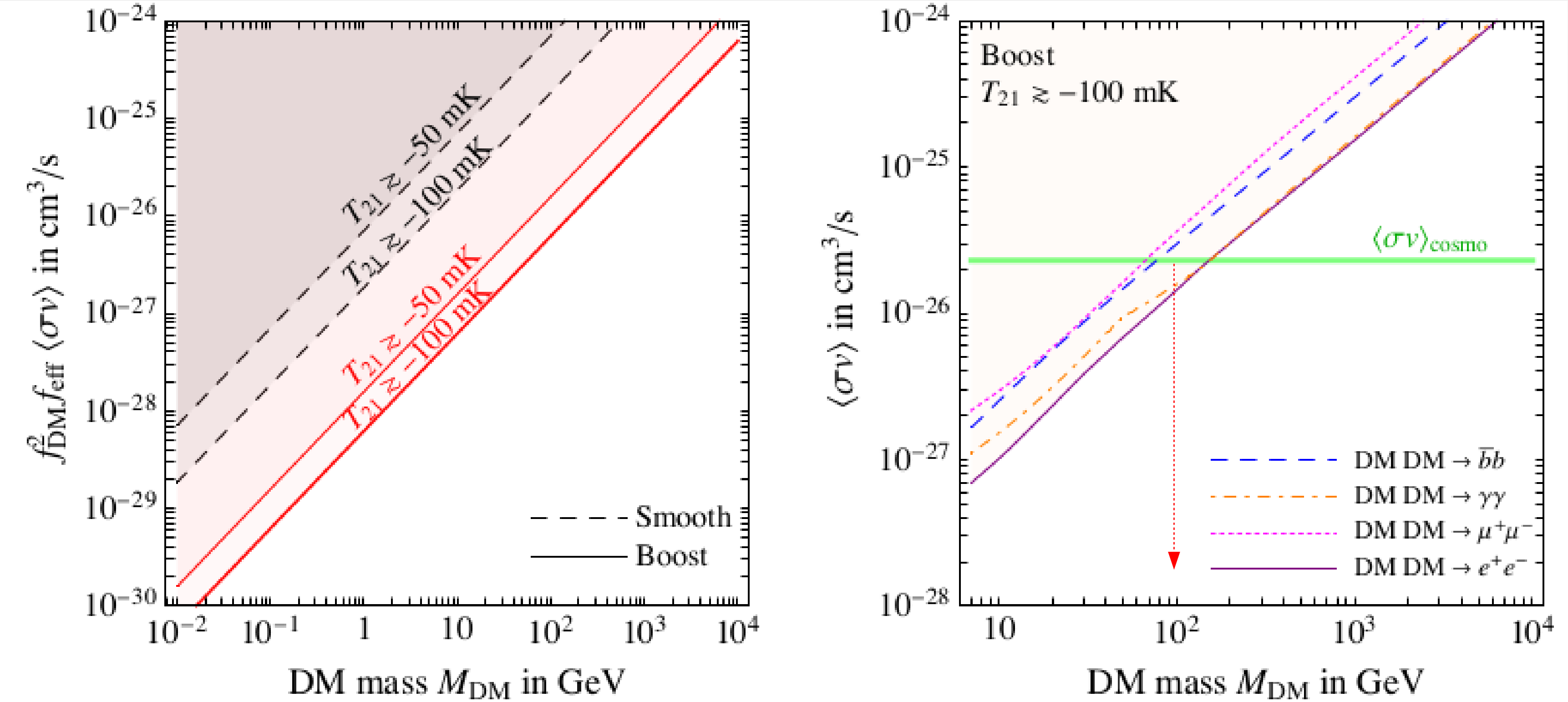
EDGES T_{21} uncertainty:
+200 mK by 95% up-fluctuation

The observation as an upper limit

On DM annihilation rates:
by requiring injection induced
ΔT_{21} < +100 or +150 mK

G. D'Amico, P. Panci, A. Strumia 1803.03629

Excluding vanilla thermal wimp below 10^2 GeV?

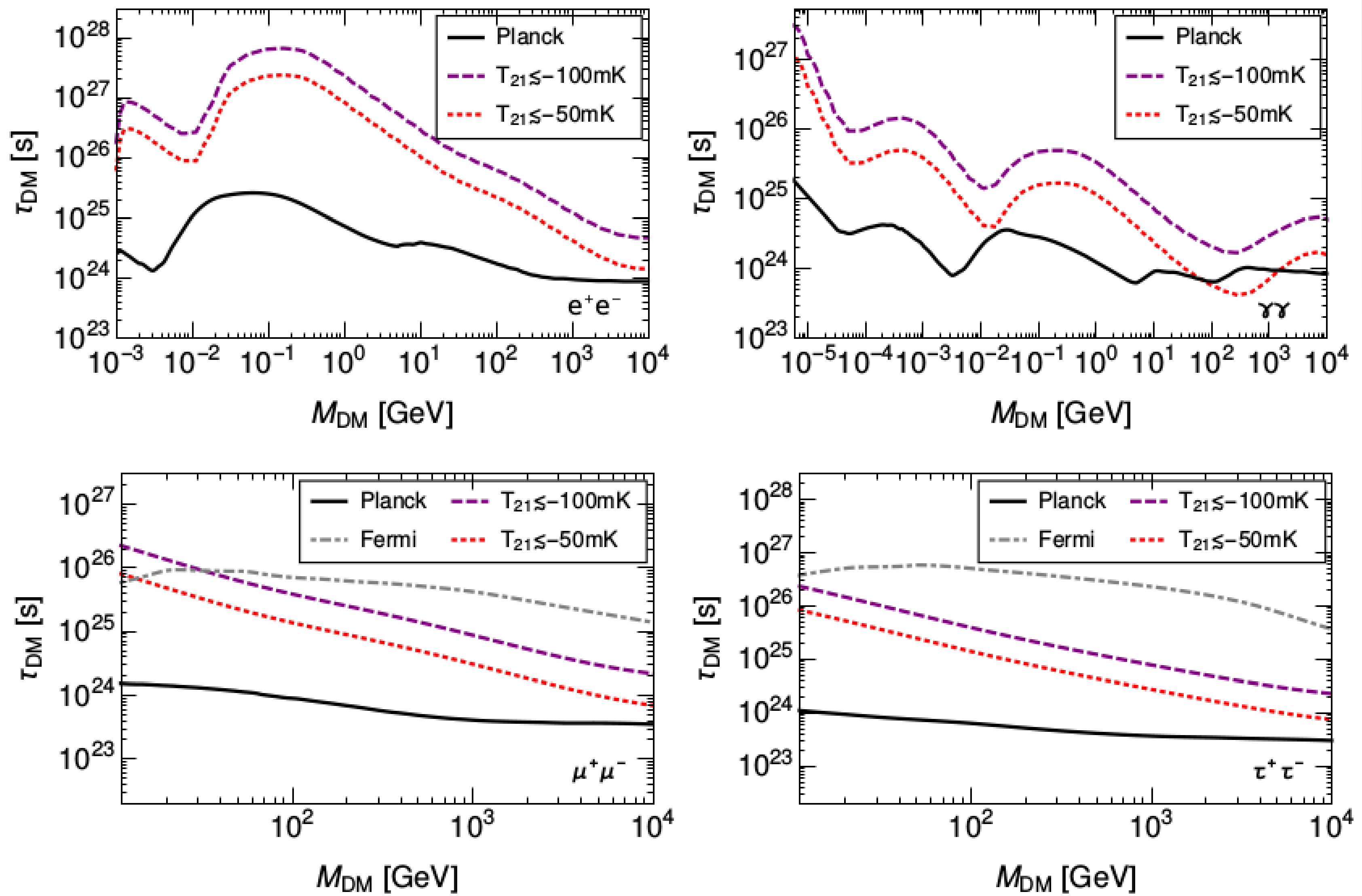


Annihilation: over-density boost B~217 at z~20

Upper limit on DM decay lifetime

ΔT_{21} < +100 or +150 mK at z=17

* 3.5 KeV line allowed



S.Clark, B.Dutta, Y.Gao, Y.-Z.Ma, L.E.Strigari, 1803.09390

Upper limit on DM decay lifetime

ΔT_{21} < +100 or +150 mK at z=17
Gives 1-2 orders of magnitude better bounds on DM injection in comparison to PLANCK TT+TE+EE+lowP data

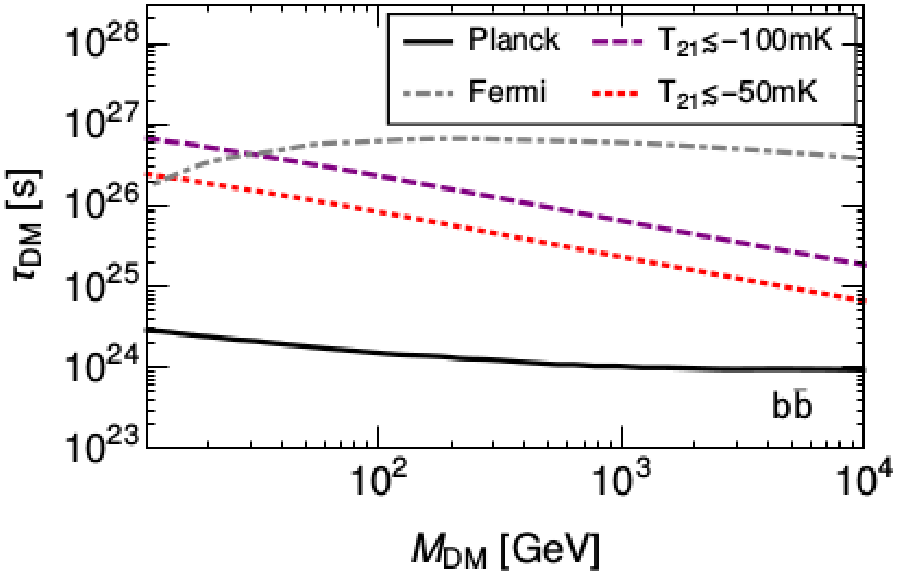
Lower sensitivity at large DM mass due to poorer effective energy loss efficiency

O(0.1-1) correction from cosmological parameter variations (PLANCK)

S. Clark, B. Dutta, Y. Gao, Y.-Z. Ma, L.E.Strigari, 1803.09390

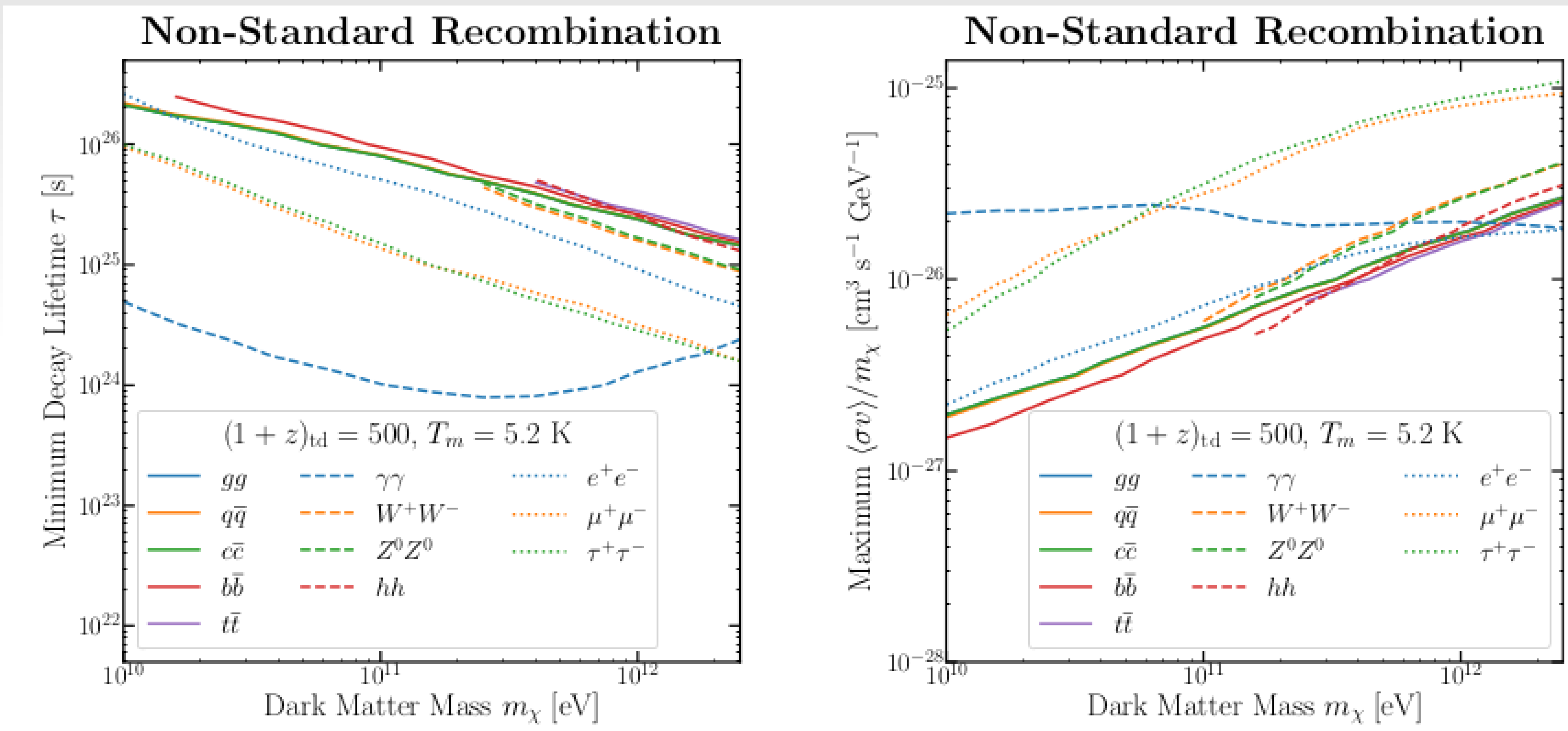
Photon line signals:

Very good limit in KeV range. Testing the 3.5 KeV line needs O(mK) T_{21} sensitivity
Comparable to X-ray line search ~10^{26}s in (sub) MeV range.
Less than Fermi-LAT's pass8 ~10^{29}s in GeV range.



Upper limit on DM annihilation & decays

$T_G(z{=}17.2)$ = 5.2K limit, no cooling

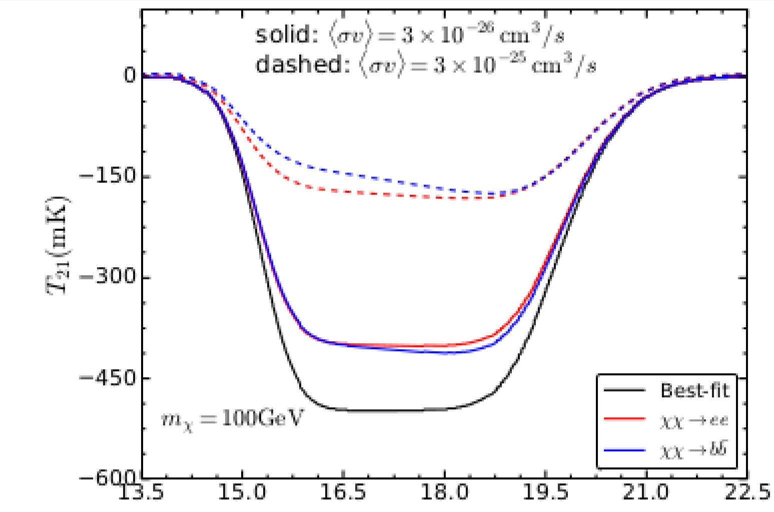


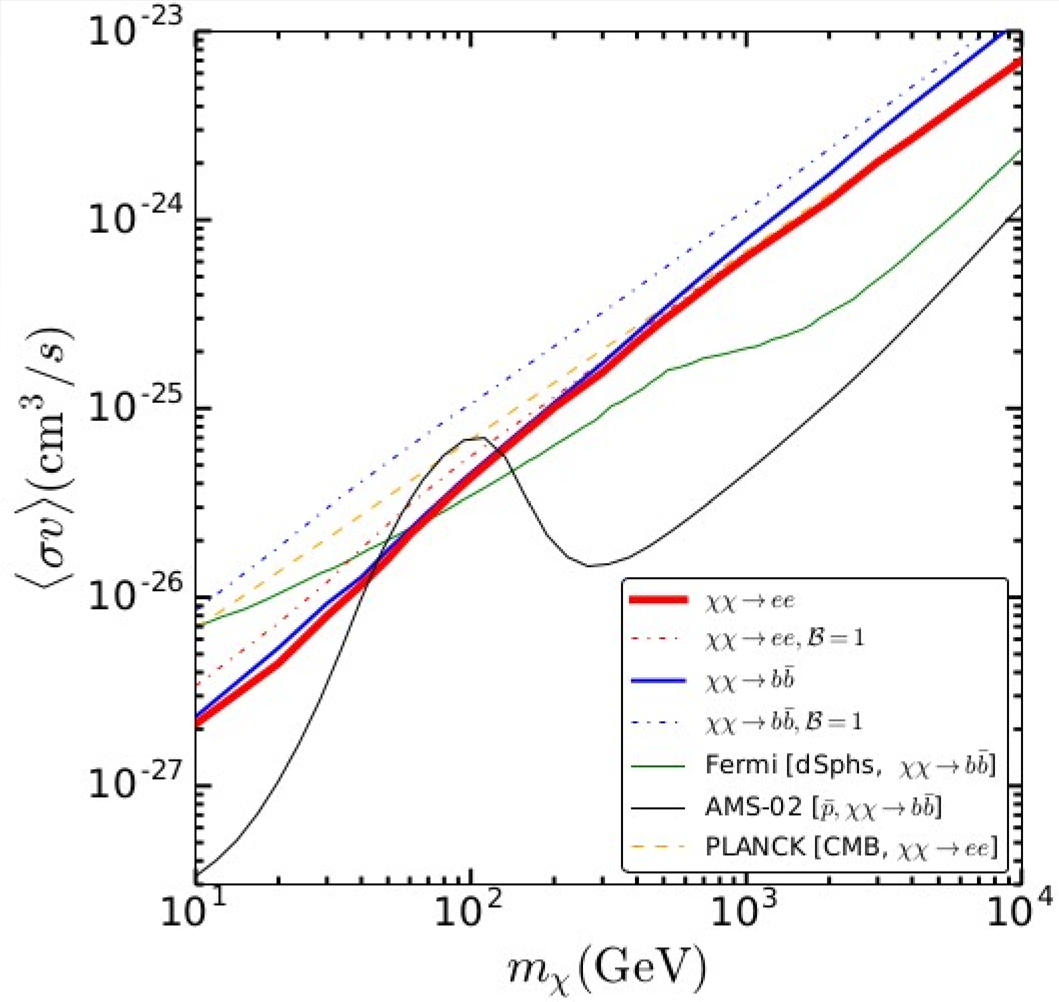
H. Liu, T. R. Slatyer, 1803.09739

Upper limit on DM annihilation

Fitting to the EDGES T_{21} history..

K. Cheung, J-L Kuo, K-W Ng, Y-L S Tsai, 1803.09398

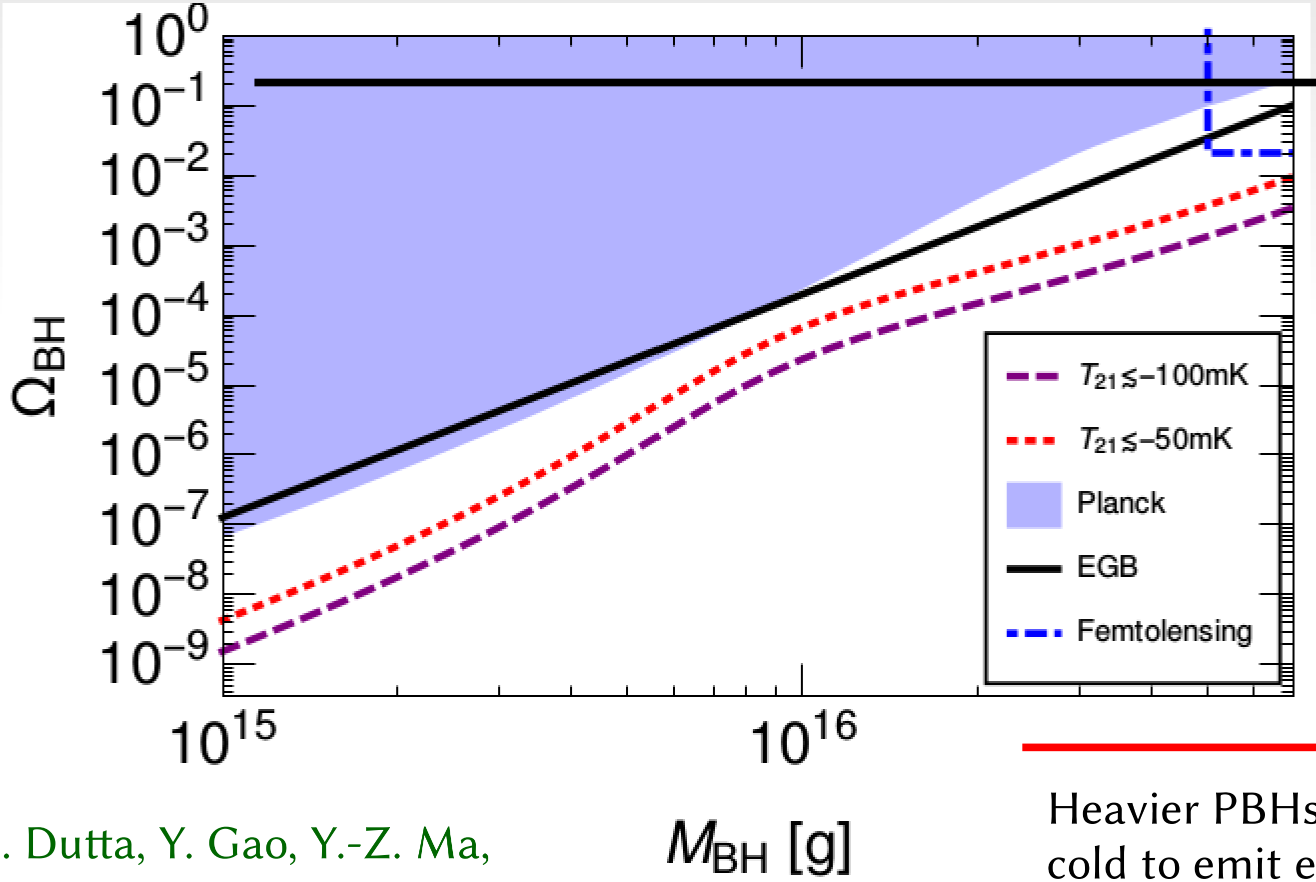




Upper limit on primordial BH's Hawking radiation

ΔT_{21} < +100 or +150 mK at z=17

Applicable to long-lived ($m>10^{15}$g), evaporating black holes



relic density

Heavier PBHs are too cold to emit electrons and have lower (~M^{-3}) evaporation rate

S. Clark, B. Dutta, Y. Gao, Y.-Z. Ma, L.E.Strigari, 1803.09390

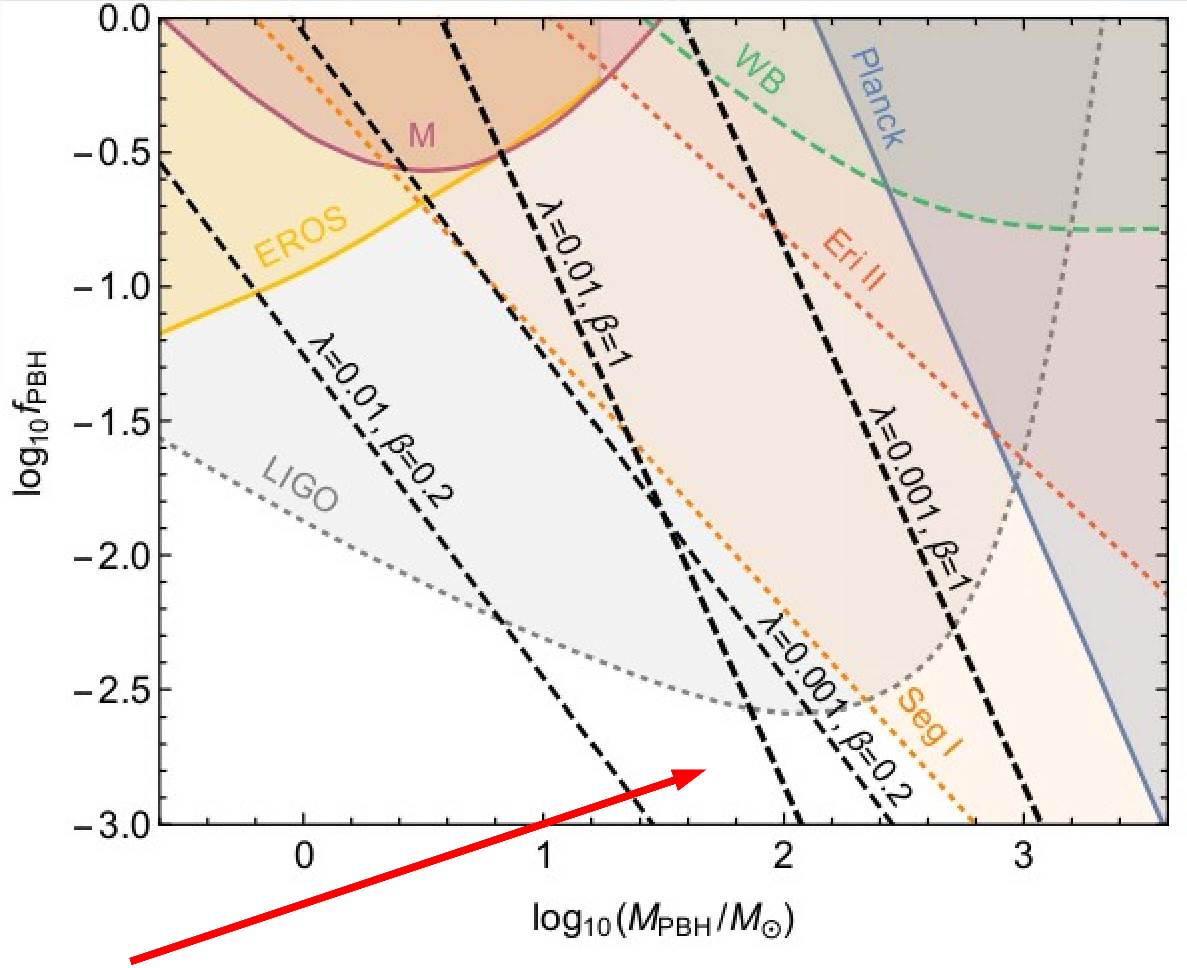
Upper limit on massive PBH's accretion

X-ray emission from BH accretion

Limit from $T_G(z\sim17.2)=8K$

Improvement over PLANCK by 1-2 orders of mag.

A. Hektor, G. Hütsi, L. Marzola, M. Raidal, V. Vaskonen, H. Veermäe, 1803.09697



Complementing the LIGO

Summary

- EDGES results can be interpreted in many ways
- Exitence of reionization epoch 21cm signal imposes a strong limit on historic energy injection
- Like CMB, injection bounds extend to low energy/mass
- 21cm bounds on DM, PBH are very powerful, even at the current proof-of-principle estimates
- Future 21cm experiments await.

Backup: Injection corrections

Extra contributions to ionization & heating

- Reduce neutral hydrogen fraction

$$\frac{dx_e}{dz} = \left(\frac{dx_e}{dz}\right)_{\rm orig} - \frac{1}{(1+z)H(z)}(I_{Xi}(z) + I_{X\alpha}(z))$$

Ionization: $I_{Xi}(z) = f_i(E,z)\frac{dE/dVdt}{n_H(z)E_i}$

Lyman-α excitation $I_{X\alpha}(z) = f_\alpha(E,z)(1-C)\frac{dE/dVdt}{n_H(z)E_\alpha}$

- IGM temperature

$$\frac{dT_{\rm IGM}}{dz} = \left(\frac{dT_{\rm IGM}}{dz}\right)_{\rm orig} - \frac{2}{3k_B(1+z)H(z)}\frac{K_h}{1+f_{\rm He}+x_e}$$

$$K_h(z) = f_h(E,z)\frac{dE/dVdt}{n_H(z)}$$

- Wouthuysen-Field

$$T_{\rm S} = \frac{T_{\rm CMB} + y_{\rm c}T_{\rm G} + y_{\rm Ly\alpha}T_{\rm Ly\alpha}}{1 + y_{\rm c} + y_{\rm Ly\alpha}},$$

$$y_{\rm c} = \frac{C_{10}}{A_{10}}\frac{T_\star}{T_{\rm G}},$$

$$y_{\rm Ly\alpha} = \frac{P_{10}}{A_{10}}\frac{T_\star}{T_{\rm Ly\alpha}},$$

For Lyman-alpha during reionization,
See B. Ciardi and P. Madau, astro-ph/0303249

- Energy deposit rate (ionization, excitations, heating)

$$I_{\mathrm{X_i}}(z) = \frac{f_{\mathrm{i}}(E,z)}{\mathrm{H}_H(z)E_{\mathrm{i}}}\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t},$$

$$I_{\mathrm{X}_\alpha}(z) = (1-C)\frac{f_\alpha(E,z)}{n_{\mathrm{H}}(z)E_\alpha}\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t},$$

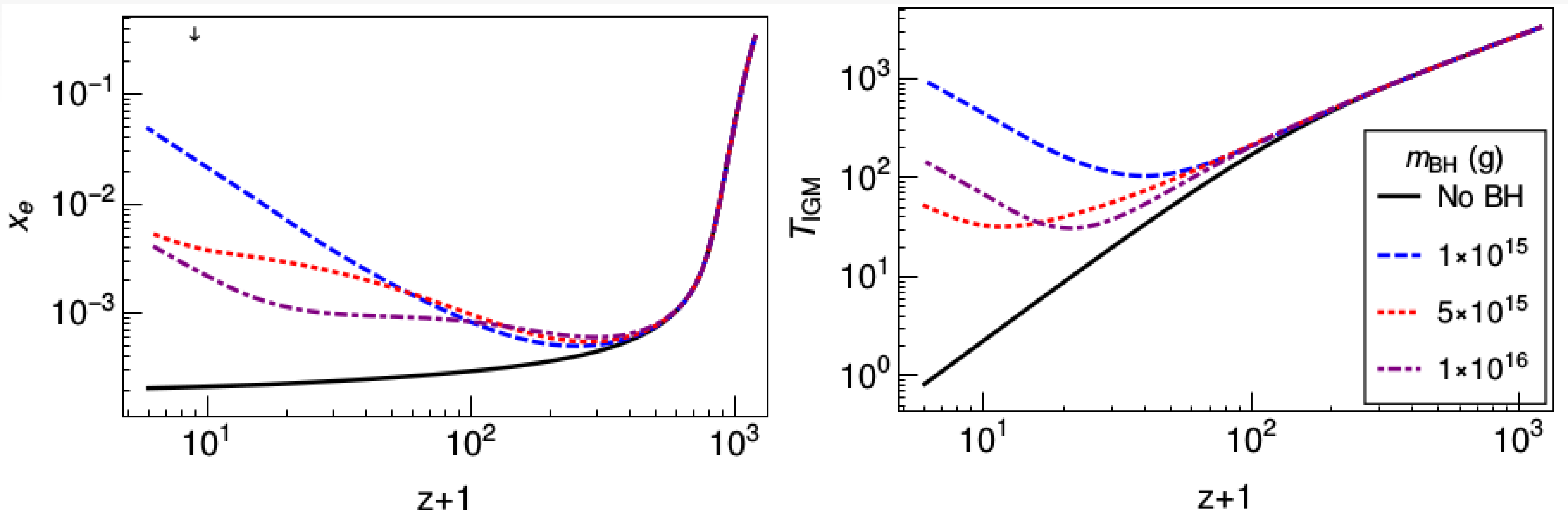
$$K_{\mathrm{h}}(z) = \frac{f_{\mathrm{h}}(E,z)}{n_{\mathrm{H}}(z)}\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}$$

$$C = \frac{1+K\Lambda_{2s,1s}n_{\mathrm{H}}(1+x_{\mathrm{e}})}{1+K\Lambda_{2s,1s}n_{\mathrm{H}}(1-x_{\mathrm{e}})+K\beta_{\mathrm{B}}n_{\mathrm{H}}(1-x_{\mathrm{e}})}.$$

- The `effective’ rate *f*(E,z) is cumulative of historic injection
- Electrons are more effective in energy deposit
- *f* is averaged over injection spectra and species

$$f_c(m_{\mathrm{DM}},z) = \frac{\sum_s \int f_c(E,z,s)E(\mathrm{d}N/\mathrm{d}E)_s\mathrm{d}E}{\sum_s \int E(\mathrm{d}N/\mathrm{d}E)_s\mathrm{d}E},$$

- Ionization history can be a powerful test (at z>7) on `prolonged’ electron & gamma ray injection from new physics



Num. computation by HyRec

Backup: PBHs (CMB)

Hawking radiation, `lifetime' and BH mass

- BH evaporates at a temperature

Hawking 75'

$$T_{PBH} = \frac{1}{8\pi GM} = 1.06\text{TeV} \times \frac{10^{10}\text{g}}{M_{PBH}}$$

- with a peak energy of radiation

$$E_\gamma = 5.71 T_{PBH}\,,\ E_\nu = 4.22 T_{PBH}\,,\ E_{e^\pm} = 4.18 T_{PBH}$$

Mass loss rate:

$$\dot{M}_{10} = -5.34 \times 10^{-5} \left(\sum_i f_i \right) M_{10}^{-2}\ \text{s}^{-1}$$

lifetime:

$$\tau(M) \sim \frac{G^2 M^3}{\hbar c^4} \sim 10^{64} \left(\frac{M}{M_\odot} \right)^3 \text{yr}$$

- BH evaporation can be a good source of cosmic rays, injection particle species determined by BH mass

A steady radiation injection below ~100 MeV

- Relevant for PBH mass above 10^{15} g, or peak radiation energy below muon mass
- Hawking evaporation after recombination yields (mostly) e+e- and gamma rays
- For M>>10^{15}g, mass loss negligible during the age of the Universe
- A steady and long-lasting injection of radiation that scales as $(1+z)^3$

Extra-galactic source

Of Hawking evaporation rate

$$\dot{M}_{10} = -5.34 \times 10^{-5} \left(\sum_i f_i \right) M_{10}^{-2} \ \mathrm{s}^{-1}$$

into light (massless) species,

$$f_0 = 0.267, \quad f_1^{\gamma} = 0.06, \quad f_{3/2} = 0.02,$$
$$f_2^{g} = 0.007, \quad f_{1/2}^{\nu} = 0.147, \quad f_{1/2}^{e^{\pm}} = 0.142$$

J.MacGibbon, PRD, 1991

and photons & electrons affects the environment with unit volume injection rate, +redshift

$$\frac{dE}{dV\,dt} = \dot{M}_{PBH}\eta(E_i, z) n_{PBH}$$
$$= \frac{\dot{M}_{10}}{M_{10}} \rho_{cr}(z) \Omega_{PBH}(z) \eta_i(E, z)$$

$\Gamma^{-1} \sim \dot{M}/M$

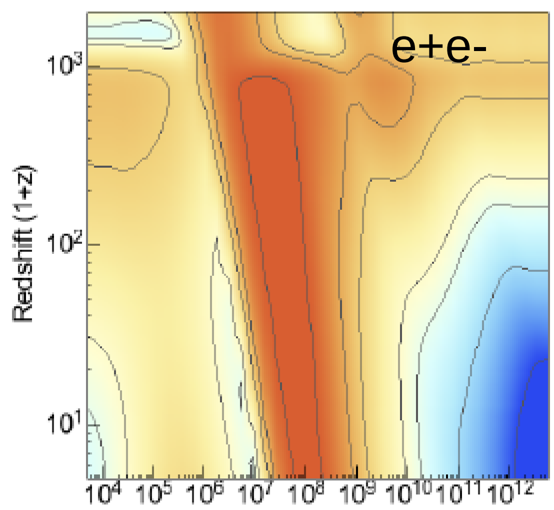
$$\left.\frac{dE}{dV\,dt}\right|_{\mathrm{BH}} = \frac{\dot{M}_{10}}{M_{10}} \rho_{cr}(z) \Omega_{PBH}(z) \eta(E_{PBH}, z)$$

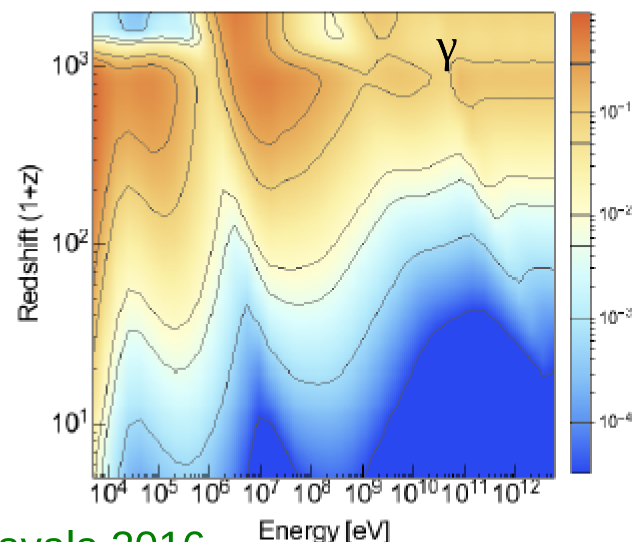
Injection vs absorption

- photons interact via Compton scattering & absorptions
- electrons lose energy by inverse C. scattering & ionization

Not all energy is efficiently absorbed by the environment (gas) esp. if particles are too energetic

Energy “fraction” into ionization (of H)



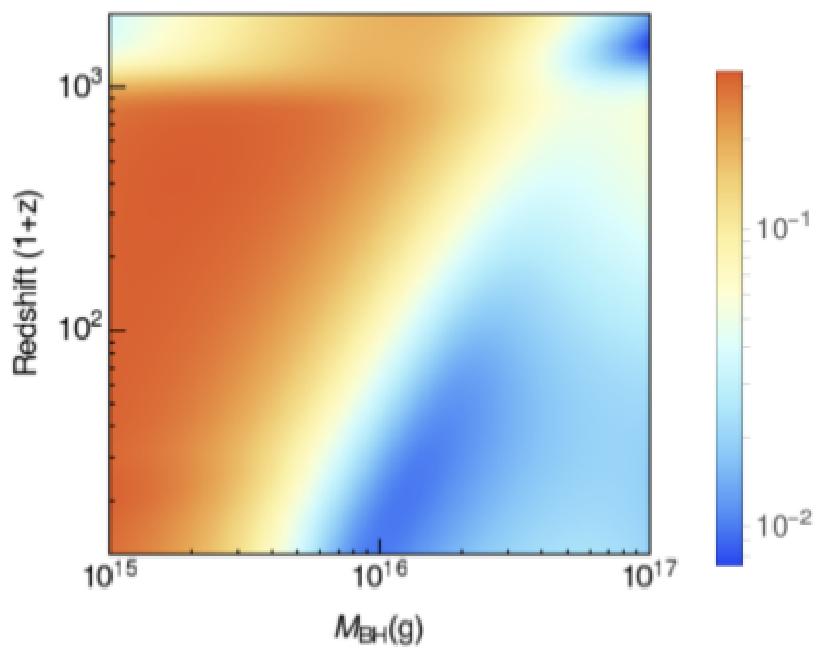


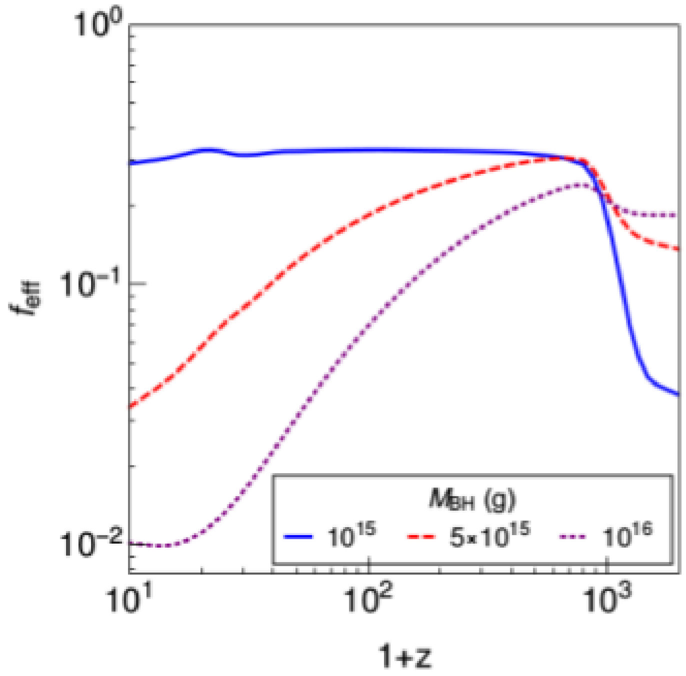
Liu, Slatyer, Zavala 2016

Also see:
Belotsky, Kirillov 2015

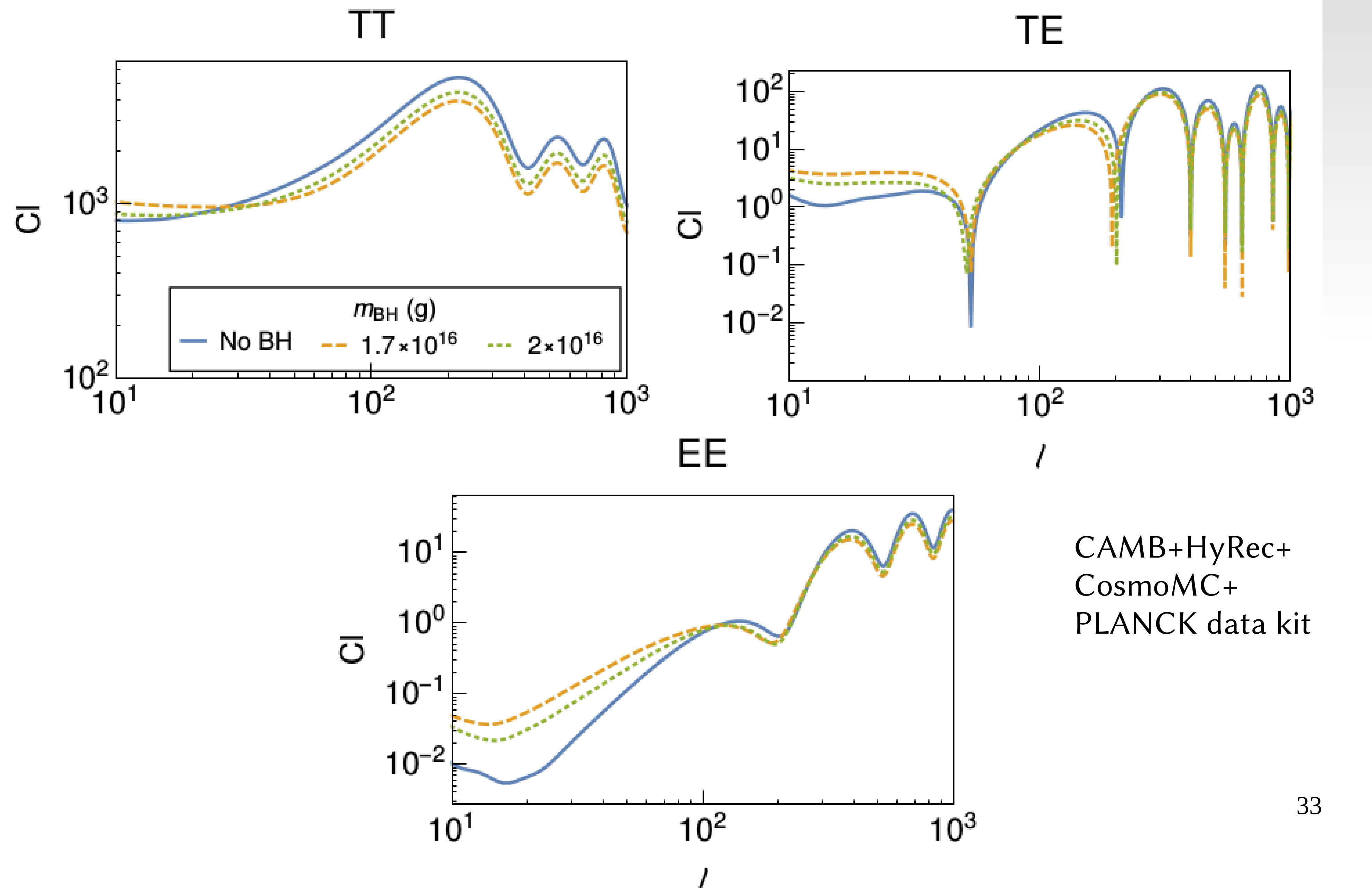
BH Effective absorption efficiencies

- Average over injection (BH radiation) energy spectrum and particle species
- Delayed injection: integrated over earlier z (up to CMB)



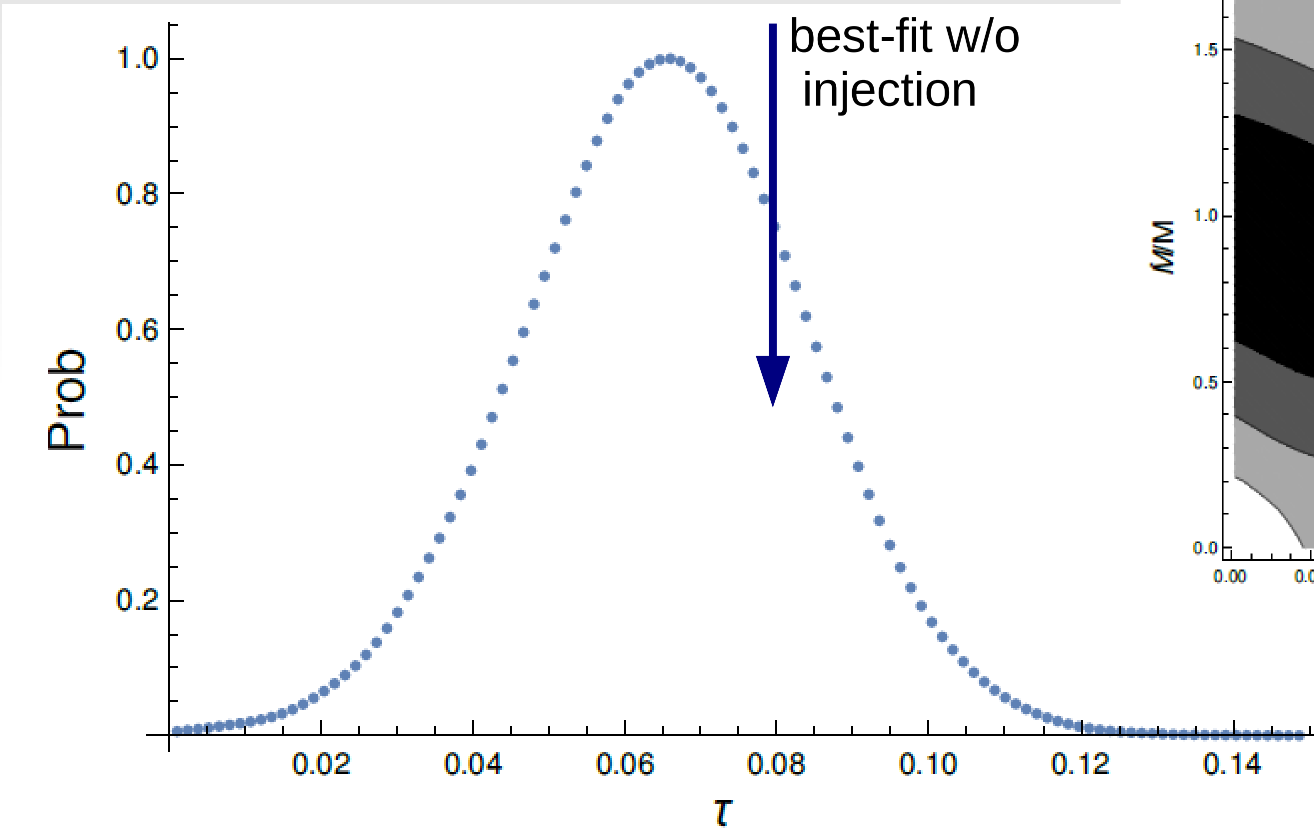


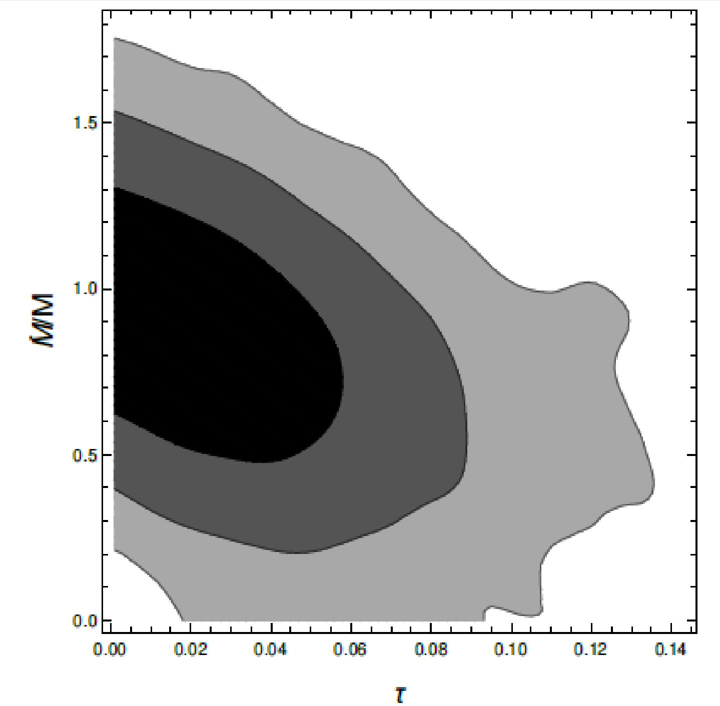
Impact on the CMB Cls



CAMB+HyRec+
CosmoMC+
PLANCK data kit

Reduced optical depth





Consistent with PLANCK no injection fit within 2σ

Consistent with reionization $z \leq 6$

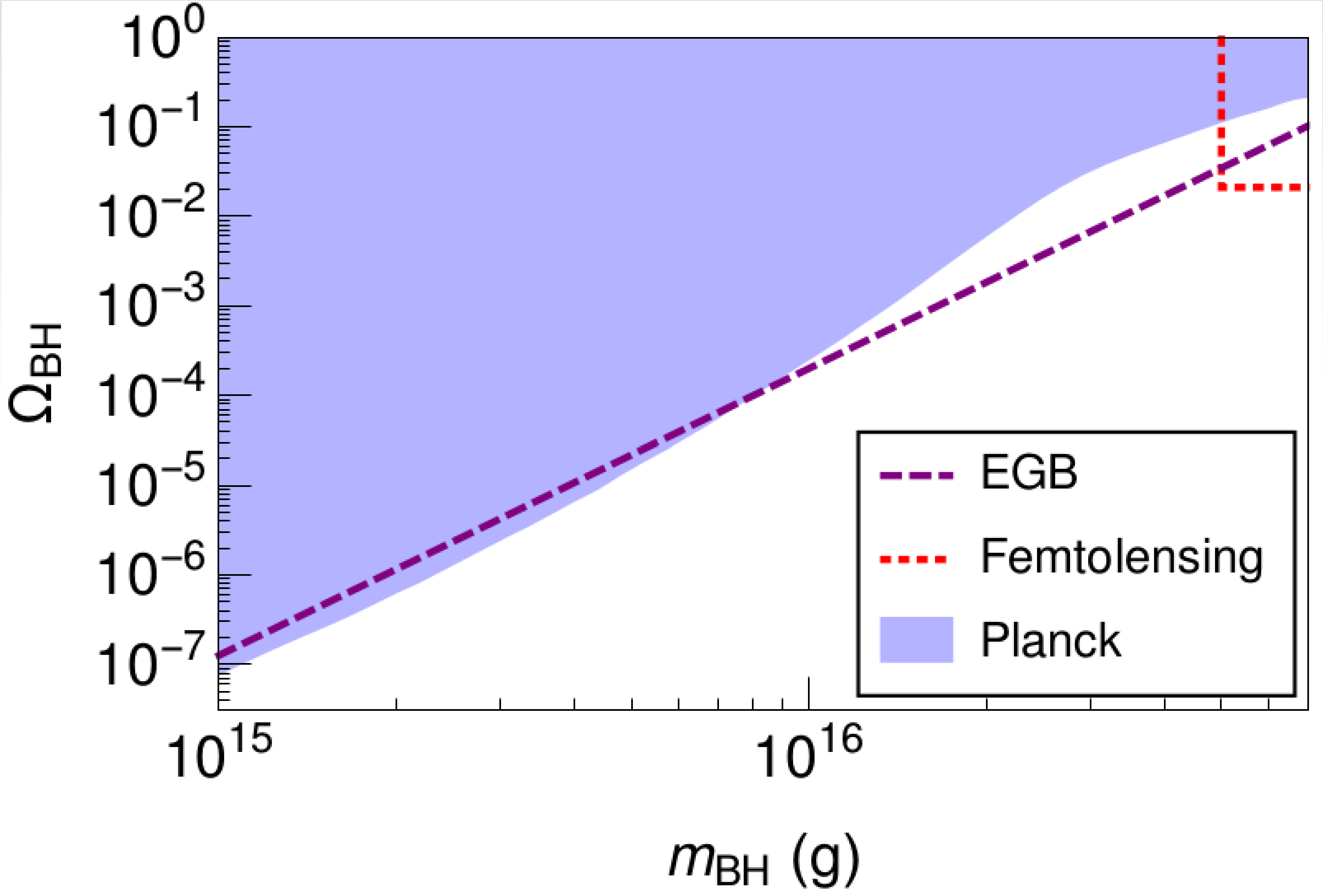
DM Decay: $\tau > 10^{24}$ s

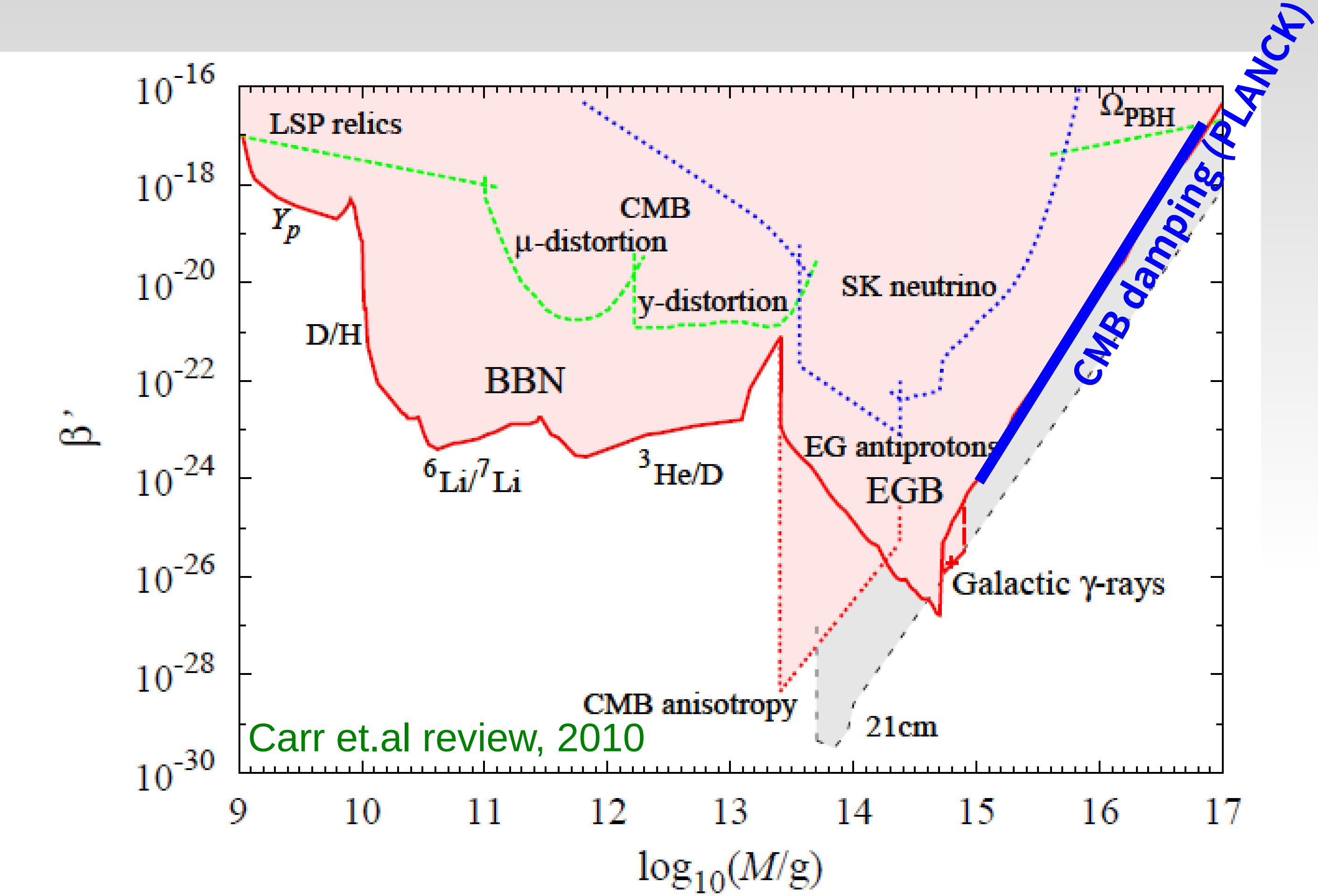
Liu, Slatyer, Zavala 2016

PLANCK: temperature + polarization

Need polarization data to break degeneracy

PLANCK constraint on PBH as `relic' abundance



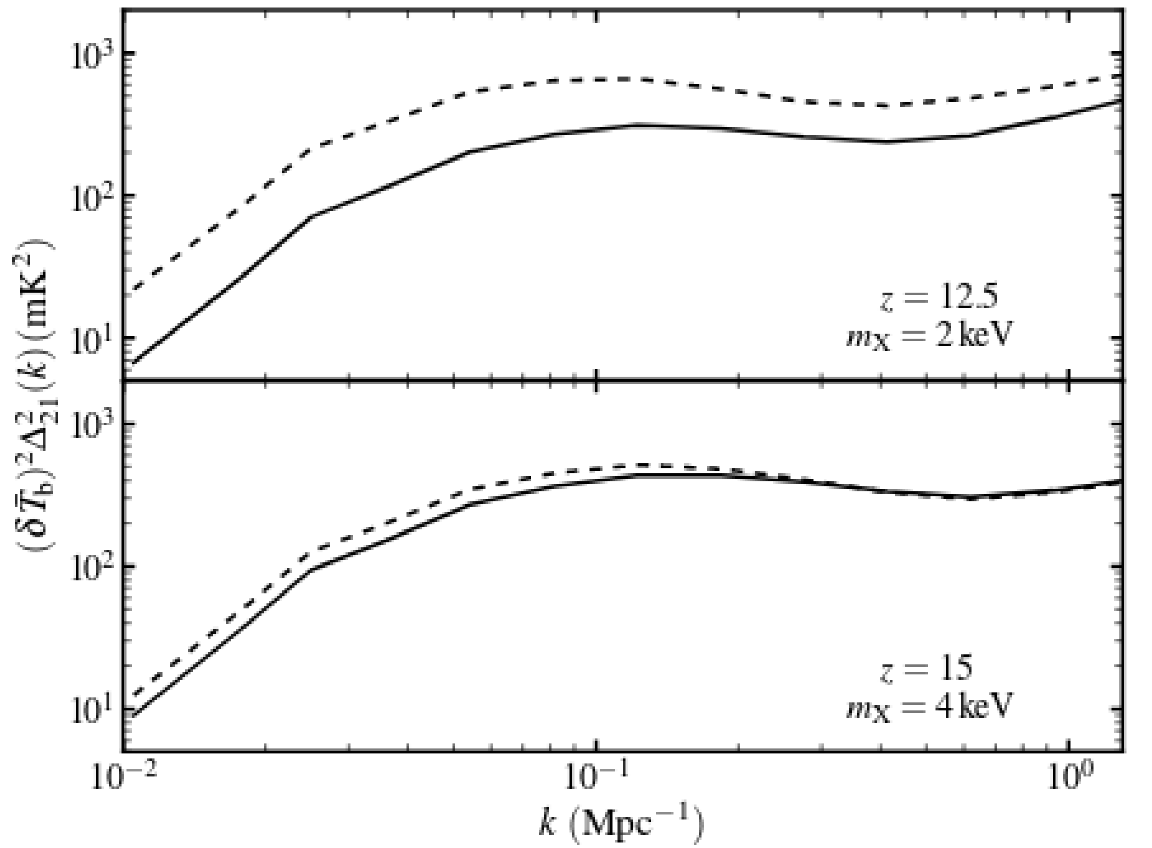


Simple scaling, assume no entropy production after PBHs and $\Omega_r \sim 10^{-4}$ at CMB time, also β' ~ β:

$$\Omega_{\rm PBH} \simeq \beta\, \Omega_{\rm r}\, (1+z) \sim 10^6\, \beta \left(\frac{t}{1\,{\rm s}}\right)^{-1/2} \sim 10^{18}\, \beta \left(\frac{M}{10^{15}\,{\rm g}}\right)^{-1/2} \quad (M > 10^{15}\,{\rm g})\,.$$

Backup: 21cm prospects

SKA: Reionization epoch power spectrum



SKA white paper

Warm versus Cold DM at 21cm

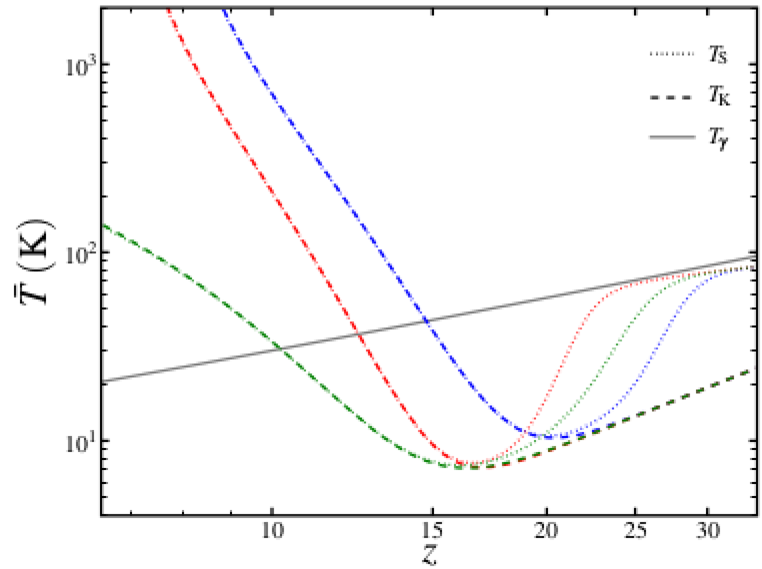


Figure 3. Mean spin temperatures $\bar{T}_S$ for CDM and WDM models. The dotted curves show $\bar{T}_S$ for our fiducial CDM model (blue), WDM with $m_X = 3\,\text{keV}$ (red), and CDM with $f_*/f_{*\text{fid}} = 0.1$ (green). In addition, the mean kinetic temperature $\bar{T}_K$ of each model is plotted with a dashed curve in the same colour used for $\bar{T}_S$. The grey solid line is the CMB temperature. Figure taken from (Sitwell et al., 2014).

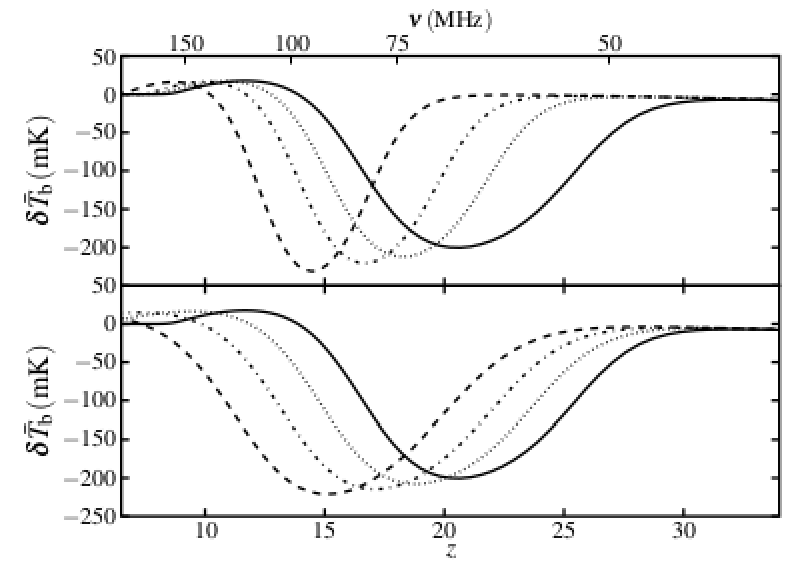


Figure 4. Mean 21 cm brightness temperature $\delta\bar{T}_b$. The solid curve is the fiducial CDM model. The upper plot shows the results of WDM runs where the dashed, dotted-dashed, and dotted curves are for $m_X = 2, 3, 4$ keV, respectively. The lower plot shows CDM runs where the dashed, dotted-dashed, and dotted curves are for CDM models with $f_*/f_{*\text{fid}} = 0.03, 0.1, 0.5$, respectively. Figure taken from Sitwell et al. (2014).

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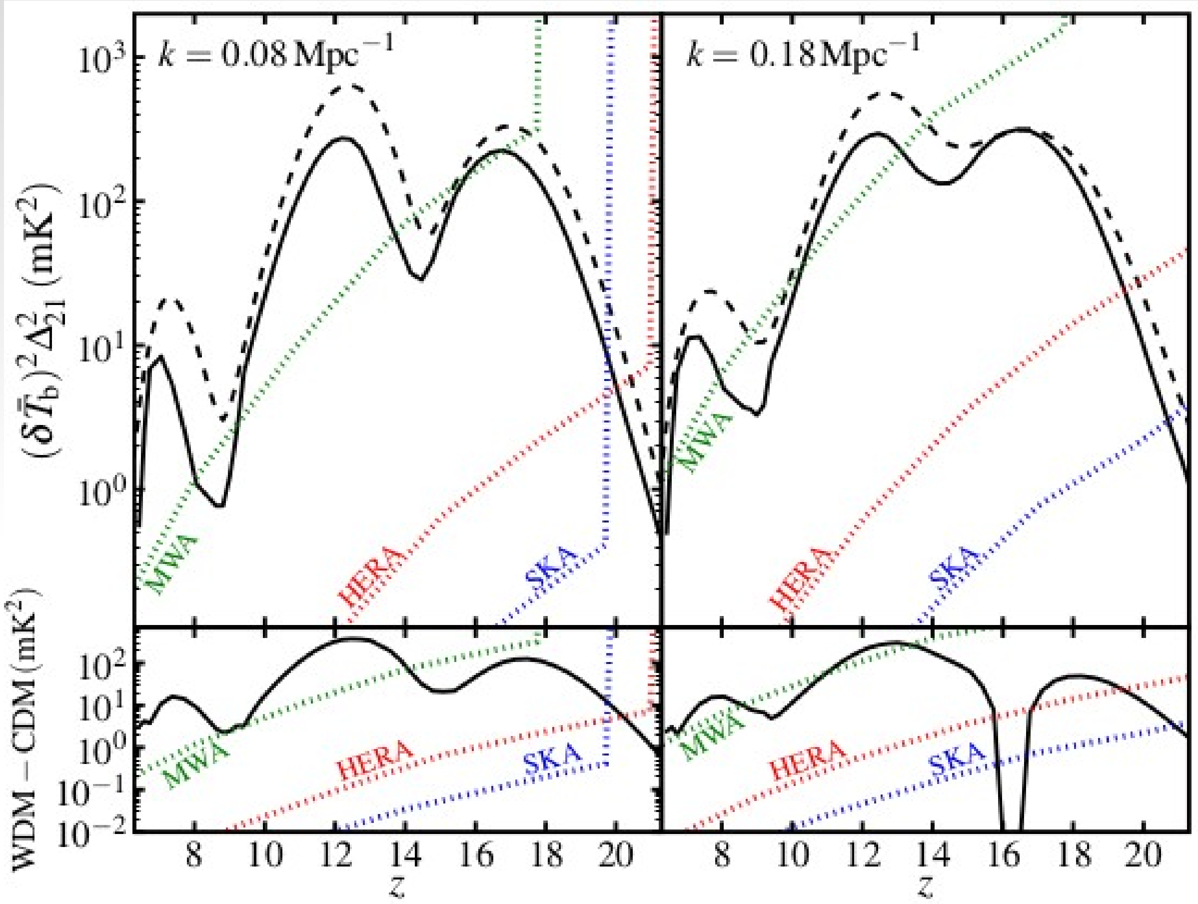

M. Sitwell, A. Mesinger, Y.-Z Ma,
K. Sigurdson, 2014 MNRAS, 438, 2664